

# A kpc-scale resolved study of unobscured and obscured star-formation activity in normal galaxies at $z = 1.5$ and $2.2$ from ALMA and HiZELS

Cheng Cheng,<sup>1,2,3</sup> Edo Ibar,<sup>3</sup> Ian Smail,<sup>4</sup> Juan Molina,<sup>5</sup> David Sobral,<sup>6</sup> Andrés Escala,<sup>7</sup> Philip Best,<sup>8</sup> Rachel Cochrane,<sup>9,8</sup> Steven Gillman,<sup>4</sup> Mark Swinbank,<sup>4,10</sup> R. J. Ivison,<sup>11</sup> Jia-Sheng Huang,<sup>1,2</sup> Thomas M. Hughes,<sup>1,3,12,13</sup> Eric Villard<sup>14</sup> and Michele Cirasuolo<sup>11</sup>

<sup>1</sup>Chinese Academy of Sciences South America Center for Astronomy, National Astronomical Observatories, CAS, Beijing 100101, China, Email: chengcheng@nao.cas.cn

<sup>2</sup>CAS Key Laboratory of Optical Astronomy, National Astronomical Observatories, Chinese Academy of Sciences, Beijing 100101, China,

<sup>3</sup>Instituto de Física y Astronomía, Universidad de Valparaíso, Avda. Gran Bretaña 1111, Valparaíso, Chile

<sup>4</sup>Centre for Extragalactic Astronomy, Durham University, South Road, Durham, DH1 3LE

<sup>5</sup>Kavli Institute for Astronomy and Astrophysics, Peking University, 5 Yiheyuan Road, Haidian District, Beijing 100871, P.R. China

<sup>6</sup>Department of Physics, Lancaster University, Lancaster, LA1 4BY

<sup>7</sup>Departamento de Astronomía (DAS), Universidad de Chile, Casilla 36-D, Santiago, Chile

<sup>8</sup>SUPA, Institute for Astronomy, Royal Observatory Edinburgh, EH9 3HJ

<sup>9</sup>Harvard-Smithsonian Center for Astrophysics, 60 Garden St. Cambridge, MA 02138, USA

<sup>10</sup>Institute for Computational Cosmology, Durham University, South Road, Durham DH1 3LE

<sup>11</sup>European Southern Observatory, Karl-Schwarzschild-Strasse 2, D-85748 Garching, Germany

<sup>12</sup>CAS Key Laboratory for Research in Galaxies and Cosmology, Department of Astronomy, University of Science and Technology of China, Hefei 230026, China

<sup>13</sup>School of Astronomy and Space Science, University of Science and Technology of China, Hefei 230026, China

<sup>14</sup>Joint ALMA Observatory/ESO Avenida Alonso de Cordova 3107, Vitacura, Santiago, Chile

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## ABSTRACT

We present Atacama Large Millimeter/Submillimeter Array (ALMA) continuum observations of a sample of nine star-forming galaxies at redshifts 1.47 and 2.23 selected from the High- $z$  Emission Line Survey (HiZELS). Four galaxies in our sample are detected at high significance by ALMA at a resolution of  $0''.25$  at rest-frame  $355\ \mu\text{m}$ . Together with the previously observed  $\text{H}\alpha$  emission, from adaptive optics-assisted integral-field-unit spectroscopy ( $\sim 0''.15$  resolution), and F606W and F140W imaging from the *Hubble Space Telescope* ( $\sim 0''.2$  resolution), we study the star-formation activity, stellar and dust mass in these high-redshift galaxies at  $\sim$ kpc-scale resolution. We find that ALMA detection rates are higher for more massive galaxies ( $M_* > 10^{10.5} M_\odot$ ) and higher  $[\text{N II}]/\text{H}\alpha$  ratios ( $> 0.25$ , a proxy for gas-phase metallicity). The dust extends out to a radius of 8 kpc, with a smooth structure, even for those galaxies presenting clumpy  $\text{H}\alpha$  morphologies. The half-light radii ( $R_{\text{dust}}$ ) derived for the detected galaxies are of the order  $\sim 4.5$  kpc, more than twice the size of submillimetre-selected galaxies at a similar redshift. Our global star-formation rate estimates — from far-IR and extinction-corrected  $\text{H}\alpha$  luminosities — are in good agreement. However, the different morphologies of the different phases of the interstellar medium suggest **complex extinction properties of the high-redshift normal galaxies**.

**Key words:** galaxies: ISM – submillimetre: galaxies – galaxies: starburst – galaxies: star formation

## 1 INTRODUCTION

Observations of star-formation activity are critically important to tackle open questions relating to galaxy formation and evolution. The most direct view of cosmic star-formation history comes from observing ultraviolet (UV) photons from the young massive stars. However, part of this radiation is absorbed by dust, and the higher

UV-energy photons are absorbed by neutral hydrogen. Gas ionised by this radiation eventually recombines, producing emission lines such as  $\text{Ly}\alpha$  and  $\text{H}\alpha$ , which have been used historically to estimate the star-formation rate (SFR, e.g., Sobral & Matthee 2019). Where the star-formation activity is shielded by gas and dust, the ionising photons may get absorbed and re-radiated at far-infrared (FIR)

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wavelengths by dust. Therefore, the rest-frame FIR emission is used to trace dust-obscured star formation in galaxies, and to derive total SFRs in combination with tracers of unobscured emission (e.g. Kennicutt 1998; Kennicutt et al. 2012).

Different ways to estimate the SFR in galaxies have their own limitations and biases. For example, the  $H\alpha$  recombination line is generated originally by the photo-ionising radiation from massive stars ( $\gtrsim 10 M_{\odot}$ ), or AGN, and is therefore sensitive to recent star formation, within  $\sim 10$  Myr, and has modest sensitivity to dust obscuration. The UV flux ( $\sim 1600 \text{ \AA}$ ), on the other hand, comes from young, massive stars but is also emitted by older 10–100 Myr stars (Kennicutt et al. 2012), and is very sensitive to dust obscuration. The FIR emission produced by heated dust grains comes from young star-forming regions but in lower luminosity sources, it can also arise from dust heated by older stellar populations. Spatially-resolved observations of local galaxies show that all of these SFR tracers are strongly correlated on  $\sim$  kpc scales (e.g., Boquien et al. 2016).

Previous observations of high-redshift galaxies have shown that their star-formation activity — as traced by UV or  $H\alpha$  emission — presents bright, clumpy star-forming complexes that could be up to  $\sim 1000\times$  more massive than those seen in local galaxies (e.g., Genzel et al. 2006, 2008; Förster Schreiber et al. 2006, 2009, 2018; Shapiro et al. 2008; Swinbank et al. 2012a,b). However, it may be that regions with fainter  $H\alpha$  and UV emissions are affected by significant dust obscuration. Previous work has shown that dust attenuation causes the SFR derived from different indicators to be inconsistent (Swinbank et al. 2004; Katsianis et al. 2017), thus it is essential to have a spatially resolved view of the ionised gas and the dust content in order to characterise the total SFR (unobscured and obscured), especially for high-redshift galaxies.

The High- $z$  Emission Line Survey (HiZELS, Sobral et al. 2012, 2013, 2015; Geach et al. 2008) was designed to study ‘normal’ star-forming galaxies selected in narrow redshift slices at 0.4, 0.84, 1.47 and 2.23 via the identification of  $H\alpha$  emission using near-infrared (near-IR) narrow-band filter imaging in extragalactic survey fields including the Subaru-*XMM* Deep Field / UKIDSS Ultra Deep Survey (UDS, Lawrence et al. 2007) and the Cosmological Evolution Survey (COSMOS, Scoville et al. 2007) fields. The HiZELS survey detects thousands of emission-line objects, samples the ‘typical’ galaxy population (Oteo et al. 2015; Cochrane et al. 2018), following the so-called ‘main sequence’ for star-forming galaxies at  $z \approx 1.47$  and 2.23 (Gillman et al. 2019), and probing below the knee of the  $H\alpha$  luminosity function ( $< L_{H\alpha}^*$ ) at these redshifts ( $L_{H\alpha}^* = 10^{42.6} \text{ erg s}^{-1}$  at  $z \approx 1.47$  and  $L_{H\alpha}^* = 10^{42.9} \text{ erg s}^{-1}$  at  $z \approx 2.23$ , Swinbank et al. 2012a; Cochrane et al. 2017). Over thirty galaxies from HiZELS have also been mapped in the follow-up near-IR using integral field unit (IFU) spectroscopy (the ‘SHIZELS’ sample) aided by adaptive optics (AO), with the Spectrograph for INtegral Field Observations in the Near Infrared (SINFONI) at the Very Large Telescope (VLT) or the Near-Infrared Integral Field Spectrometer (NIFS) at Gemini-North (Swinbank et al. 2012a,b; Molina et al. 2017; Gillman et al. 2019). These observations provide  $H\alpha$  IFU imaging at  $\sim 1$  kpc scales for galaxies at  $z \approx 1.47$  or 2.23.

In this work, we make use of the available AO-aided IFU data and the *Hubble Space Telescope* (HST) F606W, F140W data, together with new Atacama Large Millimeter/submillimeter Array (ALMA) observations targeting the continuum emission at submillimetre wavelengths at similar spatial resolution (synthesised beam of  $\sim 0''.2 - 0''.3$  full width half maximum), to characterise the spatial correlation between the the  $H\alpha$ /UV and dust emission of

high-redshift  $z \approx 1.47$  and 2.23 galaxies (the ‘ALMA-SHiZELS’ sample). The combination of the AO-aided IFU  $H\alpha$  data, HST data and the ALMA observations provides a unique opportunity to characterise the spatial correlation between the SFR tracers around the peak of cosmic star formation. We describe our observations, data reduction and analysis in Section 2. In Section 3, we present the results and discussion, and summarise in Section 4. Throughout this work, we assume a  $\Lambda$ CDM cosmology with  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_M = 0.3$ ,  $\Omega_{\Lambda} = 0.7$ , where  $0''.15$  corresponds to a physical scale of  $\sim 1.3$  kpc at both  $z \approx 1.47$  and 2.23.

## 2 OBSERVATIONS AND DATA REDUCTION, AND ANALYSIS

### 2.1 ALMA observations

Nine SHiZELS galaxies were observed with ALMA during Cycle 2, 3, 4 (projects 2012.1.00402.S, 2013.1.01188.S and 2015.1.00026.S; P.I.: E. Ibar) in Bands 6 or 7, depending on whether the galaxies were at  $z = 2.23$  or 1.47, respectively, corresponding to rest-frame  $\sim 355 \mu\text{m}$ . The observations were designed to detect continuum emission using four spectral windows (SPWs), each covering an effective bandwidth of 1.875 GHz at a spectral resolution of 15.6 MHz. Observations were taken under relatively good weather conditions with precipitable water vapour (P.W.V.) ranging from 0.2 mm to 3.9 mm, and using 29 to 45 antennas (the earlier the cycle, the smaller the number) with the longest baselines spanning 1.0–1.5 km. The phase, bandpass and flux calibrators for all observations used in this work are listed in Table 1.

Data reduction was carried out using the Common Astronomy Software Applications (CASA) and using the ALMA pipeline up to the production of calibrated  $uv$  data products. Three imaging approaches were performed using the task `TCLEAN`, exploring a Briggs weighting (`ROBUST=0.5`), natural weighting, and  $uv$  tapering such that we created a synthesized beam of  $\sim 1''$  (see Figs A1, A2). In all cases, the signal is interactively cleaned down to  $2-3 \sigma$  (r.m.s.  $\sim 25 \mu\text{Jy beam}^{-1}$ ) in regions with significant emission at the source position. The astrometric accuracy of the ALMA Band-7 image is around 1.5 mas. Information about each target is presented in Table 2.

Sources SHiZELS-7, SHiZELS-9, SHiZELS-11 and SHiZELS-14 were detected by ALMA at a significance higher than  $5\sigma$  in at least one of the three imaging approaches, while SHiZELS-21, SHiZELS-8, SHiZELS-10, SHiZELS-2 and SHiZELS-3 remained undetected, regardless of the imaging approach.

Within the field of view (FoV) of SHiZELS-7 and SHiZELS-10, we identify two serendipitous detections. We denote them as SHiZELS7-ID2 and SHiZELS10-ID2 and their properties are listed in Appendix D.

### 2.2 Global astrometry

One of the main challenges when analysing multi-wavelength high-resolution spatially-resolved observations is the global astrometric accuracy. In particular, IFU observations with a small FoV (as obtained with SINFONI, which has a FoV of about  $3'' \times 3''$ ) present a global astrometric uncertainty which is larger than the resolution elements of the observations, therefore they are difficult to anchor to other observations at high-resolution. For this reason, we perform the following astrometric corrections to analyse our data.

**Table 1.** ALMA observational set-up, ordered by the observation date.

Source List	Project ID	Observation Date	Flux Calibrator	Bandpass Calibrator	Phase Calibrator	P.W.V. (mm)	Number of antennas	Band ( $\nu_{\text{obs}}$ /GHz)	Time on Target (min)
SHIZELS-8	2012.1.00402.S	6 Nov. 2013	Uranus	J2148+0657	J0215-0222	3.9	29	7 (344)	52
		29 Nov. 2013	Uranus	J0334-4008	J0215-0222	1.3	26		43
		29 Nov. 2013	J0423-013	J0334-4008	J0215-0222	1.0	26		43
SHIZELS-7	2013.1.01188.S	11 Aug. 2015	Ceres	J0006-0623	J0208-0047	0.9	43	7 (344)	28
		29 Aug. 2015	J0334-401	J0006-0623	J0219+0120	1.5	37	6 (261)	30
		29 Aug. 2015	J0238+166	J0224+0659	J0219+0120	1.4	37	6 (261)	30
SHIZELS-3	2015.1.00026.S	28 Jul. 2016	J1058+0133	J1058+0133	J0948+0022	0.6	37	6 (261)	26
		16 Jul. 2016	J0238+1636	J0238+1636	J0217+0144	0.4	38	7 (344)	43
		26 Jul. 2016	J0238+1636	J0238+1636	J0217+0144	0.4	38	7 (344)	43
		26 Jul. 2016	J0238+1636	J0238+1636	J0217+0144	0.6	45	7 (344)	43
		10 Aug. 2016	J0006-0623	J0006-0623	J0217+0144	0.2	42		43
		10 Aug. 2016	J0238+1636	J0238+1636	J0217+0144	0.2	39		43
SHIZELS-14		2 Aug. 2016	J1058+0133	J1058+0133	J0948+0022	0.7	39	6 (261)	26

**Table 2.** Properties of the SHIZELS galaxies presented in this work. ID and  $z_{\text{H}\alpha}$  and  $M_{\star}$  are extracted from previous AO-aided IFU observations presented in Gillman et al. (2019).  $S_{355\mu\text{m}}$  and  $R_{\text{dust}}$  are the observed continuum flux density and the deconvolved half-light radius at  $355\mu\text{m}$  (rest-frame). ‡ Possible AGN.

ID	R.A. (J2000)	Dec. (J2000)	$z_{\text{H}\alpha}$	$\log(L_{\text{H}\alpha}/\text{erg s}^{-1})$	$S_{355\mu\text{m}}$ (mJy)	$\log(L_{\text{FIR}}/L_{\odot})$	$\log(M_{\star}^{\text{MAGPHYS}}/M_{\odot})$	$\log(\text{SFR}^{\text{MAGPHYS}}/M_{\odot}\text{yr}^{-1})$	$A_{\nu}$	$R_{\text{dust}}$ (kpc)	$R_{\text{dust}}^{\text{fit}}$ (kpc)
SHIZELS-7	02:17:00.34	-05:01:50.6	1.455	42.1	$0.15 \pm 0.03$	$10.5 \pm 0.4$	$10.4 \pm 0.2$	$1.6 \pm 0.3$	$0.8 \pm 0.5$	$3.7 \pm 0.3$	$5.1 \pm 0.4$
SHIZELS-9	02:17:12.99	-04:54:40.7	1.462	42.4	$0.51 \pm 0.06$	$11.2 \pm 0.4$	$10.6 \pm 0.1$	$1.9 \pm 0.2$	$1.0 \pm 0.3$	$4.9 \pm 0.3$	$4.8 \pm 0.4$
SHIZELS-11‡	02:18:21.23	-05:02:48.9	1.492	42.3	$1.34 \pm 0.14$	$11.6 \pm 0.4$	$11.5 \pm 0.1$	$1.7 \pm 0.1$	$0.8 \pm 0.1$	$5.2 \pm 0.4$	$4.2 \pm 0.5$
SHIZELS-14	10:00:51.58	+02:33:34.1	2.242	42.9	$2.31 \pm 0.23$	$12.6 \pm 0.3$	$11.1 \pm 0.1$	$2.6 \pm 0.2$	$1.5 \pm 0.2$	$4.7 \pm 0.3$	$5.5 \pm 0.4$
SHIZELS-10	02:17:39.02	-04:44:41.4	1.447	41.9	$< 0.11 (5\sigma)$	$< 11.3 (5\sigma)$	$10.1 \pm 0.1$	$2.0 \pm 0.2$	$1.4 \pm 0.2$	–	–
SHIZELS-8	02:18:20.96	-05:19:07.5	1.460	42.2	$< 0.12 (5\sigma)$	$< 11.2 (5\sigma)$	$10.3 \pm 0.1$	$1.9 \pm 0.2$	$1.4 \pm 0.4$	–	–
SHIZELS-2	02:19:25.50	-04:54:39.6	2.223	42.3	$< 0.12 (5\sigma)$	$< 11.9 (5\sigma)$	$9.9 \pm 0.2$	$1.4 \pm 0.3$	$0.3 \pm 0.3$	–	–
SHIZELS-3	10:00:27.69	+02:14:30.6	2.225	42.4	$< 0.10 (5\sigma)$	$< 11.8 (5\sigma)$	$9.0 \pm 0.1$	$0.9 \pm 0.2$	$0.2 \pm 0.1$	–	–
SHIZELS-21	02:16:45.82	-05:02:45.0	2.237	42.6	$< 0.12 (5\sigma)$	$< 11.8 (5\sigma)$	$9.8 \pm 0.1$	$1.6 \pm 0.2$	$1.2 \pm 0.3$	–	–

Firstly, we align the astrometry of the HiZELS narrow-band images by the public *Gaia* DR2 catalogue. The offset between the HiZELS catalogue generated by SExtractor and the *Gaia* catalogue is corrected to  $\Delta\text{R.A.} = 0''.004 \pm 0''.13''$  and  $\Delta\text{Dec.} = -0''.010 \pm 0''.10$ . We degrade the resolution of the H $\alpha$  moment-0 SINFONI image down to  $0''.7$  using a Gaussian kernel to get a similar resolution to the HiZELS narrow-band images. Then we fit a Gaussian profile to this low-resolution image in order to obtain the position of the peak, identify the offset with respect to the narrow-band image, and use this information to correct the astrometry of the IFU image.

As mentioned before, we have also used the available *HST* F606W and F140W observations<sup>1</sup>. The *HST* FoV is about  $2' \times 2'$ , which is too small to apply a reliable astrometric correction using *Gaia* DR2. Therefore, we correct the global astrometry of the *HST* images using the HSC DR2 catalogue available in our fields (Aihara et al. 2019). The  $5\sigma$  limiting magnitude within  $2''$  diameter apertures of the HSC *i*-band catalogue is 26.7 AB mag, which is similar to the *HST* images, and deep enough to have about 200 high-S/N sources for the alignment. The astrometric accuracy of HSC DR2 catalogue derives from the *Gaia* DR2 catalogue (Aihara et al. 2019). We match the *HST* and HSC catalogues for the offsets of R.A., Dec., and correct the astrometry of the *HST* image. After our astrometric correction, the catalogue match between *HST* and HSC catalogues is consistent with an offset of zero, with systematic errors of  $\sim 0''.04$ .

In Fig. 1, we show the true-colour image composite from *HST* F606W (in blue) and F140W (in red) images, including ALMA

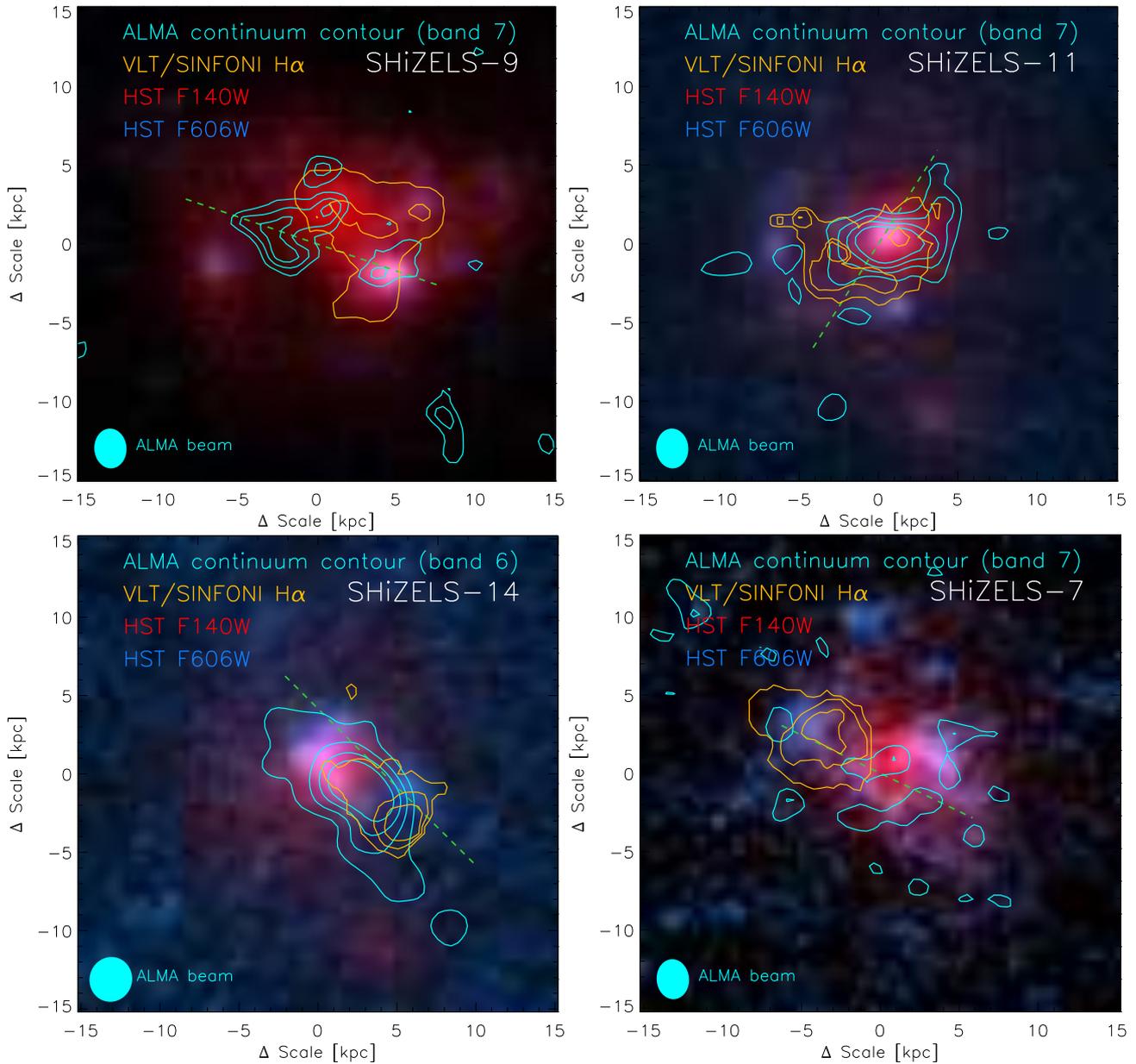
continuum contours in levels of  $2\sigma$ ,  $3\sigma$ ,  $4\sigma$ ,  $5\sigma$  for the four ALMA-detected galaxies. The galaxy centres revealed by *HST* more or less coincide with the ALMA imaging. The observations reveal a complex state for the interstellar medium in SHIZELS galaxies at  $z \approx 1.47$  and  $2.23$ . We show multi-wavelength postage stamps of the *HST*, H $\alpha$  (Gillman et al. 2019) and our ALMA images for each of the galaxies presented in this work in Figs A1 and A2.

### 2.3 MAGPHYS fitting with ALMA flux

SEDs of the SHIZELS sample were previously fitted using MAGPHYS (da Cunha et al. 2008) in Gillman et al. (2019). The ALMA detected fluxes of the ALMA-SHIZELS galaxies can help to constrain the FIR properties. Therefore, we re-fit the SEDs of the four ALMA detected galaxies including the FIR flux we obtained in this work, and list the main results in Table 2 and 3. For the ALMA non-detected galaxies, we list the dust mass given by MAGPHYS from fitting the optical-to-NIR SEDs. The stellar mass and dust mass show reasonable consistencies from the different approach.

We show the dust-to-stellar mass ratio ( $\log_{10}(M_{\text{dust}}/M_{\star})$ ) in Fig. 2 and compare our results with the local galaxies from Dustpedia (Clark et al. 2018; Casasola et al. 2020). The only ULIRG in our sample, SHIZELS-14, has a larger  $\log_{10}(M_{\text{dust}}/M_{\star}) = -2.1 \pm 0.2$  than the local galaxies, while this value is consistent with the typical value of  $\log_{10}(M_{\text{dust}}/M_{\star})$  in SMGs at redshift about 2 (Calura et al. 2017). The rest targets of our sample all have consistent or  $1\sigma$  higher  $\log_{10}(M_{\text{dust}}/M_{\star})$  values as the local galaxies, even for the targets with no ALMA detection.

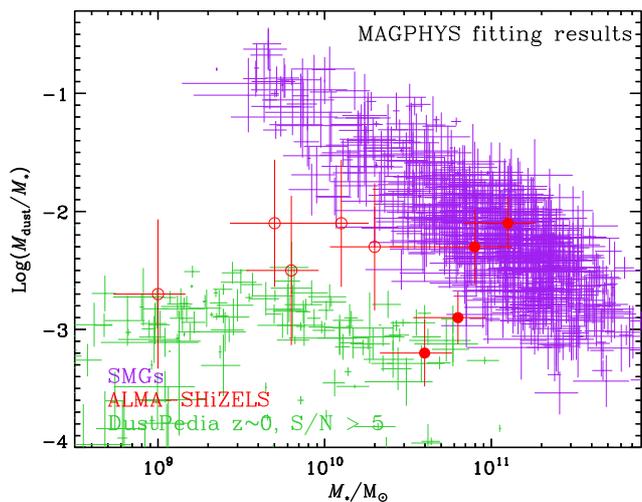
<sup>1</sup> *HST* proposal ID: 14719, P.I.: P. Best

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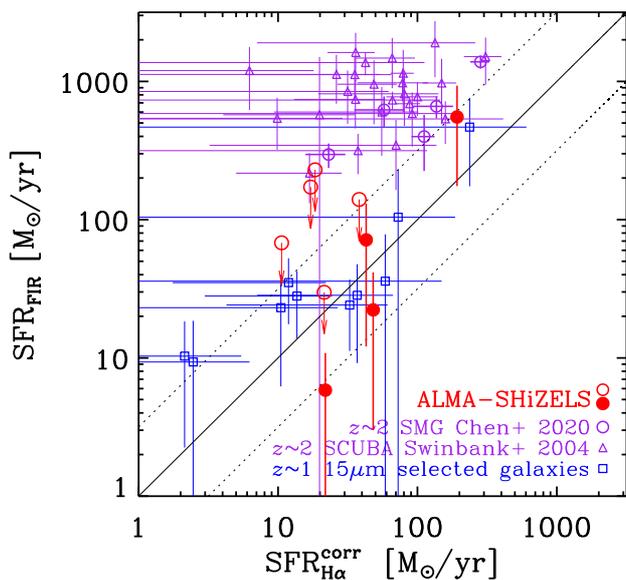
**Figure 1.** The *HST* F606W (blue) and F140W (red) true-colour composite images, including ALMA continuum contours (cyan) and H $\alpha$  contours (orange). Contour levels are shown at 2-, 3-, 4- and 5- $\sigma$ . The ALMA synthesised beam is shown in the lower left corner. The green dash lines show the direction of major kinematic axis ( $PA_{\text{vel}}$ , see the velocity map in Gillman et al. 2019). The three different tracers of star formation are clearly highlighting very different regions.

**Table 3.** Gas phase metallicities derived from previous AO-aided IFU observations (following Curti et al. (2017)), gas to dust ratio (see § 2.5), total ISM mass estimated following Scoville et al. (2016), the dust mass from gas-to-dust ratio ( $\log(M_{\text{dust}}^{\text{GDR}})$ ) and SED fitting ( $\log(M_{\text{dust}}^{\text{SED}})$ ) for the ALMA-SHiZELS galaxies presented in this work. ‡ Possible AGN. For the ALMA detected galaxy, the MAGPHYS fitting includes the new ALMA flux.

ID	[N II]/H $\alpha$	12+log(O/H)	$\delta_{\text{GDR}}$	$\log(M_{\text{ISM}}/M_{\odot})$	$\log(M_{\text{dust}}^{\text{GDR}}/M_{\odot})$	$\log(M_{\text{dust}}^{\text{SED}}/M_{\odot})$	$\log(M_{\text{dust}}^{\text{MAGPHYS}}/M_{\odot})$	$\log(M_{\text{dust}}^{\text{MAGPHYS}}/M_{*})$
SHiZELS-7	$0.43 \pm 0.05$	$8.79 \pm 0.11$	$69 \pm 24$	$9.65 \pm 0.05$	$7.8 \pm 0.2$	$7.5 \pm 0.2$	$7.4 \pm 0.2$	$-3.2 \pm 0.3$
SHiZELS-9	$0.27 \pm 0.03$	$8.64 \pm 0.11$	$97 \pm 30$	$10.17 \pm 0.04$	$8.2 \pm 0.1$	$8.1 \pm 0.2$	$7.9 \pm 0.1$	$-2.9 \pm 0.2$
SHiZELS-11‡	$0.60 \pm 0.10$	$8.87 \pm 0.12$	$57 \pm 21$	$10.59 \pm 0.06$	$8.8 \pm 0.2$	$8.5 \pm 0.2$	$8.6 \pm 0.1$	$-2.3 \pm 0.3$
SHiZELS-14	$0.60 \pm 0.05$	$8.87 \pm 0.11$	$57 \pm 21$	$11.19 \pm 0.06$	$9.4 \pm 0.2$	$8.8 \pm 0.1$	$9.0 \pm 0.1$	$-2.1 \pm 0.2$
SHiZELS-10	$0.13 \pm 0.04$	$8.47 \pm 0.17$	$143 \pm 37$	$< 9.50$	$< 7.3$	$< 7.4$	$8.0 \pm 0.5$	$-2.1 \pm 0.5$
SHiZELS-8	$< 0.1$	$< 8.42$	$> 264$	-	-	$< 7.4$	$8.0 \pm 0.5$	$-2.3 \pm 0.5$
SHiZELS-2	$0.12 \pm 0.01$	$8.45 \pm 0.11$	$150 \pm 38$	$< 9.92$	$< 7.7$	$< 7.6$	$7.3 \pm 0.6$	$-2.5 \pm 0.6$
SHiZELS-3	$0.03 \pm 0.01$	$8.13 \pm 0.29$	$310 \pm 52$	$< 9.84$	$< 7.3$	$< 7.5$	$6.3 \pm 0.6$	$-2.7 \pm 0.6$
SHiZELS-21	$0.23 \pm 0.04$	$8.60 \pm 0.16$	$106 \pm 32$	$< 9.92$	$< 7.9$	$< 7.6$	$7.6 \pm 0.5$	$-2.1 \pm 0.5$



**Figure 2.** The dust mass to stellar mass of our sample (red), the local galaxies from the Dustpedia project (green, Clark et al. 2018), and the SMGs from the ALMA SCUBA-2 UDS survey (purple, Dudzevičiūtė et al. 2020). The dust mass and the stellar mass of our sample and the comparison samples are derived from the SED fitting results by MAGPHYS. The red filled circles are the ALMA detected targets while the red open circles are the ALMA non-detected targets. The target with the stellar mass above  $10^{11} M_{\odot}$  is SHiZELS-14, which is a ULIRG.



**Figure 3.** The total SFRs derived from the extinction-corrected  $H\alpha$  (Eq. 3), and FIR luminosities for  $z \sim 2.23$  and  $z \sim 1.47$  samples, respectively. We show the ALMA detected targets as filled red circles and the ALMA non-detected targets as open red circles. We compare our findings with previously observed SMGs at  $z \sim 2$  from Swinbank et al. (2004) and Chen et al. (2020). A sample of  $15\text{-}\mu\text{m}$ -selected galaxies (at  $z \sim 1$ ) presented by Franceschini et al. (2003) are also shown. The dotted lines show 0.5 dex above and below the solid line.

#### 2.4 Global SFRs from $H\alpha$ and FIR emission

To estimate the FIR luminosity (and  $\text{SFR}_{\text{FIR}}$ ) using the ALMA observations (rest-frame continuum at  $\sim 355 \mu\text{m}$ ), we assume a FIR SED template based on previous stacking analyses for HiZELS galaxies at  $z \sim 1.47$  (Ibar et al. 2013) and  $2.23$  (Thomson et al. 2017). Thanks to the rich multi-wavelength coverage, the stellar

masses of the HiZELS galaxies can be estimated reasonably well (see Sobral et al. 2014; Gillman et al. 2019). For the purpose of this work, for the  $z \sim 1.47$  targets, we adopt the stacked SEDs derived by Ibar et al. (2013) in the stellar mass bins of  $9.9 < \log(M_*/M_{\odot}) < 10.3$  and  $10.3 < \log(M_*/M_{\odot}) < 11.8$ . For galaxies at  $z \sim 2.23$ , we consider the FIR template presented in Thomson et al. (2017).

We fit each FIR template SED following the modified-blackbody fitting method described by Beelen et al. (2006), assuming a fixed power-law index for the dust emissivity,  $\beta = 1.8$ . The fitting results show a dust temperature of about  $25 \pm 1$  K for  $z \approx 1.47$  targets and  $32 \pm 2$  K for the  $z \approx 2.23$  targets, similar to previous measurements of luminous IR galaxies at high redshift (Hwang et al. 2010; Oteo et al. 2017; Liang et al. 2019). To estimate FIR luminosities, we normalise the assumed SED to the observed ALMA flux densities. We derive dust masses following the method presented in Beelen et al. (2006), finding values in the range of  $10^{7.1} - 10^{8.9} M_{\odot}$ . From Table 3 we can see that the dust masses derived from SED fitting and the MAGPHYS are consistent. For the galaxies with no ALMA detection, we use global upper limits using  $5\times$  the r.m.s. of the tapered image.

Uncertainties for the FIR properties come from the SED fitting, FIR flux measured from ALMA results, and template assumption. We expect dust-temperature uncertainties for galaxy templates for  $M_* > 10^{10} M_{\odot}$  galaxies of about 5 K (see the Fig. 5 in Ibar et al. 2013), which leads to a systematic uncertainty in  $\log(L_{\text{FIR}}/L_{\odot})$  of about 0.3 dex. We run a Monte-Carlo simulation to sample the dust temperature, assuming a Gaussian distribution centred at  $T_{\text{dust}}$  from the FIR SED templates, and a scatter  $\sigma_{T_{\text{dust}}} = 5$  K. We also sample the ALMA flux density (or flux density limit) from a Gaussian distribution centred at the measured flux density with the observed scatter. Then we derive  $L_{\text{FIR}}$  and  $M_{\text{dust}}$ , and their r.m.s. scatter, such that the scatter of  $L_{\text{FIR}}$  includes the uncertainty in  $T_{\text{dust}}$  and ALMA flux density. Since we only have one detected FIR band for most sources, we cannot sensibly adopt a more complex model to understand the FIR properties.

The obscured SFR, as derived from the observed FIR emission, can be estimated by Kennicutt et al. (2012):

$$\log \text{SFR}_{\text{FIR}} (M_{\odot} \text{ yr}^{-1}) = \log L_{\text{FIR}(8-1000\mu\text{m})} (\text{erg s}^{-1}) - 43.47, \quad (1)$$

assuming a Chabrier initial mass function (IMF; Chabrier 2003).

On the other hand, to correct the flux observed from the narrow-band  $H\alpha$  imaging for dust extinction, we assume a parametrisation of the  $H\alpha$  extinction as a function of stellar mass following Garn & Best (2010):

$$A_{H\alpha}(M_*) = -0.09X^3 + 0.11X^2 + 0.77X + 0.91, \quad (2)$$

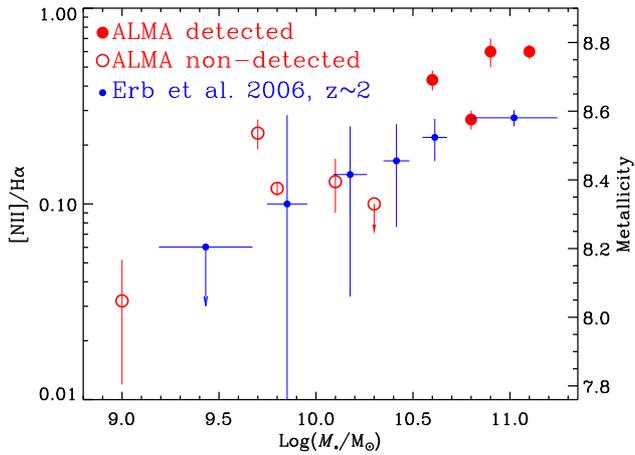
where  $X = \log_{10}(M_*/10^{10} M_{\odot})$ . This correlation between  $H\alpha$  extinction and stellar mass has also been confirmed by Sobral et al. (2012), Koyama et al. (2019) and Qin et al. (2019). Using the stellar masses presented by in Table 2, we derive the extinction-corrected  $H\alpha$  luminosity ( $L_{\text{H}\alpha}^{\text{corr}}$ ) to obtain the SFR:

$$\log \text{SFR}_{H\alpha} (M_{\odot} \text{ yr}^{-1}) = \log L_{\text{H}\alpha}^{\text{corr}} (\text{erg s}^{-1}) - 41.27, \quad (3)$$

assuming the same Chabrier IMF.

A comparison between these two SFR estimates is shown in Fig. 3. Previous studies of the SFR from FIR and  $H\alpha$  show that SMGs at  $z \sim 2$  have much larger  $\text{SFR}_{\text{FIR}}$  than  $\text{SFR}_{H\alpha}^{\text{corr}}$  (Swinbank et al. 2004; Chen et al. 2020).

Due to the sensitive  $H\alpha$  selection of our HiZELS parent sample, our sample comprises more typical ‘main sequence’ galaxies, and is not limited to the most dusty starbursts. Only one of the ALMA-SHiZELS galaxies has FIR-derived SFR above

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**Figure 4.** The stellar mass versus  $[\text{N II}]/\text{H}\alpha$  for our ALMA-SHiZELS galaxies. The ALMA-detected targets are shown in red while the non-detected targets are coloured blue. Since the metallicity could be derived from the  $[\text{N II}]/\text{H}\alpha$  ratio (Curti et al. 2017), we show the metallicity on the right-hand axis. We also show the mass-metallicity results at redshift  $\sim 2$  from Erb et al. (2006) in blue dots. Only the more massive galaxies with high metallicities are detected in continuum by ALMA.

$100M_{\odot} \text{ yr}^{-1}$ . It is therefore not surprising that the SFRs derived from  $\text{H}\alpha$  and FIR emission are more similar for our sample than for literature SMGs.

### 2.5 The role of stellar mass and gas-phase metallicity

We consider the gas-phase oxygen abundance as a proxy for metallicity, and estimate it from the  $[\text{N II}]/\text{H}\alpha$  emission line ratio (Pettini & Pagel 2004; Curti et al. 2017) from the near-IR IFU spectroscopy (see Table 3). The SHiZELS population has mainly Solar to sub-Solar values (Swinbank et al. 2012b; Molina et al. 2017). We show the mass-metallicity relation of our ALMA-SHiZELS sample in Fig. 4 and identify the ALMA-detected and non-detected targets. We show the mass-metallicity relation at redshift  $\sim 2$  from Erb et al. (2006) as a comparison. In our observations, we find that the lower the metal content or stellar mass, the weaker the ALMA continuum emission. Indeed, all of the ALMA-detected galaxies have  $[\text{N II}]/\text{H}\alpha \gtrsim 0.2$  or  $M_{*} > 10^{10.5}M_{\odot}$ , corresponding to  $12+\log(\text{O}/\text{H}) > 8.5$ . The typical  $[\text{N II}]/\text{H}\alpha$  emission line ratio for galaxies at  $z \sim 2$  ranges from 0.03 to 0.3 (Strom et al. 2017), thus our ALMA-detected sample is indeed biased towards massive and high-metallicity galaxies. As expected, dust continuum is more easily detected for a galaxy with higher stellar mass or metallicity.

The ALMA detected targets have stellar mass values about one order of magnitude higher than the ALMA non-detected targets. Since the SFR of the ALMA-SHiZELS sample is in the range of 10 to  $100 M_{\odot} \text{ yr}^{-1}$  (within one order of magnitude; Gillman et al. 2019), the ALMA detected targets in this work mainly have low specific SFR compared to the rest of the SHiZELS sample.

Using the metallicity measurements, we can also derive the gas-to-dust mass ratio ( $\delta_{\text{GDR}}$ ) following Magdis et al. (2012):  $\log \delta_{\text{GDR}} = (10.54 \pm 1.0) - (0.99 \pm 0.12) \times [12 + \log(\text{O}/\text{H})]$ . We estimate ISM mass using the ALMA-derived  $355 \mu\text{m}$  flux, following Scoville et al. (2016) (Eq 16). Our derived dust masses are presented in Table 3. As shown there, these estimates are in good agreement with the values derived from SED fitting.

### 2.6 Dust radius of the ALMA detected targets

To estimate the dust radius, we fit the  $uv$ -real visibility of the ALMA detected targets by CASA task `UVMODELFIT`, and show the results in Fig. 5. We chose the *disc* model in `UVMODELFIT` to fit the  $uv$  profile. The fitted radii of our targets are shown in Table 2. Since the FIR morphology is not regular, and thus the disk fitting results may not represent the flux distribution, we also measure the half-light radius of the ALMA images from the surface brightness distribution in the image plane and show the results in Table 2. The FIR surface brightness profiles in real plane are also fitted by Sérsic function and can be found in Section B. The different approaches adopted for the measurements of the dust radii all show that the radii are  $4\text{--}5 \pm 0.5$  kpc, larger than the typical size of the high-redshift SMGs (Ikarashi et al. 2015; Simpson et al. 2015; Fujimoto et al. 2017). Dust emission from the four ALMA-detected galaxies is found to extend up to  $\sim 2''$  ( $\sim 16$  kpc, Fig. B1).

## 3 RESULTS AND DISCUSSION

### 3.1 Individual galaxy properties

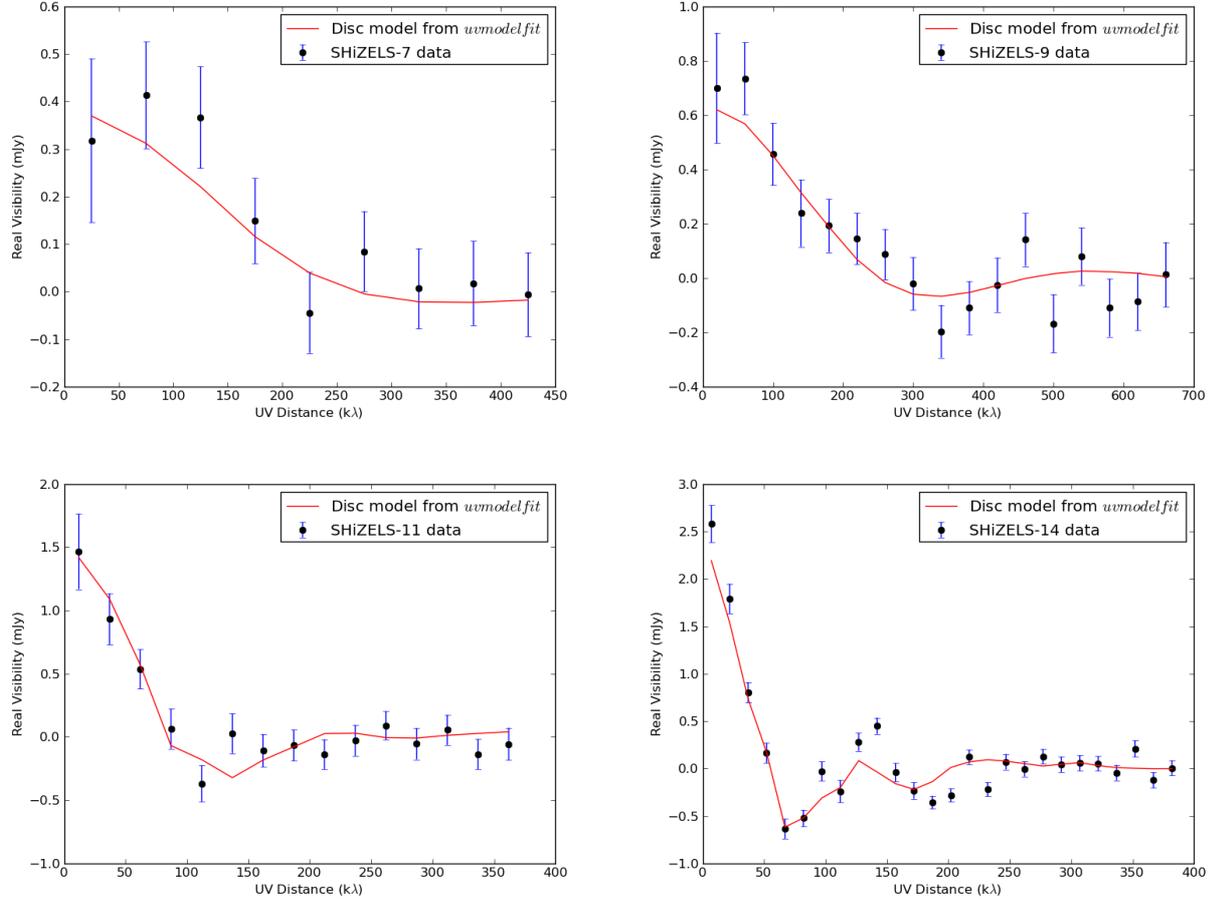
Figs 1 and A1 and A2 show the *HST*, SINFONI and ALMA images of our SHiZELS sample. For the galaxies at redshift about 1.47 and 2.23, the *HST* F606W filter corresponds to rest-frame  $2400 \text{ \AA}$  and  $1800 \text{ \AA}$ , therefore these images are tracing mainly the detectable UV emission, i.e. the star formation that is not obscured by the dust. On the other hand, the observed *HST* F140W image reveals the rest-frame *r*- or *g*-band morphology (free of strong line emission), which is more sensitive to stellar mass. Thus Figs 1, A1 and A2 show proxies for the stars (F140W), UV-traced star formation (F606W),  $\text{H}\alpha$ -traced star formation (SINFONI), and dust mass distributions (ALMA). In this section, we describe properties of our ALMA-detected targets individually<sup>2</sup>.

We caution that morphological information derived from the rest-frame UV/optical can be potentially very misleading in sources with substantial dust extinction and on-going star formation and in particular it is difficult to trace the true stellar mass (see Lang et al. 2019).

- **SHiZELS-7:** The  $\text{H}\alpha$  kinematics reveal the presence of an extended disc-like rotating structure. Two  $\text{H}\alpha$  star-forming clumps, separated by 4 kpc, are identified by Swinbank et al. (2012a), although the fainter clump is not seen in Gillman et al. (2019) because of the different S/N criteria. The FIR continuum is spread over a diffuse structure. The two  $\text{H}\alpha$  clumps are not likely to be affected by significant dust obscuration, so there might be no massive dust and gas clouds associated with the  $\text{H}\alpha$  features. The *HST* F140W and F606W images show that the least obscured regions in this galaxy show a compact stellar core and an extended UV morphology.

- **SHiZELS-9:** shows an extended rotation-dominated structure in  $\text{H}\alpha$  with three bright clumps (separated by  $\sim 3$  kpc, Swinbank et al. 2012a). The FIR continuum shows a V-like structure in the central 3 kpc, linking a bright stellar core with a fainter one. A third stellar core shows both UV and FIR emission.

<sup>2</sup> Most of our targets have two versions of the  $\text{H}\alpha$  maps (Swinbank et al. 2012a; Gillman et al. 2019). The main difference in the  $\text{H}\alpha$  morphology is caused by the different S/N criteria to create the  $\text{H}\alpha$  maps. Gillman et al. (2019) built  $\text{H}\alpha$  maps with high S/N spectra while the results in Swinbank et al. (2012a) use all the available spectra, which would then trace more extended  $\text{H}\alpha$  emission.



**Figure 5.** The  $uv$ -real visibility diagrams of the ALMA detected targets. The black dots and blue error bar are obtained from the ALMA data while the red lines are the fitting results by disk model in the CASA task `UVMODELFIT`.

- **SHiZELS-11:** This source was classified as a disk galaxy with a compact  $H\alpha$  structure (Swinbank et al. 2012a); although a newer analysis shows a marginally extended morphology (Gillman et al. 2019). SHiZELS-11 has an active galactic nucleus (AGN) identified in the X-rays by *XMM-Newton* and in the radio by the Very Large Array (VLA — Simpson et al. 2006; Ueda et al. 2008). SHiZELS-11 is also detected by *Chandra* in X-UDS survey (Kocevski et al. 2018), with X-ray luminosity is  $L_X = 10^{43.3}$  erg  $s^{-1}$ . The velocity dispersion of  $H\alpha$  is about 90 km  $s^{-1}$  (Gillman et al. 2019), much less than the typical velocity dispersion of type-I AGN, implying that the AGN in SHiZELS-11 must be obscured. Rotation is not clear from the  $H\alpha$  dynamics, and the position of the maximum velocity dispersion is offset from the brightest  $H\alpha$  pixel (Gillman et al. 2019). The *HST* F140W image shows a smaller source to the south, which may suggest a close merger or dust lane. The source has an apparently high metallicity and a steep metallicity gradient (Swinbank et al. 2012a), which might be due to contamination from an AGN. Both UV and FIR are bright in the galaxy centre, while the rest-frame optical is mainly coming from a clumpy structure surrounding the centre. The ALMA continuum shows compact and extended emission, which does not overlap with the  $H\alpha$ .

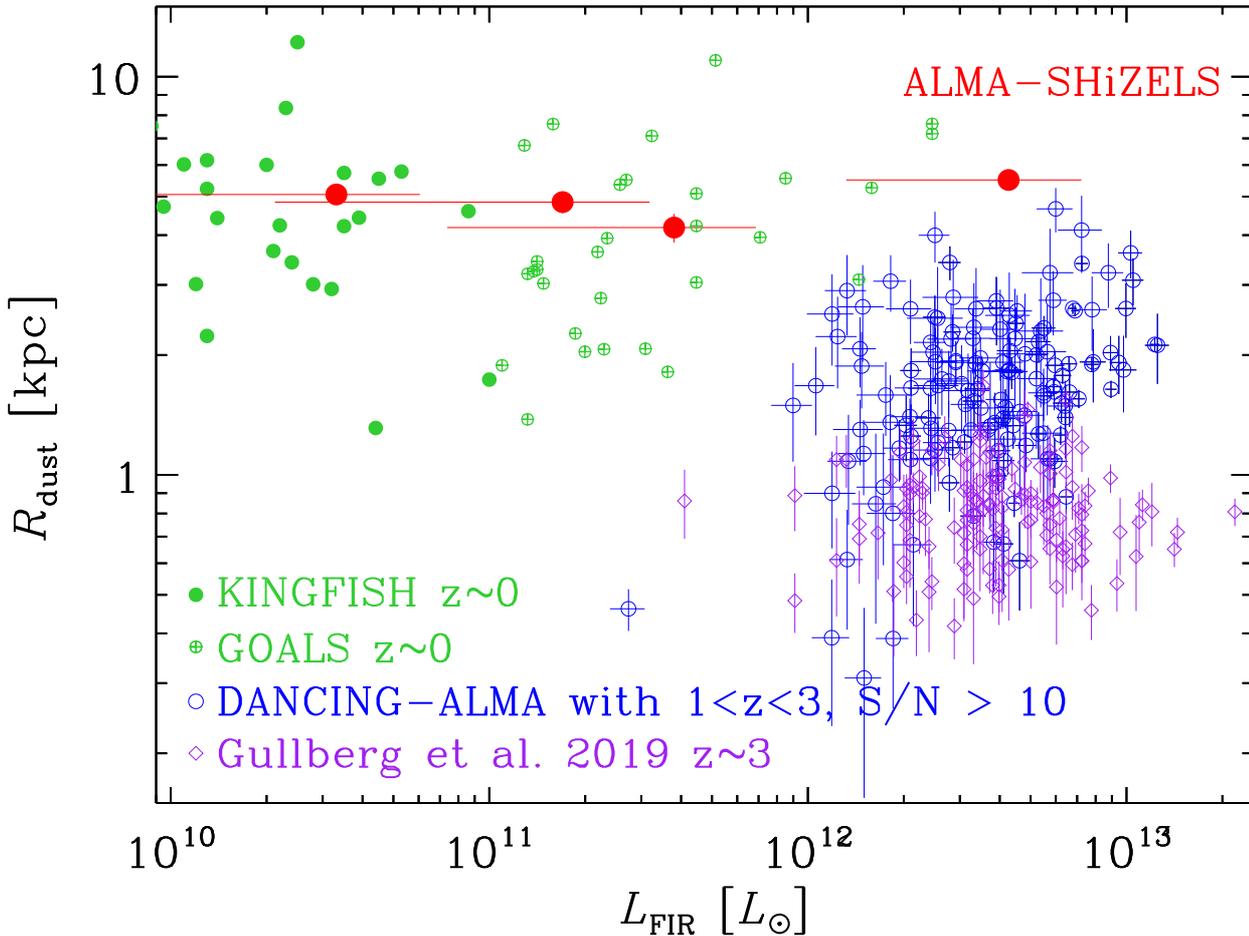
- **SHiZELS-14:** This galaxy has been identified as a merger with three  $H\alpha$  clumps separated by  $\sim 5$  kpc (Swinbank et al. 2012b). The  $H\alpha$  morphology in Gillman et al. (2019) is less clumpy, but still extends to a half-light radius of about 7 kpc (Sersic model fitting

results). The  $H\alpha$  kinematics show a velocity-dispersion-dominated system. The galaxy has the highest metallicity in the SHiZELS sample of Swinbank et al. (2012a). The full extent of the  $H\alpha$  emission is not traced by the rest-frame UV imaging. The ALMA continuum emission shows one dominant compact component at the centre, and more extended emission following a similar orientation as the F140W morphology. The FIR emission spreads up to  $2''$ , i.e.  $\sim 16$  kpc at  $z = 2.23$ . This target is comprehensively studied by Cochrane et al. (in preparation), including new Jansky VLA data.

## 3.2 Spatially resolved properties

### 3.2.1 The dust emission from $z = 1.47$ & 2.23 galaxies

Previous sub-millimetre studies of high-redshift galaxies at high spatial resolution have revealed that clumps at  $\sim 1$  kpc scales do not particularly dominate the total flux emission, as is often seen in UV or  $H\alpha$  imaging (Hodge et al. 2019). For example, Hodge et al. (2016) used ALMA to observe a sample of  $z \sim 2.5$  submillimetre galaxies (SMGs) at  $0''.16$  resolution and found that the observed  $870\mu\text{m}$  continuum morphologies are predominantly smooth and disc-like with typical radii of  $\sim 1.8$  kpc. Gullberg et al. (2019) also shows that dust emission size from SMGs is about  $\sim 1.5$  kpc at  $z \sim 1-4$ , with a larger sample.



**Figure 6.** Comparison between different surveys, local and high- $z$ , with previous spatially resolved FIR observations. We show our ALMA-SHiZELS galaxies (red dots), the galaxy sample from KINGFISH (green filled circles [Kennicutt et al. 2011](#)), GOALS (green cross filled circles [Armus et al. 2009](#); [Chu et al. 2017](#)), DANCING-ALMA with  $S/N$  of  $L_{\text{FIR}}$  higher than 5 (cross filled blue circles [Fujimoto et al. 2017](#)) and the recent ALMA observation results of the SCUBA-2 bright galaxies ([Gullberg et al. 2019](#), purple diamonds). The figure shows that our ALMA-SHiZELS are fainter than observations of typical submillimetre bright galaxies at similar redshifts, and also present larger sizes. The sizes of the SHiZELS galaxies compare better with those of low-redshift star-forming galaxies from KINGFISH and GOALS.

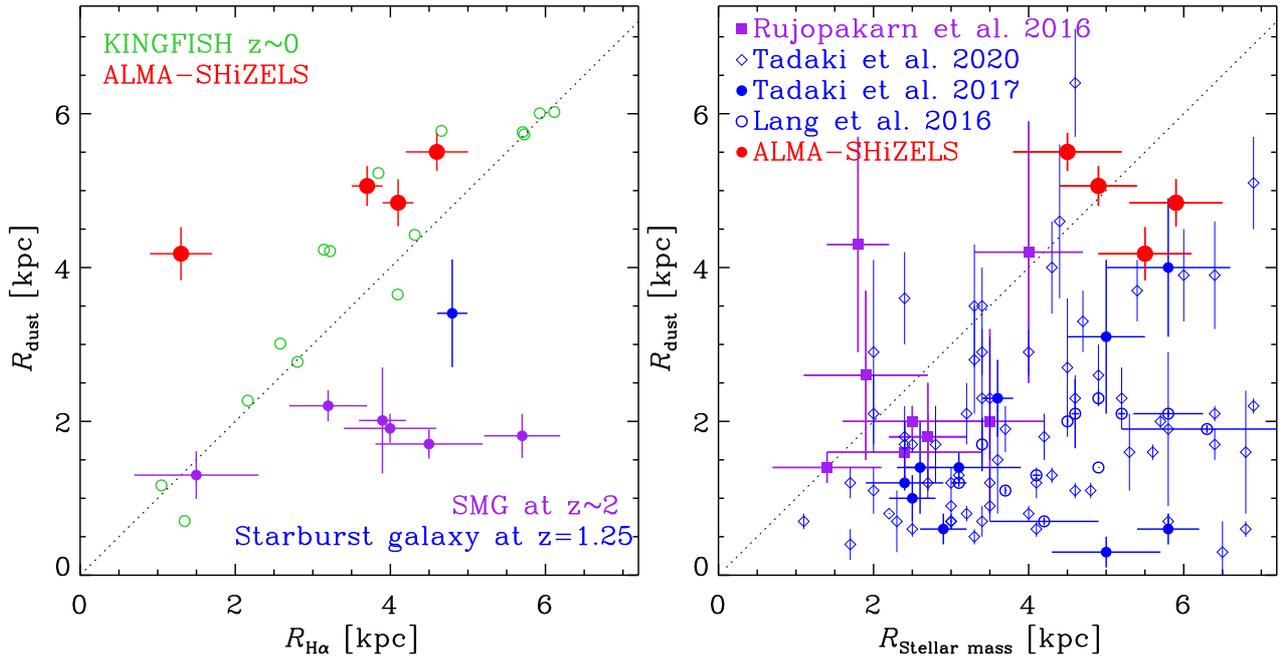
Although these previous studies have revealed the internal properties of the brightest SMGs, our focus is on the normal star forming galaxy population (i.e. galaxies selected by  $H\alpha$  flux, with many having  $\lesssim L_{\text{H}\alpha}^*$  [Swinbank et al. 2012a](#)) at high redshift. We find that our sample shows sub-millimetre continuum emission which is  $\sim 2\times$  more extended than that seen in SMGs at  $z \sim 3$ . In the local universe, it is well known that ULIRGs present compact star-formation cores, which may eventually lead to a compact stellar core. The large FIR radius of our ALMA-SHiZELS observations reveal the existence of a high-redshift galaxy population with extended star-formation activity at  $z = 1.47$  or  $2.23$ , possibly sustained in thick rotating disc-like structures.

In Fig. 6 we compare our results to previous spatially resolved FIR studies. For the galaxies at  $z = 0$  and  $L_{\text{FIR}} < 10^{11} L_{\odot}$ , the Key Insights on Nearby Galaxies: a Far Infrared Survey with *Herschel* project (KINGFISH; [Kennicutt et al. 2011](#)) surveys have found a typical radius of about  $\sim 4$  kpc with predominately disc-like morphologies ([Mosenkov et al. 2019](#)). For isolated LIRGs taken from the Great Observatories All-sky LIRG Survey (GOALS, [Armus et al. 2009](#); [Chu et al. 2017](#)) we also find relatively large FIR

sizes, similar to those from KINGFISH. To have enough spatial resolution, we measure the GOALS galaxy FIR size based on the *Herschel* PACS  $160\mu\text{m}$  image. The difference between the FIR radius in  $160\mu\text{m}$  and  $350\mu\text{m}$  is about 1.5 times or less ([Mosenkov et al. 2019](#)), which will not affect our results. Some of the LIRGs in GOALS sample are galaxy pairs or merging galaxies. To avoid the contamination from the neighbours, we only make use of the isolated galaxy sample here.

On the other hand, Demonstrating a New Census of Infrared Galaxies with ALMA (DANCING-ALMA, [Fujimoto et al. 2017](#)) project presents  $\sim 1000$  galaxies ( $0 < z < 6$  mainly ULIRGs) resolved by ALMA at 1 mm and finds a positive correlation between the radius and  $L_{\text{FIR}}$  (in Fig. 6 we only chose  $1 < z < 3$ ,  $S/N > 10$  galaxies). The recent work by [Gullberg et al. \(2019\)](#) shows compact FIR emission for  $\sim 150$  SMGs, when the IR luminosity of the galaxy exceeds  $10^{12} L_{\odot}$ .

Our ALMA-SHiZELS sample is different from the previous luminous high redshift dusty populations as we are targeting a much fainter population. As shown in Fig. 6 and Tab. 2, our sample has typically lower IR luminosity, larger dust size than the high-



**Figure 7.** **Left panel:** Comparison of the half-light radii measured in the ALMA continuum imaging with respect to those measured in  $H\alpha$  (Swinbank et al. 2012b; Molina et al. 2017). We also overplot the *Herschel* 350- $\mu\text{m}$  and narrow-band  $H\alpha$  radii (obtained after convolving both to the same resolution) from the KINGFISH (Kennicutt et al. 2011) and *Spitzer* Infrared Nearby Galaxies Survey (SINGS) projects (Kennicutt et al. 2003; Dale et al. 2005, 2007). Our high-redshift galaxy sample shows a broadly similar behaviour as is seen in these local galaxies, except SHiZELS-11 which shows a significantly larger radius in the FIR than in  $H\alpha$ . We also show the recent ALMA-resolved SMG sample at  $z \sim 2$  (Chen et al. 2020) and one starburst galaxy (GOODSN-18574) at  $z = 1.25$  (Nelson et al. 2019). **Right panel:** The stellar mass size as derived from the *HST* F140W image compared to the FIR size as derived from the ALMA observations. We show the  $H\alpha$  selected main sequence star-forming galaxies at redshift about 2.2-2.5 by Tadaki et al. (2017), the star-forming galaxies presented by Rujopakarn et al. (2016), the  $z \sim 2$  SMGs from Lang et al. (2019), and the massive star-forming galaxy sample at  $z \sim 2$  from Tadaki et al. (2020). The dot line shows the one-to-one relation. Our targets are roughly the largest galaxies in both rest-frame optical band and FIR radius, and may evolve into massive disk galaxies.

redshift SMGs, which suggests that our survey are targeting the high-redshift star-forming galaxy population, rather than the high-redshift starburst galaxies.

The extension of the dust emission ( $R_{\text{dust}}$ ) could be roughly estimated from a theoretical point of view, assuming  $L_{\text{FIR}} = 4\pi R_{\text{eff}}^2 \sigma T_{\text{mbb}}^{4.32}$ , where  $\sigma$  is the Stefan-Boltzmann constant, and the dust temperature ( $T_{\text{mbb}}$ ) and  $L_{\text{FIR}}$  are derived from SED templates (see § 2.4; Ma & Yan 2015; Yan & Ma 2016; Dudzevičiūtė et al. 2020). This effective radius can be considered as the lower size limit of the galaxy FIR emitters, hence the difference with respect to the half-light radius could be understood as the spatial occupation of dust structures below the resolution elements. Using this equation, we derive effective radii for our SHiZELS galaxies of 0.3–1.1 kpc, much smaller than the observed half-light radii of about 4.5 kpc, implying typical dust filling factors from 5 to 10 for our sample. Hodge et al. (2019) have found that at 500 pc resolution they can see tentative evidence of the spiral and bar structures of SMGs ( $z = 1.5 - 4.9$ ) at  $\sim 250 \mu\text{m}$  (rest-frame). Subject to surface brightness limitations, higher resolution ALMA imaging might reveal similarly complex structures in our galaxies.

### 3.2.2 Spatial correlation between dust and ionised gas

The observed morphologies of dust and  $H\alpha$  emissions are well known to be correlated (at kpc scales) for samples of local star-forming galaxies, such as the KINGFISH project (Kennicutt et al. 2011). In the left panel of Fig. 7 we show that if we compare the half-light radii measured by ALMA with respect to the  $H\alpha$  emission, we

see good agreement for sources SHiZELS-7, -9 and -14, suggesting spatial coexistence (at  $\sim$  kpc scales) of the dust and the ionised gas. In contrast, the source presenting an AGN, SHiZELS-11, shows a significantly larger FIR radius than that seen in  $H\alpha$ .

We compare our results to the recent high resolution ALMA and  $H\alpha$  observations of a sample of  $z \sim 2$  SMGs (Chen et al. 2020) and one starburst galaxy (Nelson et al. 2019) in the left-hand panel of Fig. 7. We find that our sample has a similar  $H\alpha$  radius to previously observed SMGs, suggesting extended star formation for these high-redshift galaxies. Despite this agreement, the dust emission from SMGs is typically more compact than our galaxies (Gullberg et al. 2019). The starburst galaxy in Nelson et al. (2019) has a dust radius higher than the SMGs, but still lower than our sample.

Although the SHiZELS sample shows a larger dust radius than the SMGs and starburst galaxy, most of the high-redshift galaxies in Fig. 7 have similar  $H\alpha$  radius of about 4 kpc. So the SMGs appear also to have an extended star formation region, as well as a compact dusty core in the galaxy center (e.g., Fig. 6). SMGs may obscure  $H\alpha$  emission especially in the center kpc region, so the  $H\alpha$  distribution in the SMG center might be more flatter, leading to a larger  $H\alpha$  half-light radius. Since  $H\alpha$  in SMGs can be bright and extended (Swinbank et al. 2004, 2006), while the dust morphology is compact (Rujopakarn et al. 2019; Chen et al. 2020), the large dust radii we observe in ALMA-SHiZELS sample may be caused by the low SFR ( $L_{H\alpha} < L_{H\alpha}^*$ ) of the SHiZELS sample, rather than with the  $H\alpha$ -selection method.

The strong dust extinction in SMGs may also affect the  $H\alpha$

emission such that the SFR derived from  $H\alpha$  is lower than the  $SFR_{FIR}$  (Fig. 3), even though both SFR indicators may only reflect one aspect of the star formation and may still lower than the intrinsic SFR in galaxies.

The SHiZELS sample was selected on  $H\alpha$ , down to a flux limit that traces below the knee of the luminosity function, and therefore the bulk of the sample will be less dusty than submm selected galaxies. Our results show that the SHiZELS galaxies have very similar dust and  $H\alpha$  radii as the star-forming galaxies from KINGFISH at low redshift (Fig. 6 and 7). In the low-redshift universe, most of the star-forming galaxies are found to be spiral galaxies with low star-formation efficiency over an extended large disk (e.g., Leroy et al. 2008; Cheng et al. 2018), while ULIRGs usually have a more violent compact nuclear starburst in the galaxy center (Solomon et al. 1997; Downes & Solomon 1998). The similarity between the bulk of the ALMA-SHiZELS sample and the KINGFISH sample in Fig. 6 and 7 is in line with the two samples tracing more similar modes of main sequence star formation, compared with the extreme starburst modes witnessed in SMGs.

$H\alpha$  clumpy features are commonly found in high-redshift galaxies (Genzel et al. 2006; Swinbank et al. 2012b). However, our ALMA continuum images do not show clear counterparts of the  $H\alpha$  clumps, at least at  $\sim$  kpc scales (Fig. 1). It may be that the observed clumps are not tracing true star-forming clumps, but less dusty regions of the galaxy, from which the  $H\alpha$  can escape (Swinbank et al. 2004).

### 3.2.3 A comparison between the HST and ALMA morphologies

The HST F140W images of our sample which trace the rest-frame optical emission (Fig. A1 and A2), reveal that only SHiZELS-9 has an apparently dual stellar core (the two cores have similar fluxes), which may be evidence of a major merger, or caused by a dust lane. SHiZELS-11 and -14 may both be undergoing or recently have undergone a minor merger event.

Spatial offsets between the rest-frame optical and FIR observations have been found in low and high-redshift (U)LIRGs (Charmandaris et al. 2004; Hodge et al. 2016). In Fig. 1 we see the spatial offsets and the overlap between the F140W and ALMA emission. For SHiZELS-9, the two major (rest-frame)  $V$ -band cores present dust emission, while additional FIR emission comes from a region connecting to the third  $V$ -band core. The  $V$ -band flux distribution does not always follow the  $H\alpha$  or the dust morphology in our sample. A detailed study of spatially resolved dust extinction maps and the IRX- $\beta$  relation (the FIR and UV luminosity ratio versus the UV SED slope, Meurer et al. 1999) of our high-redshift star forming galaxies will be presented in a forthcoming paper.

The rest-frame UV images also show complex morphologies. Compared to the detected FIR emission from SHiZELS galaxies, the UV morphologies are extended and clumpy. For the ALMA non-detected galaxies, SHiZELS-10, -2, 3, and -21, the rest-frame UV and optical band images show compact cores at their centres (Fig. A2), evidencing compact UV star formation cores or regions of lower obscuration which is also found in the low-redshift low mass galaxies (Cheng et al. 2020). We see that  $z \approx 2.23$  galaxies display more compact rest-frame UV morphologies than the  $z \approx 1.47$  sample, which is consistent with previous studies (Paulino-Afonso et al. 2017).

We show the half-light radii of the F140W emission in the right panel of Fig. 7. The F140W radii (adopted from Gillman et al. 2019)) have typical effective radii of  $\sim 4.5$  kpc, consistent with the typical size of the full HiZELS sample (Stott et al. 2013). Us-

ing the stellar masses for our targets to predict half-light radius based on the mass-size relation (Suess et al. 2019), we find that the size revealed by the F140W imaging are consistent with the predicted radius within  $1\sigma$ . We also present the galaxy sample from Rujopakarn et al. (2016) composed of star-forming galaxies at redshift  $\sim 2$ , observed by ALMA ( $870 \mu\text{m}$ , 1.3 mm) and VLA (5 cm) at  $0''.4$ , the  $H\alpha$  selected main sequence star-forming galaxies from Tadaki et al. (2017), the SMGs at redshift 2 observed by ALMA  $870 \mu\text{m}$  from Lang et al. (2019), and the recent ALMA  $870 \mu\text{m}$  observation results of the massive star-forming galaxies at redshift 2 (Tadaki et al. 2020). In the right panel of Fig. 7, Tadaki et al. (2017)'s sample of proposed elliptical progenitor galaxies has systematically smaller FIR size similar to other SMGs, suggesting the future formation of compact red cores. Our targets are roughly the largest galaxies in both rest-frame optical band and FIR radius. The large FIR radius indicate the stellar mass is assembling at a larger radius, thus our targets would evolve into massive disk galaxies.

## 4 CONCLUSION

We present high-resolution ( $0''.25$ ) ALMA continuum observations (rest-frame  $\sim 355 \mu\text{m}$ ) of nine star-forming galaxies at redshift 1.47 and 2.23, taken from the HiZELS survey. These galaxies have been observed with previous AO-aided IFU  $H\alpha$  spectroscopy and HST at similar resolution, facilitating a spatial exploration of the star formation at high-redshift at kpc scale resolution. Our sample comprises a population of star-forming galaxies at redshift 1.47 and 2.23, which are mainly 'main sequence' galaxies, and have  $H\alpha$  luminosities close to or below  $L_{H\alpha}^*$  at their redshift.

We detect four out of nine galaxies with ALMA. Their morphologies present extended faint structures, out to 16 kpc in diameter, much larger than the typical FIR size of sub-millimetre galaxies at high redshift. Our ALMA observations also reveal two serendipitous detections within the ALMA primary beams. Both of these are detected in previous optical and NIR surveys, located at different redshifts. For the ALMA non-detected galaxies, we find they mainly have either lower stellar masses ( $M_* < 10^{10.5} M_\odot$ ) and lower  $[N II]/H\alpha < 0.25$  ratios, corresponding to  $12+\log(O/H) < 8.5$ .

The SFR derived from FIR and dust corrected  $H\alpha$  are consistent with each other within  $1\sigma$ . However, the spatially resolved FIR and  $H\alpha$  morphologies do not show a similar distribution, implying a complex distribution of the ISM state in these galaxies. At least at  $\sim$  kpc scales, the  $H\alpha$  emission do not show a clear spatial correlation with respect to the FIR emission. We find that the brightest  $H\alpha$  clumps, previously identified via AO-aided IFU spectroscopy, are not significantly spatially correlated with the dust continuum emission which traces the bulk of the ISM, which appears smooth.

Our sample of SHiZELS galaxies have a typical rest-frame FIR size twice larger than the SMGs at high-redshift, while the  $H\alpha$  emission size of our sample and SMGs are similar. The similarity between the extent of the dust continuum emission, and IR luminosities of our sample and the local star-forming disk galaxies is in line with our sample being drawn from the 'normal' star-forming galaxy population at  $z=1.5-2.2$ . The dust continuum emission in these systems is dominated by a cool extended component, while in more active galaxies such as SMGs, compact nuclear starbursts dominate the dust emission (Gullberg et al. 2019)

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## REFERENCES

Aihara, H., AlSayyad, Y., Ando, M., et al. 2019, PASJ, 106  
 Armus, L., Mazzarella, J. M., Evans, A. S., et al. 2009, PASP, 121, 559  
 Beelen, A., Cox, P., Benford, D., et al. 2006, ApJ, 642, 694  
 Boquien, M., Kennicutt, R., Calzetti, D., et al. 2016, A&A, 591, A6  
 Calura, F., Pozzi, F., Cresci, G., et al. 2017, MNRAS, 465, 54  
 Casasola, V., Bianchi, S., De Vis, P., et al. 2020, A&A, 633, A100  
 Chabrier, G. 2003, PASP, 115, 763  
 Charmandaris, V., Le Floch, E., & Mirabel, I. F. 2004, ApJ, 600, L15  
 Chen, C.-C., Harrison, C. M., Smail, I., et al. 2020, A&A, 635, A119  
 Cheng, C., Ibar, E., Hughes, T. M., et al. 2018, MNRAS, 475, 248  
 Cheng, C., Xu, C. K., Xie, L., et al. 2020, A&A, 633, A105  
 Chu, J. K., Sanders, D. B., Larson, K. L., et al. 2017, ApJS, 229, 25  
 Clark, C. J. R., Verstocken, S., Bianchi, S., et al. 2018, A&A, 609, A37  
 Cochrane, R. K., Best, P. N., Sobral, D., et al. 2017, MNRAS, 469, 2913  
 Cochrane, R. K., Best, P. N., Sobral, D., et al. 2018, MNRAS, 475, 3730  
 da Cunha, E., Charlot, S., & Elbaz, D. 2008, MNRAS, 388, 1595  
 Curti, M., Cresci, G., Mannucci, F., et al. 2017, MNRAS, 465, 1384  
 Dale, D. A., Bendo, G. J., Engelbracht, C. W., et al. 2005, ApJ, 633, 857  
 Dale, D. A., Gil de Paz, A., Gordon, K. D., et al. 2007, ApJ, 655, 863  
 Downes, D. & Solomon, P. M. 1998, ApJ, 507, 615  
 Dudzevičiūtė, U., Smail, I., Swinbank, A. M., et al. 2020, MNRAS, 494, 3828  
 Erb, D. K., Shapley, A. E., Pettini, M., et al. 2006, ApJ, 644, 813

Franceschini, A., Berta, S., Rigopoulou, D., et al. 2003, A&A, 403, 501  
 Fomalont, E., van Kempen, T., Kneissl, R., et al. 2014, The Messenger, 155, 19  
 Förster Schreiber, N. M., Genzel, R., Lehnert, M. D., et al. 2006, ApJ, 645, 1062  
 Förster Schreiber, N. M., Genzel, R., Bouché, N., et al. 2009, ApJ, 706, 1364  
 Förster Schreiber, N. M., Renzini, A., Mancini, C., et al. 2018, ApJS, 238, 21  
 Fujimoto, S., Ouchi, M., Shibuya, T., & Nagai, H. 2017, ApJ, 850, 83  
 Genzel, R., Tacconi, L. J., Eisenhauer, F., et al. 2006, Nature, 442, 786  
 Genzel, R., Burkert, A., Bouché, N., et al. 2008, ApJ, 687, 59  
 Garn, T., & Best, P. N. 2010, MNRAS, 409, 421  
 Geach, J. E., Smail, I., Best, P. N., et al. 2008, MNRAS, 388, 1473  
 Gillman, S., Swinbank, A. M., Tiley, A. L., et al. 2019, MNRAS, 486, 175  
 Gullberg, B., Smail, I., Swinbank, A. M., et al. 2019, MNRAS, 490, 4956  
 Hodge, J. A., Swinbank, A. M., Simpson, J. M., et al. 2016, ApJ, 833, 103.  
 Hodge, J. A., Smail, I., Walter, F., et al. 2019, ApJ, 876, 130  
 Hwang, H. S., Elbaz, D., Magdis, G., et al. 2010, MNRAS, 409, 75  
 Ibar, E., Sobral, D., Best, P. N., et al. 2013, MNRAS, 434, 3218  
 Ikarashi, S., Ivison, R. J., Caputi, K. I., et al. 2015, ApJ, 810, 133  
 Katsianian, A., Blanc, G., Lagos, C. P., et al. 2017, MNRAS, 472, 919  
 Kennicutt, R. C. 1998, ARA&A, 36, 189  
 Kennicutt, R. C., Jr., Armus, L., Bendo, G., et al. 2003, PASP, 115, 928  
 Kennicutt, R. C., et al. 2009, ApJ, 703, 1672  
 Kennicutt, R. C., Calzetti, D., Aniano, G., et al. 2011, PASP, 123, 1347  
 Kennicutt, R. C., Evans II, N. J. 2012, ARA&A, 50, 531  
 Kocevski, D. D., Hasinger, G., Brightman, M., et al. 2018, ApJS, 236, 48  
 Koyama, Y., Shimakawa, R., Yamamura, I., et al. 2019, PASJ, 71, 8  
 Lang, P., Schinnerer, E., Smail, I., et al. 2019, ApJ, 879, 54  
 Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, MNRAS, 379, 1599  
 Leroy, A. K., Walter, F., Brinks, E., et al. 2008, AJ, 136, 2782  
 Liang, L., Feldmann, R., Kereš, D., et al. 2019, MNRAS, 489, 1397  
 Ma, Z., & Yan, H. 2015, ApJ, 811, 58  
 Magdis, G. E., Daddi, E., Béthermin, M., et al. 2012, ApJ, 760, 6  
 Mehta, V., Scarlata, C., Capak, P., et al. 2018, ApJS, 235, 36  
 Meurer, G. R., Heckman, T. M., & Calzetti, D. 1999, ApJ, 521, 64  
 Molina, J., Ibar, E., Swinbank, A. M., et al. 2017, MNRAS, 466, 892  
 Molina, J., Ibar, E., Smail, I., et al. 2019, MNRAS, 487, 4856  
 Mosenkov, A. V., Baes, M., Bianchi, S., et al. 2019, A&A, 622, A132  
 Nelson, E. J., Tadaki, Ken-ichi, Tacconi, L. J., et al. 2019, ApJ, 870, 130  
 Oteo, I., Sobral, D., Ivison, R. J., et al. 2015, MNRAS, 452, 2018  
 Oteo, I., Smail, I., Hughes, T., et al. 2017, arXiv e-prints, arXiv:1707.05329.  
 Paulino-Afonso, A., Sobral, D., Buitrago, F., et al. 2017, MNRAS, 465, 2717  
 Pettini, M., & Pagel, B. E. J. 2004, MNRAS, 348, L59  
 Qin, J., Zheng, X. Z., Wuyts, S., et al. 2019, ApJ, 886, 28  
 Rujopakarn, W., Dunlop, J. S., Rieke, G. H., et al. 2016, ApJ, 833, 12  
 Rujopakarn, W., Daddi, E., Rieke, G. H., et al. 2019, ApJ, 882, 107  
 Scoville, N., Aussel, H., Brusa, M., et al. 2007, ApJS, 172, 1  
 Scoville, N., Sheth, K., Aussel, H., et al. 2016, ApJ, 820, 83  
 Shapley, K. L., Genzel, R., Förster Schreiber, N. M., et al. 2008, ApJ, 682, 231  
 Simpson, C., Martínez-Sansigre, A., Rawlings, S., et al. 2006, MNRAS, 372, 741  
 Simpson, J. M., Smail, I., Swinbank, A. M., et al. 2015, ApJ, 799, 81  
 Sobral D., Best P. N., Matsuda Y., Smail I., Geach J. E., Cirasuolo M., 2012, MNRAS, 420, 1926  
 Sobral D., Smail I., Best P. N., Geach J. E., Matsuda Y., Stott J. P., Cirasuolo M., Kurk J., 2013, MNRAS, 428, 1128  
 Sobral, D., Best, P. N., Smail, I., et al. 2014, MNRAS, 437, 3516  
 Sobral D. et al., 2015, MNRAS, 451, 2303  
 Sobral, D., & Matthee, J. 2019, A&A, 623, A157  
 Solomon, P. M., Downes, D., Radford, S. J. E., et al. 1997, ApJ, 478, 144  
 Stott, J. P., Sobral, D., Smail, I., et al. 2013, MNRAS, 430, 1158  
 Strom, A. L., Steidel, C. C., Rudie, G. C., et al. 2017, ApJ, 836, 164  
 Suess, K. A., Kriek, M., Price, S. H., et al. 2019, ApJ, 877, 103  
 Swinbank, A. M., Smail, I., Chapman, S. C., et al. 2004, ApJ, 617, 64  
 Swinbank, A. M., Chapman, S. C., Smail, I., et al. 2006, MNRAS, 371, 465  
 Swinbank, A. M., Sobral, D., Smail, I., et al. 2012a, MNRAS, 426, 935

Swinbank, A. M., Smail, I., Sobral, D., et al. 2012b, *ApJ*, 760, 130  
 Tadaki, Ken-ichi, Genzel, R., Kodama, T., et al. 2017, *ApJ*, 834, 135  
 Tadaki, Ken-ichi, Belli, S., Burkert, A., et al. 2020, arXiv:2009.01976  
 Thomson, A. P., Simpson, J. M., Smail, I., et al. 2017, *ApJ*, 838, 119  
 Ueda, Y., Watson, M. G., Stewart, I. M., et al. 2008, *ApJS*, 179, 124  
 Yan, H., & Ma, Z. 2016, *ApJ*, 820, 16

## APPENDIX A: STAMP IMAGES OF ALMA-HIZELS TARGETS

We present the *HST*, VLT/SINFONI, and ALMA images in Fig. A1 and Fig. A2.

## APPENDIX B: DUST SURFACE BRIGHTNESS PROFILES

We derive the dust continuum surface brightness profile by using the peak flux pixel as the centres of the tapered ALMA images. Then we perform aperture photometry by using a series of ring apertures in steps of  $0''.25$  (similar to the synthesised beam) in the natural-weighted ALMA images. Then we derive the surface brightness in each annulus for all of the four ALMA-detected galaxies, and show the results in Fig. B1. We also show the surface brightness profile of the ALMA synthesised beam, normalised to the maximum value for each source in Fig. B1. We fit the ALMA surface brightness profile with a Sersic function, and present the results in each panel. Sersic indices  $n$  of the ALMA images are lower than  $n = 2$ , implying a disk-like morphology of the dust emission. Based on the aperture photometry in each apertures, we also obtained the flux growth curve of each ALMA-detected galaxy, and derive the half-light radius in Tab. 2.

## APPENDIX C: ALMA NON-DETECTED TARGETS

### C1 Individual properties

- **SHiZELS-21**, also known as UDS-10 in Molina et al. (2017). This galaxy is a compact rotating galaxy. The  $H\alpha$  velocity dispersion profile is double peaked at about 1.5 kpc from the galaxy center. The *HST* F606W image shows a double UV core structure which is barely resolved in the F140W image.

- **SHiZELS-8**: This galaxy is dominated by rotation and present clumpy features (Swinbank et al. 2012a; Gillman et al. 2019). None of the three  $H\alpha$  clumps is detected by ALMA. Its metallicity is relatively low ( $12 + \log(\text{O}/\text{H}) < 8.3$ ) and the metallicity gradient is flat (Swinbank et al. 2012a). A simple dynamical description using carbon monoxide is presented by Molina et al. (2019). The *HST* F606W image shows an extended UV morphology, while the F140W image shows a compact core in the galaxy. The clumpy structures seen in the  $H\alpha$  map are not recovered by the F606W image.

- **SHiZELS-10**: This galaxy is compact in  $H\alpha$  ( $\sim 2.3$  kpc) and identified as a merger (Swinbank et al. 2012a; Gillman et al. 2019). The ALMA continuum emission is undetected, probably due to limitations in the surface brightness. On the other hand, the *HST* F140W, F606W and the VLT/SINFONI images show a good spatial consistency with bright dots and a long tail.

- **SHiZELS-2**: The  $H\alpha$  IFU observations evidence a clear rotation curve, including two clumps at the centre with a separation of  $0''.2$  (about 1.5 kpc Gillman et al. 2019). The fainter clump locates in the rotation center. The orientation of the *HST* morphology is

similar to that revealed by rotation in  $H\alpha$ . The *HST* images show a compact stellar and star formation distribution. The *HST* F606W image also shows that the star formation in the galaxy center bright in rest-frame UV (Gillman et al. 2019).

- **SHiZELS-3**: The  $H\alpha$  velocity map shows a clear rotation feature, while the velocity dispersion map shows two peaks separated by  $\sim 2.5$  kpc, suggesting a complex dynamics probably associated to merging activity (Gillman et al. 2019). *HST* images show that the stellar mass and rest-frame UV star formation morphologies are compact.

### C2 Stacking analyses

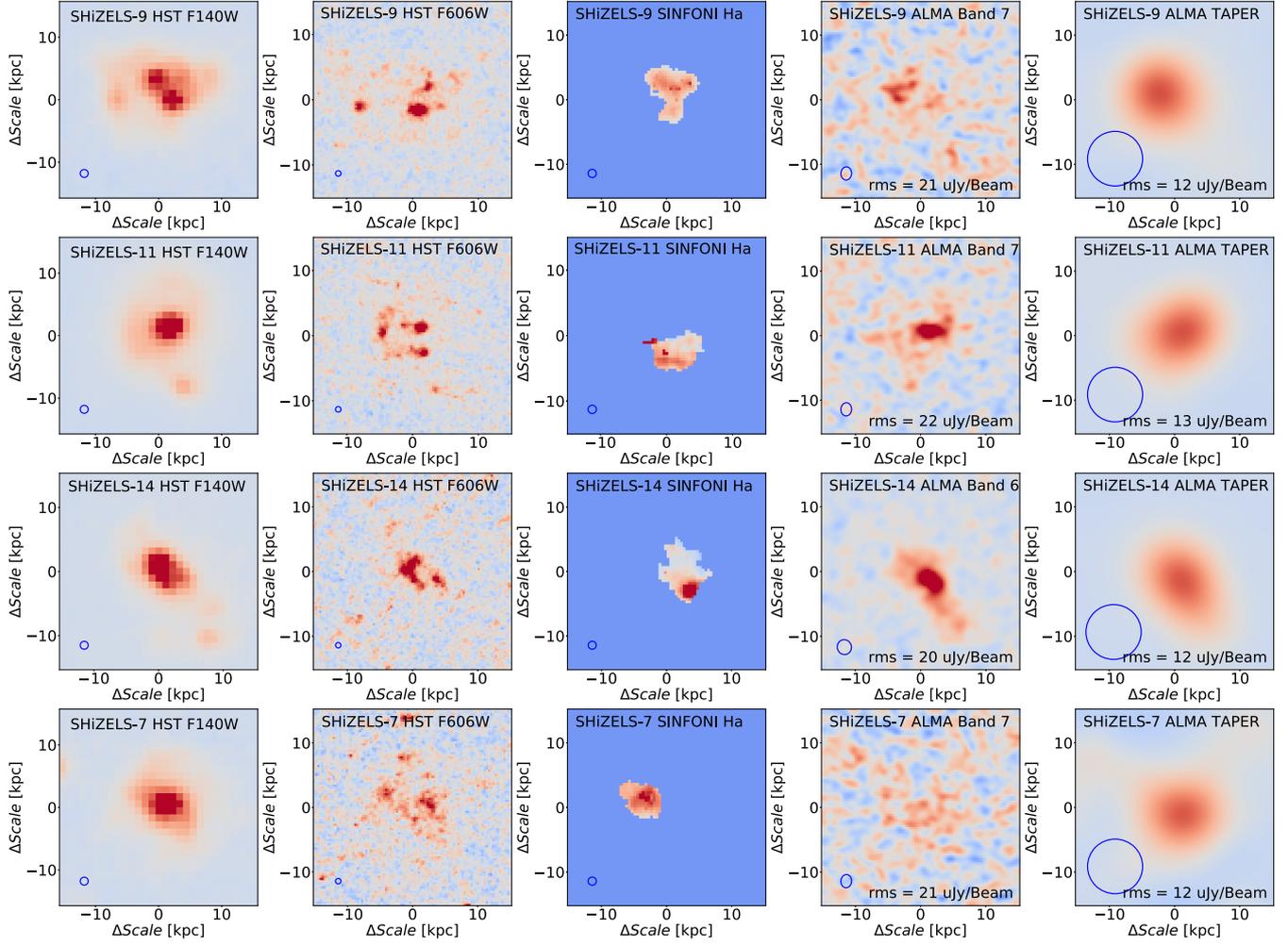
In this appendix we explore the possibility to extract information from the non-detected sources via a stacking approach. To do this, we stack the ALMA tapered image which are generated at  $1''0$  resolution. We generate postage stamps (of  $8'' \times 8''$ ) for each ALMA continuum image, centred at the optical RA and Dec., to stack them based on median and average statistics (see Fig. C1). We reach rms values from 3.5 to 5.5  $\mu\text{Jy beam}^{-1}$  in these stacks. At  $z = 1.47$ , we detected three galaxies out of five, and both the average and median show a significant emission at the image centre. For the  $z = 2.23$  population, only one out of four targets is detected, hence the clear detection in the mean stack is clearly biased by the brightest galaxy. This significant detection disappears when we look at the median estimate. We also combined all non-detections together (mixing  $z = 1.47$  and 2.23 galaxies) in the right panels of Fig. C1.

To estimate the significance of the stacks of non-detected ALMA images, we use a peak to noise criterion. The peak values are obtained from a 2-D Gaussian profile using a fixed centre and fixed width (FWHM) at  $1''$  (assuming point like detections). We measure the peak flux densities of 12.7  $\mu\text{Jy beam}^{-1}$  for the mean stacked image, and 13.1  $\mu\text{Jy beam}^{-1}$  for the median one. Comparing these values with the background noise, these peaks are only at  $\sim 3\sigma$  significance. We consider these stacks as non-detections.

Based on the rest-frame frequencies for these stack measurements, and considering the Rayleigh–Jeans tail (at  $850\mu\text{m}$  flux; Equ. 16 in Scoville et al. 2016), we can derive global ISM mass limits for our SHiZELS targets. Considering median stacks, and  $5\sigma$  upper limits for the  $z = 2.23$  population, we derive ISM masses of  $\log(M_{\text{ISM}}/M_{\odot}) = 9.5$  at  $z = 1.47$  and  $\log(M_{\text{ISM}}/M_{\odot}) < 9.2$  at  $z = 2.23$ .

## APPENDIX D: SERENDIPITOUS ALMA DETECTION

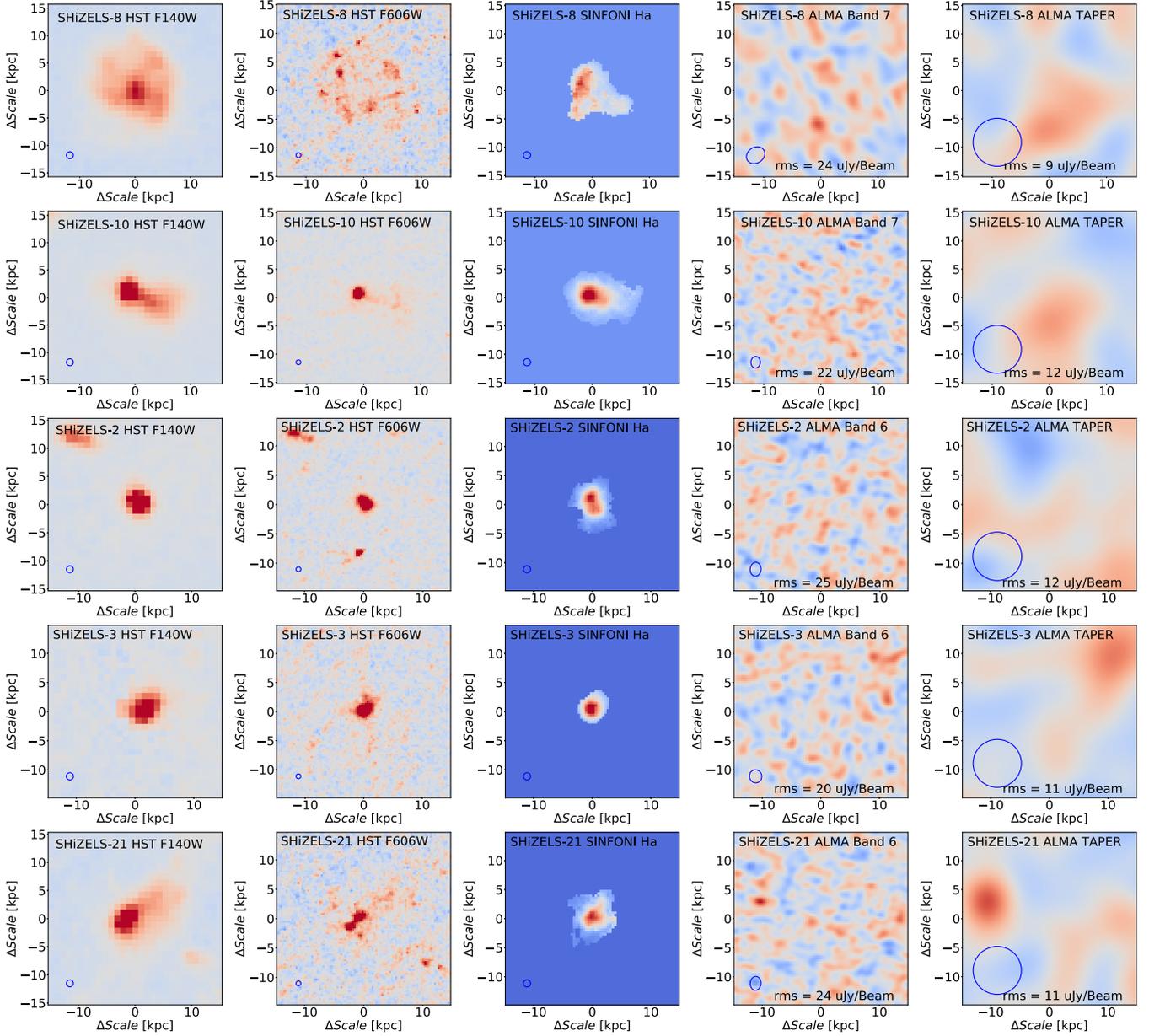
In this appendix, we briefly characterise the properties of the serendipitous galaxies SHiZELS7-ID2 and SHiZELS10-ID2. Both sources have been identified in the Multi-wavelength Photometric Catalog of the *Spitzer* Large Area Survey with Hyper-Suprime-Cam (SPLASH) in the Subaru XMM-Newton Deep Field (SXDF) (Mehta et al. 2018). SHiZELS10-ID2 presents a spectroscopic redshift at  $z_{\text{spec}} = 1.126$  while SHiZELS7-ID2 only has a photometric redshift estimate at  $z_{\text{phot}} = 2.03$ . In Fig. D1 we show the ALMA continuum  $355\mu\text{m}$  rest-frame contours on top of a fake colour optical image. The ALMA flux densities are considered for a SED fitting approach including  $U, G, R, I, Z, Y, J, H, Ks$  photometric bands taken from Mehta et al. (2018). We derive the stellar mass by MAGPHYS and the properties of these two targets are presented in Table D1 and Fig. D1. Considering the ALMA flux, our stellar masses are consistent with the previous results within 0.5 dex. More properties of these two targets can be found in Mehta et al. (2018).



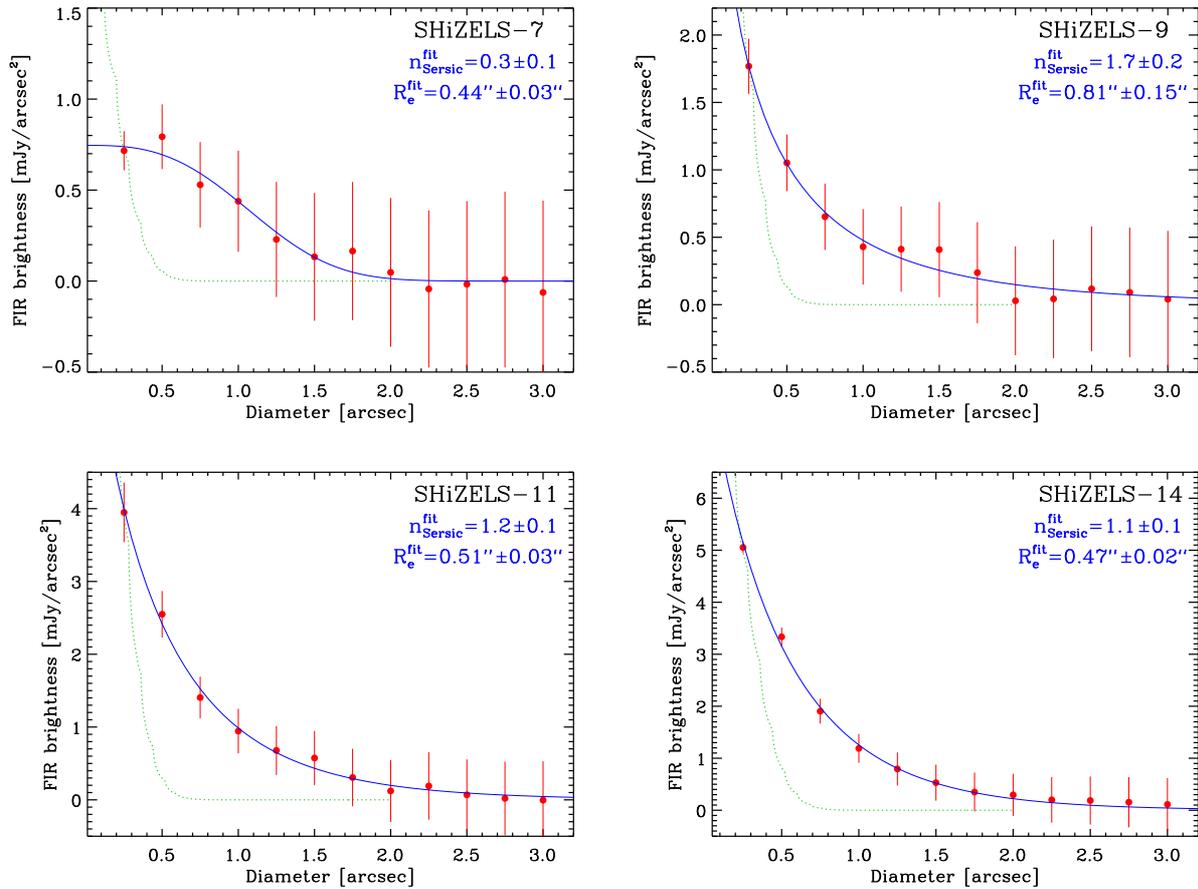
**Figure A1.** Postage stamp ( $30 \text{ kpc} \times 30 \text{ kpc}$ ) images for the SHiZELS galaxies presented in this work. From left to right images: *HST* F140W, *HST* F606W, VLT/SINFONI  $H\alpha$  moment-0, Briggs weighting ( $\text{robust}=0.5$ ) ALMA continuum at  $\sim 355 \mu\text{m}$  rest-frame, and tapered ALMA image (synthesised beam of  $\sim 1''$ ) to highlight extended emission. The image shows the complexity of the different phases of the ISM in these high-redshift galaxies

**Table D1.** Properties of the serendipitous ALMA detections found in the field of view of the SHiZELS galaxies presented in this work.

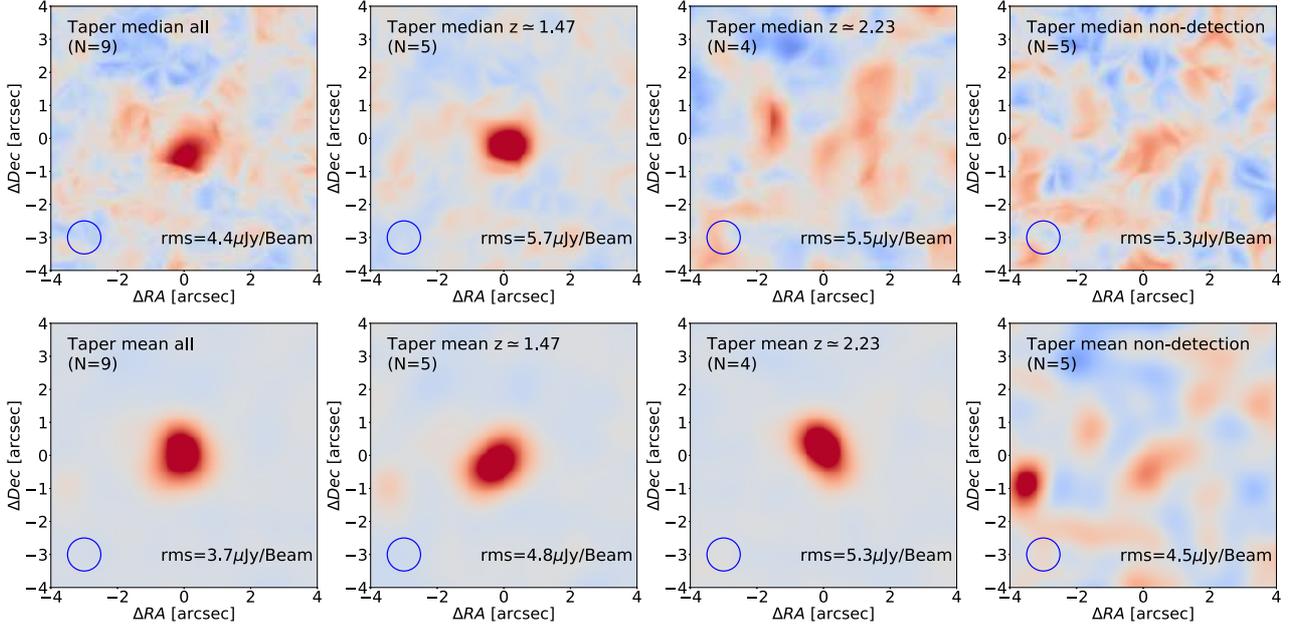
ID	ID From Mehta et al. (2018)	RA (J2000)	Dec (J2000)	Redshift	ALMA flux (mJy) 873 $\mu\text{m}$	$\log(M_*/M_\odot)$ By MAGPHYS	$\log(M_*/M_\odot)$ From Mehta et al. (2018)
SHiZELS10-ID2	954698	02:17:39.261	-4:44:42.33	$z_{\text{spec}} = 1.126$	$0.20 \pm 0.03$	$10.4 \pm 0.1$	10.2
SHiZELS7-ID2	874393	02:16:59.969	-5:01:53.49	$z_{\text{phot}} = 2.033 \pm 0.045$	$0.32 \pm 0.03$	$10.3 \pm 0.1$	10.6

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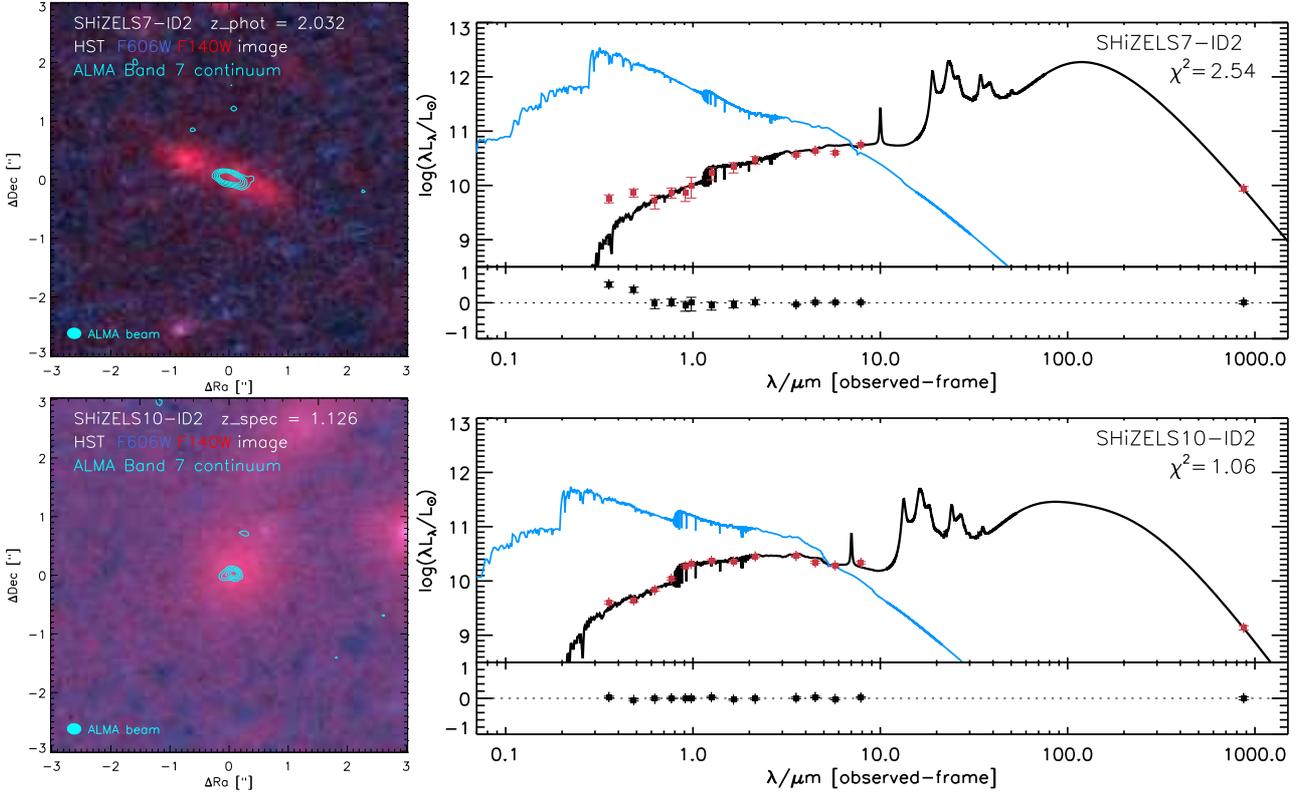
**Figure A2.** Same as Fig. A1, but for the undetected ALMA continuum sources.



**Figure B1.** The rest-frame  $355 \mu\text{m}$  surface brightness profile of the ALMA-detected galaxies. The interval of each aperture annuli is  $0''.25$  with the aperture centered at the peak pixel of the ALMA-tapered images. The green dotted line shows the surface brightness profile of the ALMA synthesised beam. We can see that continuum emission extends up to typically  $\sim 2''$ , i.e. about 16 kpc in diameter. The blue lines are the Sersic function fitting results. We denote the fitting results in each panel.

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**Figure C1.** Postage stamps ( $8'' \times 8''$ ) for the stacked continuum tapered (FWHM =  $1''$ ) ALMA images. Upper and lower panels are the median and average stacks. From left to right are the stacks for different populations: all sources, only  $z = 1.47$ , only  $z = 2.23$ , all non-detected images (mixing galaxies at  $z = 1.47$  at  $z = 2.23$ ).



**Figure D1.** Left panels: Fake color image (blue for F606W and red for F140W) with ALMA contour (3, 4, 5, 6, 7  $\times$  rms). Right panels: Redshifts and SEDs taken from [Mehta et al. \(2018\)](#) as well as our new ALMA observations are shown in red dots. We fit the SED using MAGPHYS ([da Cunha et al. 2008](#)). The initial stellar spectra are shown in blue lines, and the model SED after considering the dust extinction are shown in black lines.

1  
2  
3 Reviewer's Comments:

4 Reviewer: 1  
5

6 Comments to the Author

7 This is a concise and very well written paper, presenting a multi-  
8 faceted study of star-forming galaxies during the cosmic noon, taking  
9 advantage of high-resolution and state of the art observations that go  
10 beyond the usual characterization of the global properties of high-z  
11 galaxies or of luminous star-bursting systems. While the results are  
12 not unique among the literature and the reached conclusions (in my  
13 view) do not offer any pioneering insights, such efforts that combine  
14 multi-wavelength, high-resolution observations that trace, at the same  
15 physical scale, the stellar and the ISM components of high-z galaxies,  
16 are still scarce and very valuable for making progress in the field.  
17 The data, the methods and the analysis are clearly presented and the  
18 authors have achieved a good balanced by moving some technical (yet  
19 interesting for several readers) aspects in the Appendix, greatly  
20 enhancing the clarity and the overall flow of the paper. I only have a  
21 few minor comments/suggestions that the authors might want to take into  
22 consideration before I can recommending the manuscript for publication.  
23

24 We thank the referee for the helpful comments and suggestions. We  
25 respond to all the comments below.  
26

27 Comments:

28  
29 1. Data. What is the achieved rms of the high-resolution ALMA map? If I  
30 understand correctly the authors are only providing the rms/beam for  
31 the 1'' map. Adopting a conversion of the ALMA rms to  $M_{\text{dust}}$ ,  $M_{\text{gas}}$ , LIR  
32 or SFR, the authors can also provide the limiting dust ( LIR, SFR etc)  
33 surface density, of their observations.  
34

35 We have included the rms values in Figs. A1 & A2. ALMA continuum  
36 detections and upper limits, values taken from tapered maps, have been  
37 taken into account in the MAGPHYS SED fitting approach.  
38

39 2. Section 2.3. The comparison of the  $\text{SFR}_{\text{FIR}}$  and  $\text{SFR}_{\text{Ha\_corr}}$  makes  
40 sense only if the dust uncorrected  $\text{SFR}_{\text{Ha}}$  is only a small fraction of  
41 the  $\text{SFR}_{\text{Ha\_corr}}$ , since the  $\text{SFR}_{\text{FIR}}$  only traces the obscured, not the  
42 total SFR. I guess that this is the case for the current sample, but  
43 this should be included in the text, since in general the  $\text{SFR}_{\text{FIR}}$   
44 should be compared to  $\text{SFR}_{\text{Ha\_corr}} - \text{SFR}_{\text{Ha}}$ . Also, the authors state  
45 that the two SFR tracers are consistent within  $1\sigma$ . However, in Figure2  
46 three out of the four ALMA detected sources have  $\text{SFR}_{\text{FIR}}$  and  
47  $\text{SFR}_{\text{Ha\_corr}}$  values that deviate from the 1-1 relation by more than  $1\sigma$ .  
48  
49

50 The  $\text{SFR}_{\text{FIR}}$  and  $\text{SFR}_{\text{Ha\_corr}}$  trace the obscured and obscuration-  
51 corrected SFR, respectively, and it is true that the linear  
52 relationship (Kennicutt 98, Calzetti 2020) between both quantities hold  
53 for relatively low  $H\alpha$  extinction levels. Note, however, that the  
54 SMGs from the literature shown in Fig3 are bright submm-selected  
55 galaxies, for which the SFR is almost entirely dust-enshrouded (as in a  
56 calorimetry model, actually tracing the total SFR). For these dusty  
57 SMGs, there is no much sense in direct estimates of  $\text{SFR}_{\text{Ha\_corr}}$  because  
58 For these dusty SMGs, there is not much sense in direct estimates of  
59  $\text{SFR}_{\text{Ha\_corr}}$  from the little H-alpha that escapes the galaxy because H-  
60 alpha is not tracing the bulk of the star-formation. On the other hand,

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3 unlike the literature SMGs, our sample is H $\alpha$ -selected, and  
4 therefore a significant fraction of the total SFR is traced by H $\alpha$ .  
5 We have added a paragraph in Section 2.4 clarifying this point.  
6

7 A proper estimation of extinction (without Balmer decrements) is not  
8 trivial, nevertheless we estimate  $A_{\text{H}\alpha} \sim \lambda_{\text{mag}}$  values as based on  
9 Mstar estimates (following Garn & Best 2010). This is discussed in a  
10 following paper (Cochrane et al submitted), where we were able to  
11 describe the spatially resolved H $\alpha$  extinction in one galaxy.  
12

13 Please note that in Fig3, the filled red points correspond to the  
14 detected galaxies, while the unfilled remained undetected (upper  
15 limits). Only one detected galaxy is barely outside the  $1\sigma$   
16 threshold.  
17  
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19

20 3. FIR modeling. As the authors correctly state, with a single data  
21 point in the R-J tail ( $\lambda_{\text{rest}} > 300\mu\text{m}$ ) the LIR estimate is  
22 highly uncertain if not problematic. While the average  $T_d$  and its  
23 scatter adopted by authors could in principle be used when dealing with  
24 large samples, it is dangerous when applied to individual sources. In  
25 fact, if anything else, based on the current data the  $M_d$  estimate (that  
26 scales linearly to  $T_d$  for an optically thin FIR emission) is more  
27 secure than the LIR estimate presented by the authors, for which the  
28 uncertainty could reach  $1\text{dex}$  in absence of a data point close to the  
29 peak of the IR SED.  
30

31 Thanks for pointing this out. We agree that with only one FIR data  
32 point, it would be very uncertain to derive the IR luminosity and the  
33 dust mass.  
34

35 We agree that more rest-frame FIR data is needed to have a better  
36 constrain about the dust properties, especially near the peak of the  
37 far-IR emission.  
38

39 In this work, we estimate the dust mass using three different methods.  
40 We show these results in Table3, allowing the reader to use whichever  
41 they prefers. We find a relatively good agreement between them all.  
42  
43

44 4. ALMA detection vs  $M^*$  and  $Z$ . Although the trend between ALMA flux  
45 density and  $M^*$  and  $Z$  is indicated in Figure 3, I am wondering if the  
46 underlying reason for the non-detection is linked to the SFR ( $M_{\text{gas}}$ )  
47 and/or the surface brightness of the sources, given that galaxies with  
48 lower  $M^*$  tend to have lower SFRs (or  $M_{\text{gas}}$ ) based on the  $M^*$ -SFR main  
49 sequence relation. What is the position of the sources with respect to  
50 the main sequence at their corresponding redshift? Do the authors find  
51 a similar detection trend with sSFR?  
52

53 Following Gillman et al. (2019, Fig. 1), our data show that the ALMA  
54 detected targets (filled circles) have higher stellar mass, and lower  
55 sSFR:  
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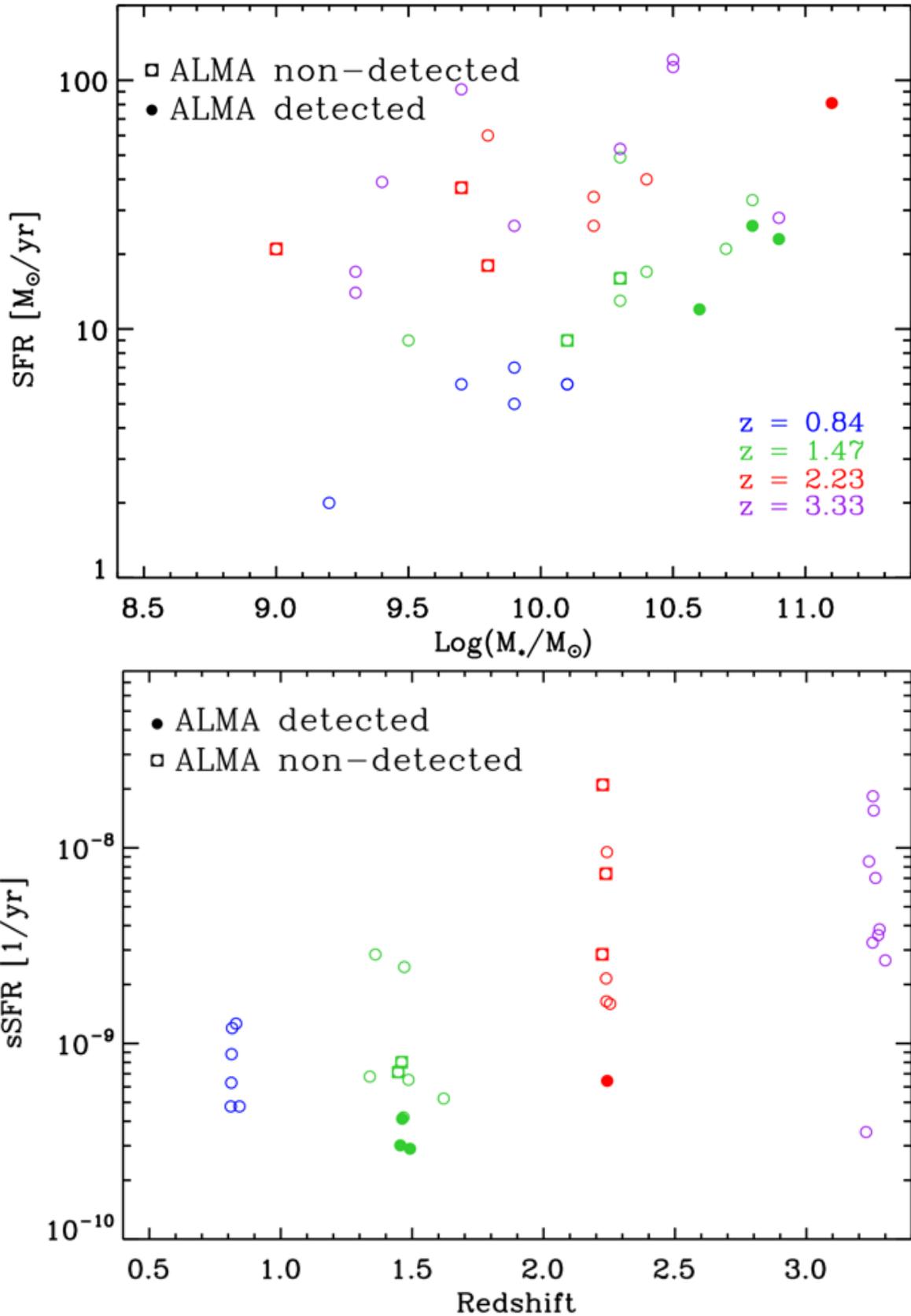


Fig. 1. upper panel: stellar mass v.s. SFR diagram. Lower panel: redshift v.s. sSFR diagram. We show all the SHiZELS sample in Gillman+ 2019 by open circles, and the ALMA observed targets in box. The open circles were in the SINFONI sample but not targeted with ALMA. The ALMA detected targets are shown by filled circles. SFR is estimated by Ha

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3 Based on the H-alpha emission, the targeted galaxies have SFRs representative of  
4 the whole HiZELS samples. At  $z=2.23$ , only the most highly star-forming galaxy is  
5 detected by ALMA, and at  $z=1.47$ , the three detected galaxies are three of the most  
6 H-alpha luminous and most massive.

7  
8 Since the stellar mass range of our sample is about  $10^9$ - $10^{11}$  Msun,  
9 while the SFR is in the range of 10 to 100 Msun/yr (within one order of  
10 magnitude), the ALMA detected galaxies have typically lower sSFR (a  
11 long standing star-forming mode). No clear trend is identified for this  
12 sample in terms of sSFR. We mention this point in Sec. 2.5.

13  
14 We thought that surface brightness may be linked to the non-detection,  
15 nevertheless since four out of five ALMA non-detected targets have  
16 compact F606W morphology, we conclude that (most probably) there is no  
17 compact dust content at the center, but low dust surface density.

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20 5. Comparison samples. Caution should be exercised when comparing the  
21 optical and FIR properties of different samples in the literature as  
22 different SED fitting methods could introduce systematics. Have the  
23 authors homogenized the properties of the local and high- $z$  literature  
24 samples to those derived based on their approach? Based on the  
25 evolution of the gas fraction of star forming galaxies reported in a  
26 plethora of literature studies I find it surprising that their sample  
27 has similar  $M_d/M^*$  (and therefore similar  $M_{\text{gas}}/M^*$ ) to that of local  
28 star-forming galaxies.

29  
30 The  $M_d/M^*$  of the local galaxy sample comes from Dustpedia. Since the  
31 stellar and dust mass of Dustpedia sample are derived by CIGALE, with a  
32 Salpeter IMF, while our results are derived by MAGPHYS with a Chabrier  
33 IMF, to have a consistent comparison, we re-fit the Dustpedia  
34 photometry with MAGPHYS and show the results in the new version.

35  
36 6. Given the spatial offset between the FIR, stellar and Ha emission,  
37 how do the authors justify the use of MAGPHYS (or of any other energy  
38 balance treatment) for the derivation of the FIR properties?

39  
40 The spatial offset shows that the energy balance assumption is not  
41 trivial for normal high- $z$  galaxies. Indeed, this opens a very  
42 interesting discussion about the reliability of such energy-balance  
43 analyses. This is mentioned in the text and it is being addressed in an  
44 upcoming following paper (Cochrane et al submitted).

45  
46 7. Sizes: Are the literature IR sizes presented in the various plots  
47 measured at the same/similar rest-frame wavelength? Depending on the  
48 rest-frame wavelength some sizes could correspond to star formation and  
49 some other to dust/gas distribution. The relevant footnote for the  
50 GOALS sample regarding this aspect should be moved to the main text  
51 along with a similar description for the high- $z$  literature samples.  
52 Moving to the stellar size it is worth mentioning if their sources  
53 follow the mass-(optical) size relation of the star forming galaxies at  
54 their corresponding redshift.

55  
56 The IR sizes we selected from the literature were mainly observed by  
57 ALMA 870um for the  $2 < z < 3$  sample (Tadaki+2017, Rujopakarn+2016,  
58 Chen+2020), corresponding to about rest-frame 250um. On the other hand  
59 those measured from the NOEMA data at 1.1mm (Nelson+2019) for the  $z =$   
60 1.25 targets, is about rest frame 500 um. The local galaxy samples were  
measured by Herschel at 160um or 350um, which are similar to the same

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3 rest-frame wavelength of our ALMA measurements. Based on the results  
4 given by Mosenkov+ 2019, the rest-frame 250, 500  $\mu\text{m}$  to the 350  $\mu\text{m}$  size  
5 ratio is about 1.2. Therefore, although the IR size comparison sample  
6 in this work were not measured in the same wavelength, this should not  
7 cause substantial bias. We moved the footnote to the context in the new  
8 version.

9  
10 For the stellar mass size, our F140W size is consistent with the half-  
11 light radius that is predicted from the mass-size relation (Suess+  
12 2019) within  $1\sigma$ . We add this result in the new version.

13  
14 8. Do the authors have in hand resolved metallicity maps, like those  
15 presented in Swinbank 2012? If so it would be really interesting to see  
16 the metallicity maps ( e.g. for SHiZELS-11) overlaid with the dust  
17 continuum contours.

18  
19 Although seems relatively trivial, the global astrometric uncertainties  
20 are those that dominate the conclusions from such analyses. We explore  
21 such analysis in an upcoming paper, Cochrane et al. (submitted).

22  
23 9. I guess that there is dynamical information available for the  
24 sources. It would be nice (but not necessary) to explore if the derived  
25  $M_{\text{gas}}$  estimates are consistent with the dynamical mass of the systems.  
26 In other words do you find  $M_{\text{dyn}}(\text{Re}) > M_{\text{gas}}(\text{Re}) + M_{\text{star}}(\text{Re})$ ? What is the  
27 corresponding dark matter fraction?

28  
29 We did not work on the dynamical mass in this work because the gas mass  
30 which are derived from one RJ point are significantly uncertain.  
31 Dynamical properties of the SHiZELS sample using direct CO emission was  
32 studied in Molina+2019. By modeling the mass distribution, we  
33 constrained the dark matter fraction as well as the CO-to-H<sub>2</sub> conversion  
34 factor. We are following up Molina's work using a larger sample of CO-  
35 detected galaxies.

36  
37 10. A relevant publication that the authors might want to consider is  
38 Rujopakarn et al. 2019 that presents (very) high resolution  
39 observations of the dust continuum of normal star forming galaxies at  
40  $z \sim 3$ .

41  
42 We cite this work in the newer version. Thanks for informing us.  
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