

Statistical correlation analysis of field-aligned currents measured by Swarm

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Key Points:

- For the first time, we infer the orientations of the current sheets of FACs.
- This study clearly reflects two different domains of FACs, small-scale which are time variable and large-scale which are rather stationary.
- We may find a way to obtain the information of Pi1B waves using low-earth orbit (LEO) satellites.

Abstract

We investigate the statistical, dual-spacecraft correlations of field-aligned current (FAC) signatures between two Swarm spacecraft. For the first time, we infer the orientations of the current sheets of FACs by directly using the maximum correlations obtained from sliding data segments. The current sheet orientations are shown to broadly follow the mean shape of the auroral boundary for the lower latitudes and that these are most well-ordered on the dusk side. Orientations at higher latitudes are less well ordered. In addition, the maximum correlation coefficients are explored as a function of MLT and in terms of either the time shift (δt), or the shift in longitude (δlon) between Swarm A and C for various filtering levels and choice of auroral region. We find that the low-latitude FACs show the strongest correlations for a broad range of MLT centred on dawn and dusk, with a higher correlation coefficient on the dusk-side and lower correlations near noon and midnight. The positions of maximum correlation are sensitive to the level of low-pass filter applied to the data, implying temporal influence in the data. This study clearly reflects the two different domains of FACs: small-scale (some tens of km), which are time variable and large-scale (>50 km) which are rather stationary. The methodology is deliberately chosen to highlight the locations of small-scale influences which are generally variable in both time and space. In addition, we may also have fortuitously found a way to obtain the information of Pi1B waves (pulsations with periods between 1-40 seconds).

1 Introduction

The Earth's field-aligned currents (FACs) are the dominant process by which energy and momentum are transported between the magnetosphere and the ionosphere-thermosphere system [e.g. *Foster et al.*, 1983; *Lu et al.*, 1998, *Yu et al.*, 2010], and therefore FACs are fundamentally important for the understanding of magnetosphere-ionosphere coupling. The upward FAC is responsible, at least in part, for the heating of the ionospheric electrons, although it is less clear whether the downward FAC cools the ionosphere [*Pitout et al.* 2015; *Wing et al.* 2015].

Both large- and small-scale FACs have been observed in the auroral zone extending over several degrees of magnetic latitude. Large-scale FACs (the Birkeland current system), with perturbations on spatial scales larger than 50 km at low-Earth orbit (LEO) satellite altitudes, have been described by *Iijima and Potemra* [1976a] in terms of 'Region 1' (R1) and 'Region 2' (R2) systems, which couple the external magnetospheric currents to the high-latitude ionosphere and the inner magnetosphere to the auroral ionosphere. *Iijima and Potemra* [1978] later found that the field-aligned currents flow into Region 1 on the dawn-side and away from Region 1 (out of the ionosphere) on the dusk-side. They also found that the current flow in Region 2 is reversed with respect to Region 1 at any given local time except in the Harang discontinuity region, ~20:00-24:00 MLT [*Harang*, 1946], where the flow patterns are more complicated. There is evidence that the large-scale FACs are generated by the 'long-term' interaction of the solar wind with the magnetosphere (for recent work, see *Wing and Johnson*, [2015], which showed upwards R1 currents can be driven by solar wind velocity shears at the magnetopause), although these current sheets can often have complicated spatial and temporal variations (here 'sheet' refers to the discussion of the azimuthal extent of R1/R2 FACs in this paper). Small-scale FACs are usually characterized by quasi-equal, parallel sheets of current into and out of the ionosphere with latitudinal thicknesses of tens of km at low Earth orbit (LEO) altitudes and with typical timescales of order 10 seconds or less. These small-scale FACs are associated with 'short-lived' plasma processes within the magnetosphere such as discrete auroral arcs [*Anderson and Vondrak*, 1975], field-line resonances [*Rankin et al.*, 1999; *Pitout et al.*, 2003; *Waters and Sciffer*, 2008], bursty bulk flows in the plasma sheet [*Merkin et al.*, 2013; *Yu et al.*, 2017] and associated Pi2 [*Cao et al.*, 2008, 2010], as well as Pi1 waves which will be discussed in detail in section 3.3.

Nevertheless, separation of the temporal and spatial nature of both small- and large-scale FACs has been notoriously difficult [for example see *Lühr et al.*, 2015; *Stasiewicz et al.* 2000] since these currents are both highly dynamic and vary in size, while single spacecraft estimates generally require assumptions of either geometry (such as infinite sheets, as adopted by, for example, *Anderson and Vondrak* [1975]; *Marshall et al.* [1991]) or some degree of time stationarity (to apply dB/dt to a spatial estimate, where multi-spacecraft estimates are unavailable [Dunlop et al. 1988], for comparison). Despite this problem, since the first identification of FACs [Zmuda et al., 1966; 1967; *Iijima and Potemra*, 1976a], many previous, typically statistical, studies have been performed, using single and multi- spacecraft methods (see Higuchi and Ohtani 2000), or indirect observations, to probe their global natures [e.g. *Anderson et al.*, 2000; *Gjerloev et al.*, 2011; *Dunlop et al.*, 2015a,]. An investigation of the characteristics of FACs which are restricted in both their spatial and temporal variations between multiple spacecraft positions has also recently been carried out by *Forsyth et al.*, [2017] (through the development of a rigorous test of purely static, 1-D normal current sheets) and has been applied recently in a study by *McGranaghan et al.*, [2017].

The alignment of current sheets of large-scale FACs is generally along the boundary of the auroral oval, but can be noticeably distorted during very disturbed periods [*Iijima and Potemra*, 1978]. Nevertheless, it has been argued that the basic pattern may often be maintained [*Gjerloev and Hoffman*, 2014], although the intensity of currents varies from event to event. Here, we have used the recently acquired Swarm multi-spacecraft data set to investigate the MLT dependence of the correlations between the two spacecraft field-aligned current sheets with a new method using statistical analysis of the inter-spacecraft maximum correlations between FAC signatures, which also shows directly the auroral alignments of the current sheets. The sensitivity of this analysis to the filtering of the data and both the time delay and longitudinal separation between the spacecraft are explored. The statistical work shows differences between large-scale FAC sheets which occur mainly in the dawn- and dusk- sectors and more localised current sheets possibly associated with the NBZ [as defined by *Iijima et al.*, 1984] and cusp currents [*Iijima and Potemra* 1976b], also referred to as Region 0 currents [*Bythrow, et al.*, 1988].

2 Methodology

The Swarm mission [*Friis-Christensen et al.* 2008] consists of three spacecraft (A, B and C) flying in phased, circular, low-Earth polar orbits since launch on 22 November 2013. The data set used here was mainly the FAC signals derived from the Swarm A and C observations during the final constellation phase (operations from 17th April 2014), where the two spacecraft had orbital periods of ~94 minutes, flying side-by-side at a mean high-latitude altitude of about 470 km, and sampling all local times in about 132 days. The third spacecraft Swarm B flies at a slightly higher orbit at ~531 km altitude, with a slightly different orbital period of ~95 minutes and drifts in MLT with respect to Swarm A and C, which remain close together throughout the time period studied here. The three Swarm spacecraft move through the auroral regions and across the polar cap as a result of their near polar orbits.

We use the official 1 Hz Level-2 OPER (Routine Operations of file class) FAC data taken from the Vector Fluxgate Magnetometer (VFM) [*Friis-Christensen et al.*, 2008, *Ritter et al.*, 2013, *Stolle et al.*, 2013] on Swarm. To minimize the non-linear variation of the magnetic field gradients, these data are processed by initial subtraction of the model “mean field” (the core, crustal and magnetospheric fields at the satellite altitude) to obtain the residual data [see *Dunlop et al.*, 2015b]. This generally results in a 5% uncertainty [*Ritter et al.*, 2013] in the estimates of the FACs as a result of non-physical errors. These estimates are provided as part of the standard Swarm level 2 data products (<https://earth.esa.int/web/guest/swarm/data-access>) with a cadence of 1 second and are obtained using a single-spacecraft method assuming that an infinite (1-D) current sheet approximation applies locally to each spacecraft (i.e. that the local structure sampled is approximately a planar sheet on temporal and spatial scales which are consistent with the 1 sec cadence). Here we also apply a low pass filter to this 1 second data with both 20 and 60 second cut-offs (removing higher frequency signals and maintaining

cadence) to obtain the large-scale (i.e. 150/450 km; corresponding to 20/60 second cut-offs respectively) FAC data. This filtering also serves to clarify the inter-comparison of spacecraft A and C data, which have a spatial separation of ~150 km (see discussion below). For our presentations in magnetic latitude and MLT we use APEX coordinates [Richmond, 1995] throughout. To probe the duration and extent of the FAC sheets, especially in different MLT regions, we statistically analyse the correlations between FACs observed by Swarm A and C during the period 17th April 2014 to 30th April 2016 when both spacecraft were flying side-by-side with apex longitude difference less than 3 degrees and a lagging time (from one spacecraft to the other) less than 20 seconds. Latitude is considered only through the auroral region (see below).

Figure 1 demonstrates the method we have adopted for data selection and the correlation analysis. Figures 1(a) and 1(b) show FACs along several Swarm orbit tracks within time period of 09:29-20:31 UT on 6th, June, 2014, projected onto APEX coordinates. We can see that similar FAC signals on Swarm A and C were seen for several hours, revealing corresponding current sheets distributed over some longitudes, but slowly changing in time and orbit track. Although the signals observed by Swarm A and C are very similar, differences are observed between them, even though the time delay for each spacecraft to arrive at the same APEX latitude varies from a few to about 14 seconds, and the difference in longitude (δlon) is ~1-3 degrees between Swarm A and C. The total time difference between the dual-spacecraft segments of maximum correlation (bold orbit segments in Figure 1c-f), indicates the time difference of arrival at the same current sheet, δt (see below).

To obtain the correlations of FACs observed by the two satellites we separate the regions between the modelled poleward and equatorward auroral boundaries (as defined by the method of Xiong and Lühr, [2014]) into two broadly equal intervals predominantly containing ‘R1’ and ‘R2’ signals (each containing approximately the same range of latitudes). The modelled poleward and equatorward auroral boundaries on 14:30 UT 6th June 2014 are shown in Figure 1 (a, b) by the magenta dashed curves. In Figure 1(c-f), however, for each orbit track the specific auroral boundaries are indicated by vertical dashed lines. The effective total time shift (δt) and the longitude difference (δlon) between the two spacecraft when the positions of the maximum correlations are found are denoted in each panel. Maximum correlations are obtained for 60 second sliding orbit segments of Swarm A and C within the R1 or R2 intervals. The segments with maximum correlation adopted finally are shown in Figure 1(c-f) by bold traces for each spacecraft (where the maximum correlation and longitude difference are indicated in blue text) with two different low pass filters, 20 seconds and 60 seconds. The filtering of the data defines the optimum data segments for that resolution and tests the temporal content, i.e. the degree of stationarity in the data signal is expected to decrease with decreasing scale size of activity.

We use 60-second length segment windows for the correlation to get rid of any influence from variable lengths on the computation of the maximum correlations. From tests using different segment lengths, we found that using longer segments can reduce the effectiveness in finding the max-correlation between two tracks when the ‘R1’ or ‘R2’ contain too many points within the segments, and it also can increase the likelihood of non-regular, shorter tracks occurring, introducing systematic errors in the maximum correlations. On the other hand, segments with too few points can decrease the confidence level of the correlation. After some experimentation, we selected a 60-second sliding window to maximise correlations and minimise the effects of systematic errors. When there is less than 72 points in an orbit track we remove the track.

The traces in Figure 1(c-f) show two different orbits for the correlation analysis, organized by apex latitude. In Figure 1c, the bold traces at higher latitude are close together and the traces generally fall on top of one another with only small-scale differences between Swarm A and C. Therefore, the large-scale current sheet appears to be well aligned in latitude within the broad region ‘R1’. This high degree of similarity, furthermore, also results in the maximum correlation fixing to a different orbit segment within ‘R1’ in the case of the 60seconds filtered trace (Figure 1d), i.e. implying that the particular segments used

are not critical for the application of different filter windows. In the lower pair we see that the current sheet is also well aligned at the Apex latitude for the ‘R1’ interval but not well aligned for the ‘R2’ interval (since Swarm A and C see similar profiles at different latitudes). Note that in this lower case the δlon is larger (~ 3 degrees), whereas the δt varies (~ 20 seconds for the 20s filtered trace and ~ 5 seconds for the 60s filtered trace), so that in the top pair the small differences are probably temporal whereas in the lower pair the differences are spatial. Each example defines a particular (δlon , δt) for the MLT of the orbit track. Building up the statistics allows us to study the correlation trends as a function of MLT and the differences in APEX longitude or time between Swarm A and C, and to define an approximate orientation of the large-scale FACs (see section 3).

3 Correlation Analysis

3.1 Current sheet orientation

We can see from Figure 1(c-f) that the 2 orbits chosen show distinct situations. The first shows a time delay of ~ 13 seconds between Swarm A and C, with small difference in longitude, and the FAC signals are seen at the same latitude. The second shows a different time delay in a different region (which changes for different filters), together with a larger difference in longitude. Figure 1 (e-f) shows a time delay in R1 of 12 seconds, and in R2 of 20 seconds (for 20s filtered data) and 5 seconds (for 60s- filtered data). The reason for different time delays (20 seconds in Figure 1e and 5 seconds in Figure 1f in R2) is that the maximum correlations appear at different latitudes when using different low pass filters, which select different temporal, and hence spatial, scales. If we draw a line between the centres (average positions) of the two orbit segments which give the maximum correlation between Swarm A and C, then this provides an estimate of the orientation of the current sheet. Inevitably there are influences arising from any temporal evolution of the current sheet between Swarm A and C (depending on the filtering used [Lühr *et al* 2015; Forsyth *et al.*, 2017]), such as influence arising from any propagation of the current sheet during the shifted time, δt . In addition, spatial structure on the scale of the spacecraft separation will also influence the estimate [Dunlop *et al.* 2016, 2018]. Nevertheless, the important point to note here is that for high levels of cross-correlation and hence for large-scale structures which do not significantly evolve on the scale of δt and $\delta long$, the estimates are more accurate, so that fluctuations and variance in the orientations highlight the presence of small-scale FACs. From statistics, we find that the 20s and 60s filtered data show very similar results. Thus, these current sheet orientations of R1 and R2 at 110 km altitude are drawn on polar maps in Figure 2a and 2b only for the 20 second filtered dataset.

It needs to be pointed out that, as mentioned earlier, we use the model estimated auroral boundaries from Xiong and Lühr [2014] here to split each auroral track into two regions, and indicate them as higher latitude regions (labelled loosely as ‘R1’ in this paper) and lower latitude regions (labelled approximately as ‘R2’ in this paper). Since we are using only these broad definitions of the intervals for simplicity, we expect that the ‘R1’ set actually contains other currents than purely R1, and that the ‘R2’ set contains some R1 currents in actuality, and indeed there are other currents around noon. Nevertheless, from a statistical perspective we expect that the main characteristics of the large-scale currents will dominate each region. The polarity of the currents in R1/R2 also does not affect our results since we are considering only the ordering of the current sheet orientations (and the pattern of correlations with MLT in section 3.2).

Figure 2a,b show inferred current sheet orientations (estimated by the method described above and as described in the caption) for each region using 20 seconds filtered data from 17th Apr. 2014 to 20th Aug. 2014, during which time period the Swarm A and C orbits have covered the Earth for nearly a full range of 24 hrs LT. Figure 2a, denoted by ‘R1’, shows the current sheet orientations found for the higher latitude regions. Figure 2b, denoted by ‘R2’, is for the lower latitude regions. The magnitude of the longitudes (0° , 90° , 180° and -90°) in Apex coordinates are denoted on each figure and the latitudes in the same coordinates are denoted at the line of 135° longitude. The sets of line segments shown in

each panel, representing the current sheet orientations, are drawn for those estimates at correlation values over 0.97. At this threshold the patterns are most clearly visible and show the distinction in the ordering in each region (see below). For lower threshold values of the correlation more vectors would be included and these contain more influence from temporal and spatial effects.

The implication of Figure 2b is that the large-scale current sheets in the lower latitude region broadly follow the oval on the dusk side and also are well consistent (although slightly less well ordered) with the oval pattern on the dawn-side. In the dawn-dusk regions, therefore, this ordering of the current sheets is very apparent and appears to be more significant on the dusk side oval. Near noon in particular, however, the current sheet orientations are more randomised and perhaps reflect the influence from other current systems in that region and/or smaller scale structures. It also appears that the alignment of the current sheets is better ordered as we move to the lowest latitudes in the distribution. Figure 2a shows that the orientation of the current sheets for the higher latitude region is less well ordered to the oval. This suggests that the character of the poleward current sheets is less dominated by large-scale structures and that this region contains more than one current system.

As a further check on the stability of these estimates, Figure 2c shows the comparison of these current sheet orientations to those implied for a 1-D current sheet inferred from maximum variance estimates (MVA, see *Sonnerup and Scheible*, [1998]) of the orientations in terms of the intersection angle (the difference between the orientations for each method). The data segments for MVA were taken from those used for the maximum correlations for each Swarm A position. The current sheet orientations obtained from maximum correlation are quite similar to the estimates using the MVA method. The plot shows that the intersection angles of the average current alignments derived from these two methods (MVA and maximum correlation) for Swarm A and C data from 17th Apr. 2014 to 30th Apr. 2016 and these are all less than 0.8 deg (and less than this in the auroral region). The intersection angles are lowest in the dawn side region. In fact, the difference between the two methods will arise naturally since the MVA measurement is centred on the Swarm A position whereas the maximum correlation result is an average centring on a position midway between A-C. We might expect that these will agree best in regions dominated by large-scale structure and this is indeed the case although there is an asymmetry in the extent of the agreement from dawn to dusk. The close agreement on the dawn side suggests that the orientations from both methods on the dawn side are more stable over a wider range of MLT and that the effect of the differing positions is less, perhaps due to the simpler shape of the oval on the dawn-side. On the other hand, the close average alignment of the maximum correlation orientations on the dusk side suggests that the large-scale ordering is most dominant there. Such asymmetric character from dawn to dusk is also seen in the correlation trends discussed below.

Figure 2d, denoted by 'R1&R2', shows the current sheet orientations for both the higher and lower latitude regions combined, superimposed on the average distribution of FACs for Swarm A from 17th Apr. 2014 to 30th Apr. 2016. In the underlying pattern of FACs we see the Region 2 and Region 1 up-down (in the field-aligned sense this corresponds to: negative-positive) currents as well as the Region 0 and NBZ regions near noon. This overlay plot shows that the current sheet orientations of 'R1&R2' can cover the whole oval region, as well as the Region 0 and NBZ region. The alignment of the current sheets reflects the large-scale features in the polar map of average FACs closely. Although further work is required to quantify the characteristics, the mean position separating of Region 1 and Region 2 can be seen. In addition, the cluster of differing orientations near noon corresponds to the average currents seen there. The intensity of the average current does not correlate with the alignment of the sheets in general.

Further details of the FAC current sheet orientations (in particular, separating the behaviour in terms of activity and other external drivers and exploring further the stability of the orientations for different correlation levels) will be discussed in future work. Here, we focus simply on the R1/R2 alignment in order to compare with the correlation trends described below.

3.2 Correlation trends

Using the methodology described in section 2 we analysed the data from 17th Apr. 2014 to 30th Apr. 2016, where the spacecraft pair A-C covered a close range of both cross-track (local time longitude, δlon) and along track (time differences between the spacecraft, δt) positions (see Figure 1). The range of δt and δlon are 0-0.3 minutes along track and 0-3 degree longitude (APEX coordinates) across track. These ranges allow us to separate the cross correlations between the spacecraft into both δt and δlon bins independently, so as to explore the MLT dependence of the correlations. We have explored these correlations for both filtered and unfiltered data to understand the effect of large- (>150 km) and small- (~7.5 km) scale structures, and their trends. The filtered data, as discussed earlier, allows consistent comparisons on the scales of the inter-spacecraft separation (i.e. ~150 km) and above. The lower choice of 20 secs matches the cadence used for the dual spacecraft FAC product in the Swarm level 2 data and therefore was used for the estimates of current sheet orientation in the previous section. It should be emphasised that it is not always possible to completely separate spatial and temporal behaviour and small-scale FACs in general depend both on space and time. Nevertheless, single spacecraft FAC estimates can still be valid locally at each spacecraft within certain criteria [Lühr *et al.* 1996] and the Swarm products are calculated at the higher smoothed cadence of 1 sec (what we term unfiltered data here). Although some types of behaviour are problematic, and the estimates can be quantitatively in error, variations on the spacecraft separation scale can be monitored through their effect on the correlation trends. We use the single spacecraft estimates in this sense here to reveal the locations and some characteristics of the smaller scale currents, through comparison of filtered and unfiltered signals.

Figures 3a and 3b show two 3×3 arrays of MLT trends separated by regions as described in section 2. For ‘R1&R2’ the analysis is performed across the whole interval, rather than for the separated ‘R1’ and ‘R2’ intervals. These intervals are chosen to capitalize predominantly Region 1 and 2 signals separately. We expect that the R2 interval should relate to actual region-2 FACs but also contain some actual R1 currents, even some other currents at noon, as shown by Figure 2b. Meanwhile the R1 interval will include many of the other high latitude current systems from time to time. It is instructive to consider this separation for both unfiltered and filtered data. Figure 3a shows the trends with respect to the time delay from the Swarm A and C spacecraft. The top panels of Figure 3a show the total correlations for the unfiltered magnetic residuals, i.e. those including signals from both small and large-scale structures. The panels headed ‘R1’ and ‘R2’ show clearly distinct trends in both MLT and δt , δlon , which actually are broadly maintained for each of the filtered datasets, consistent with the predominant nature of these regions. The R1&R2 panel is shown for context and represents the strongest effect of the signals seen in the whole auroral and some of the polar regions.

For the unfiltered, 1s resolution data, the R2 correlations remain relatively high for a broad range of MLT and are obviously lower for the range 9-15 MLT (i.e. around local noon). There is also a minor dip in the strength of the correlations from 0-3 MLT (i.e. at local midnight). This trend is maintained for most of the range of δt and is consistent with the expectation that R2 FACs are stable for a large range of MLT, centred on pre-dawn and post-dusk. The correlations on the dusk-side are higher, and extend for the maximum range of δt , suggesting a dawn-dusk asymmetry in the stability of R2 FACs. This is probably associated with the high correspondence between particle precipitation at dusk and R2 FACs (see Korth *et al.*, [2014]). By contrast, the R1 correlations peak during the ranges 15-21 MLT and 3-9 MLT, i.e. dusk and dawn, and are maintained for a smaller range of δt . Thus, the correlations are lower (less than 0.83) for a broad range of MLT around local midnight. These R1 correlations also peak at δt near 0.13 mins, or 8 seconds, which may indicate that Pi1B waves [e.g., Heacock 1967; Arnoldy *et al.*, 1998] probably can be revealed by LEO satellite by this method. This will be discussed in section 3.3). These trends are consistent with the expectation that R1 FACs will be more temporally unstable overall and there is some indication that at either side of noon the signatures are more stable: the lower correlation around noon is possibly a result of the presence of other FACs, such as the NBZ currents or Region 0 currents which also are called cusp currents. The fact that the higher correlations extend away from noon is consistent with the average dawn-dusk signature of R1 currents, while the minimum post-midnight may be consistent

with the presence of the diffuse aurora, which is most likely composed of field-aligned plasma sheet electrons scattered by the very low frequency (VLF) whistler-mode chorus waves [Wing *et al.*, 2013] and see the suggestions of Newell *et al.*, [2009], Korth *et al.*, [2014], or McGranaghan *et al.*, [2016].

The 20 secs and 60 secs filtered data shows similar trends, but a higher value of correlations, to those for R1 and R2 separately. This appears to suggest that the medium to large-scale FACs dominate the MLT trends, but other work has indicated this may not always be the case [Neubert and Christiansen, 2003; McGranaghan *et al.*, 2017]. We see, moreover, that the combined region R1&R2, for 20 secs filtering maintains the combined distribution, suggesting that it is the smaller scale currents which affect the loss of correlation in the unfiltered data. Broadly, the trends with MLT for R1 and R2 are more similar for the filtered data to each other, peaking away from both noon and midnight. This is consistent with the general pattern of large-scale FACs for both R1 and R2, which follow the well-known upward and downward pattern for a broad range of local times surrounding dawn and dusk (Iijima and Potemra 1978). However, the 20 secs and 60 secs filtered data for R1 shows some additional structure, i.e. the correlations sometimes (e.g., at 15-18 MLT for 20 secs and 18-21 MLT for 60 secs filtered data) increase instead of decrease as δt decrease at the lowest δt (0-0.05 min bins). This implies the trend modulated by the wave is defeated by the expected peak at low δt , which is an obvious trend for the unfiltered correlation in Figure 3c.

The trends in Figure 3b are shown for the δlon separation between Swarm A and C, which are different from the trends for δt . Firstly, we see that the R1 trends are highly localised to small δlon (0-0.5 degs) and rapidly fall off as δlon increases. The correlations of the two highest correlated, or steadier, FACs regions, 15-21 and 3-9 MLT for R1, fall from about 0.94 to 0.78, for the unfiltered data, and ~0.995 to ~0.975 for the filtered data. This suggests that the FAC profiles are very sensitive to shifts in longitude. This effect lessens significantly for the 20 secs and 60 secs data, as might be expected for larger scale FACs. This can be attributed to the lower applicability of the infinite current sheet approximation to the small-scale currents. This also suggests that the correlations seen in δt are dominated by the periods when A-C have a small difference in longitude.

R2 currents exhibit similar but weak trends. After examining the number of cases in each bin, we suggest the peaks around $\delta lon \sim 1$ degree are probably from the rare cases in the lowest of all valid bins (containing not less than 5 cases). The higher correlations and less sensitive to shifts in longitude in 'R2' region for the 20 secs data revealed by Figure 2b is consistent with the more aligned current sheets.

We have also investigated the combined correlations between spacecraft pair A-B and B-C. The range of δt and δlon are much higher than those of spacecraft pair A-C, because of the distinct orbit of 'B'. Figure 3c is similar to Figure 3a but shows δt in the range of 0-44 minutes along track with data limited by $\delta lon < 10$ degrees. Figure 3d is similar to Figure 3b but shows a δlon of 0-15 degree across track with data limited by $\delta t < 30$ min. The MLT dependence can still be found, with correlations higher at the dawn and dusk side for both R1 and R2, but is a little more night-ward for R2. The unfiltered data shows that the correlations decrease as δt increases, but this trend is not so clear for the filtered data. Figure 3d reveals only the MLT dependence of both filtered and unfiltered data but no obvious trend by different δlon . This may be because of the influence from different altitudes of each spacecraft, i.e. 470km, versus 531km, which compete with the δt or δlon sensitivity. However, it can probably confirm that the large-scale R1/2 FACs are relatively stable in tens of minutes, otherwise the lower two panels of Figure 3c should exhibit some descending trend with δt .

The δt versus δlon dependence of correlations taken at different MLT regions (not shown here) demonstrates that the peak in δt seen for 'R1', corresponding to temporal variations of order ~8 seconds (as outlined above), doesn't come from a specific low δlon by chance. This probably suggests that the modulation of the currents by Alfvén waves is notable in the 'R1' region [Ma *et al.*, 1995; Stasiewicz *et*

al., 2000; *Liu et al.*, 2009]. In fact, we see that for the unfiltered combined interval R1&R2, nearly only the overlapping region between 18-21 MLT remains high (more than 0.85), and is also centred around $\delta t \sim 0.13$ mins (8 seconds). We will discuss this phenomenon in detail in section 3.3.

3.3 Pi1B waves

The occurrence of ground based Pi1B is well-documented [e.g., *Heacock* 1967; *Arnoldy et al.*, 1998]. Here ‘B’ is the abbreviation for ‘Burst’. Although there is evidence of both an ionospheric origin, there is still no agreement on the origin of the Pi1B waves. Heacock suggested that ground Pi1B waves were not generated in space because of the lack of frequency dispersion in the ground events. Subsequently, several studies have shown the association of Pi1B waves with different types of ionospheric activity [*Arnoldy et al.*, 1998 and papers therein] indicating that the waves probably result from ionospheric current fluctuations. Using magnetic field data, *Arnoldy et al.*, [1998] showed that Pi1B waves observed by the geosynchronous GOES (Geostationary Operational Environmental Satellites) satellites were nearly simultaneously observed on the ground and appeared to be initiated by the dipolarization process of the night-side tail magnetic field at the onset of substorms. *Arnoldy et al.* further commented that with induction antennas sampling up to 10 Hz, there is indeed evidence of dispersion in the higher frequency Pi1B waves.

Other work has also suggested that Pi1B waves are associated with substorms, as well as FACs. *Lessard et al.* [2006] suggested that they were excited by reconnection or some other processes, and were compressional in nature, at least at geosynchronous orbit, implying either fast or slow mode. It should be noted that slow mode waves would be quickly damped so do not propagate to the ionosphere. However, the fast mode waves can propagate isotropically, cutting across the magnetic field obliquely in the vicinity of the GOES satellites. They noticed that at FAST (Fast Auroral Snapshot Explorer) altitudes, the waves are of shear-mode, so must have undergone mode conversion in the region between GOES 9 and FAST. They suggested it was possible that as the waves approached the higher latitude regions of the magnetosphere, they gradually became increasingly parallel to the background field, where they may take on the properties of a shear wave (a guided wave) and follow the field lines to the ionosphere. Other models have also been suggested [*Pilipenko et al.*, 2008; *Lessard et al.*, 2011] to interpret how propagating compressional fast magnetosonic (FMS) modes transform into running Alfvén waves.

Since Pi1B waves may change mode as they propagate, and are not well studied, it is important to investigate them at different altitudes. Under normal conditions, the curl-free part of the ionospheric, horizontal current, as well as FACs, can’t directly produce any magnetic field disturbances below the ionosphere [*Fukushima*, 1976] and thus these currents are usually magnetically invisible on the ground, but can be detected by satellites above the ionosphere. Therefore, although it is known that in general ground stations usually observe Pi1B waves between 2100 and 0200 MLT [*Arnoldy et al.*, 1998, *Posch et al.*, 2007], ground magnetometer data alone cannot define the ionospheric phenomena. Nevertheless, it is hard to observe ultra-low frequency (ULF) waves in LEO satellites since a satellite generally moves fast at this altitude so that distinguishing the temporal from the spatial variations is challenging; particularly using single spacecraft measurements. The maximum correlation method, introduced here, however, can give us the spatial and temporal variation of the FACs, so potentially providing information on Pi1B properties, which are associated with the upward [*Bösinger et al.*, 1981] and downward [*Milling et al.*, 2008] FACs.

Figure 3 reveals apparent evidence of a correlation maximum around 8 second, which is about 0.13 Hz. In order to clarify this, Figure 4a represents the number of cases found in each bin of Figure 3a. Even though the data points are not equally distributed in each bin, Figure 4a only shows some overall trend for the specific orbits of Swarm A and C. The peaks around 8 seconds are not consistent with the Figure 4a, so do not arise from the basis time phasing between Swarm A and C (shifted by 8 seconds for orbit cross-over). Figure 4b is another way to express the behaviour shown in Figure 3a, where the correlations are no longer shown by colour, but by the y axis on the left of each panel. Different colours represent

different MLTs and the δt difference is now shown in the x axis. Through Figure 4b, we can see the correlation peak around 0.13 minute (about 8 second), nearly at all MLT, and most pronounced for the 15-21 MLT regions. *Bösinger and Wedeken*, [1987] have mentioned that Pi1B wave enhancement at 0.08–0.25 Hz was frequently observed at each of their 6 stations in both horizontal components. We note here that this is just around 4-12 second in the time domain with the centre at just 8 second, as revealed by our correlation method. In contrast to the local range in longitude of the 6 stations mentioned in their paper, the Pi1 band phenomenon revealed by the maximum correlation method here exhibits a global property and the wave character is most obvious in R1, which is considered to map mostly to the boundary plasma sheet (BPS). Additionally, the 60s filtered data of R1&R2 (the lower left panel of Figure 3a) shows another peak at around 0.26 minutes (~16 second), which can be treated as a secondary harmonic of the 8 second wave. It is reasonable that the harmonic wave appears when the data are low-pass filtered.

Although the Pi1B waves observed on the ground have maximum amplitude when they lie underneath active auroral forms (*Haldoupis et al.*, 1982; *Bösinger and Wedeken* 1987; *Arnold et al.*, 1998; *Danielides et al.*, 2001; *Milling et al.*, 2008), thereby suggesting they were locally generated [*Posch et al.*, 2007]. a number of studies have mentioned that Pi1B waves can extend in latitude and longitude/MLT (for example: 12° in latitude and 20° in longitude, as mentioned by *Arnoldy et al.*, [1998]; 7° of magnetic latitude and 4 h of MLT, as mentioned by *Posch et al.*, [2007]; 5° in latitude and less than 30° in longitude, as mentioned by *Parkhomov and Rakhmatulin*, [1975]) For these extended distributions, only the brightest auroral onsets can be associated with Pi1B observations at more than 5° in latitude and 2 h in MLT distance. Their onsets have been seen to occur earlier at the auroral zone latitude at magnetic midnight. The horizontal ducting of wave power has been put forward, as well as the westward delay consistent with the Pi1B and initiated by the westward auroral surge were discussed by *Arnoldy et al.*, [1998], and the expansion is as rapid as 1 hour (MLT) per ~20 second [*Milling et al.*, 2008]. Currently, however, there is no report on whether Pi1B waves can expand to the dayside.

From both Figure 3 and 4, the maximum correlations here peak around ~8s and exhibit a global property, although they are strongest around the 15-21 MLT regions. *Lee et al.*, [2001] has shown that impulsive field-aligned currents are strongly excited near the boundary between magnetospheric cold plasma and plasma sheet hot plasmas. This corresponds to circumstances when the Alfvén speed undergoes a rapid variation, and thus intensive shear Alfvén waves can be excited through mode conversion. The indications here need further analysis to separate the effects of wave propagation and temporal amplitude variation in order to confirm the behaviour. From these signals, however, we can put forward a possible scenario that either FMS waves or shear Alfvén waves (which may be generated in the BPS) can propagate to the ionosphere either obliquely or field aligned, and therefore could be observed by Swarm in addition to the FACs at all MLT; with the strongest signal centring around 15-21 MLT. We can't confirm, but it is possible that the Swarm LEO is at just the appropriate altitude for the Pi1B to spread globally and where dayside waves have not been completely damped. In turn, this may be the cause for the absence of dayside Pi1B observations in ground based data. This global characteristic observed by Swarm may broaden our horizon on the association of ULF waves and FACs, as well as the sub-storms. We will clarify this potential capability in future work.

4 Conclusions

To explore the local time dependence and stability of FACs at Swarm altitudes, we have investigated their statistical, dual-spacecraft correlation signatures between two Swarm spacecraft, flying side by side from 17th Apr. 2014 to 30th Apr. 2016, using a method which links the correlation intervals to model estimated auroral boundaries (after *Xiong and Lühr*, [2014]; *Xiong et al.* [2014]). Thus, the segments are targeted relative to auroral boundaries defining the limit of current intensity from the ordinary R1 and R2 current systems. The interval between these boundaries is split into intervals most likely to contain R1 and R2 currents respectively. In fact, the R1 intervals cover latitudes which may contain influences from other

current systems, for example, Cusp currents (Region 0), NBZ currents, and the R2 intervals may contain some ordinary region 1 signals and, indeed, some other currents at noon. It is difficult to separate and distinguish these at the higher latitudes though this analysis. The unfiltered FAC data adopted here is the official Level-2 FAC data of Swarm which is obtained by using single-spacecraft methods, which assume an infinite current sheet approximation can be applied locally to each spacecraft in the manner detailed by [Ritter *et al.*, 2013]. We have then applied 20 and 60 seconds filtering to yield the low-pass filtered data and indicate the large-scale FACs (150/450 km along orbit). Cross-correlations are performed mainly on data obtained from the Swarm A and C spacecraft. The results show that the maximum correlations obtained from sliding data segments show clear trends in MLT. By connecting the average mid-positions of the two intervals from Swarm A and C used to estimate the maximum correlations, we show the current sheet orientation for LEO altitude directly for the first time. It is obvious that the large-scale current sheets closely follow the oval on the dusk-side and are also well consistent with the oval pattern on the dawn-side and this ordering is consistent with the correlation trends found. The orientations are estimated using a high (0.97) correlation level. It was noted that setting lower correlation thresholds for the current sheet orientations will introduce more influences from small-scale currents.

The results show that the R2 currents (referring to all FAC signatures at latitudes in the lower auroral boundary as defined by Xiong and Lühr [2014]) show the strongest correlations for a broad range of MLT, centred on pre-dawn and post-dusk, with a higher correlation coefficient on the dusk-side and lower correlations near noon and midnight. This is consistent with the results for the current sheet alignments, where the ordering relative to the auroral oval is strongest at lower latitudes and strongest on the dusk side. The FAC profiles are very sensitive to shifts in longitude, especially for the unfiltered data, which can be attributed to the lower applicability of the infinite current sheet approximation to the small-scale currents [Forsyth *et al.*, 2017]. Correlations are much higher for the filtered data and are more stable for up to 0.3 minutes, i.e. 20 seconds, time difference (δt) between Swarm A and C. It thus reflects the predominantly large-scale dominance of R2 FACs and little influence from the small-scale currents in this region. In contrast, the R1 currents (actually all high latitude currents) peak mainly at the dawn and dusk side, and are maintained for a shorter range of δt , consistent with the expectation that R1 currents are more temporally variable.

Evidence is also found for the influence from other current systems such as Region 0 and NBZ currents in the R1 region. Correlations between spacecraft A-B and B-C show little δt or δlon sensitivity, despite persistent variabilities below 44 minutes, down to 0 minutes. This may be because of the influence from different altitudes of each spacecraft, i.e. 470km, versus 531km, which compete with the δt or δlon sensitivity. However, another possibility is the temporal stability of the large-scale FACs R1/2 FACs.

To test the influence from geomagnetic activity, however, the data has been separated into two segments (AE index below 110 nT and above 110 nT, with nearly the same data number in each segment). However, the behaviour is very similar, where the only obvious effect is that the correlations are higher during geomagnetic activated time. This result is not shown here.

The evidence further suggests that the higher latitude FACs are modulated by ULF waves, which seem to be Pi1B waves in the Alfvén mode with a frequency of ~8 seconds. The trends are prominent for the unfiltered data set, indicating a relationship between the small-scale currents and the Pi1B waves. However, secondary harmonic waves seem to appear for the 60s filtered FAC data. This analysis illustrates a new way to reveal pulse observations using LEO satellites. This result arises from a statistical study and is hard to be found from case-by-case studies because of the fast motion of the LEO satellites. However, more work needs to be done to clarify this result.

The methodology, based on the correlation of single spacecraft estimates, was deliberately chosen to highlight the locations of small-scale influences, where these add to the larger scale trends. Generally speaking, therefore, this study clearly reflects the two different domains of FACs: small-scale (some tens of km) which are time variable and large-scale (>100 km) which are rather stationary. The study is very

supportive of the dual-SC FAC approach introduced by Ritter et al and explored recently by others (e.g. Dunlop et al., 2015b). The study suggests the time shifts and filters used in multi spacecraft techniques are generally suitable for accurate determination of the FACs and perhaps allows the conditions where these break down to be further investigated. The evidence further suggests that the higher latitude FACs are modulated by ULF waves, which seem to be Pi1B waves in the Alfvén mode with a frequency of ~8 seconds. The trends are prominent for the unfiltered data set, indicating a relationship between the small-scale currents and the Pi1B waves. However, secondary harmonic waves seem to appear for the 60s filtered FAC data. This analysis illustrates a new way to reveal pulse observations using LEO satellites. This result arises from a statistical study and is hard to be found from case-by-case studies because of the fast motion of the LEO satellites. However, more work needs to be done to clarify this result.

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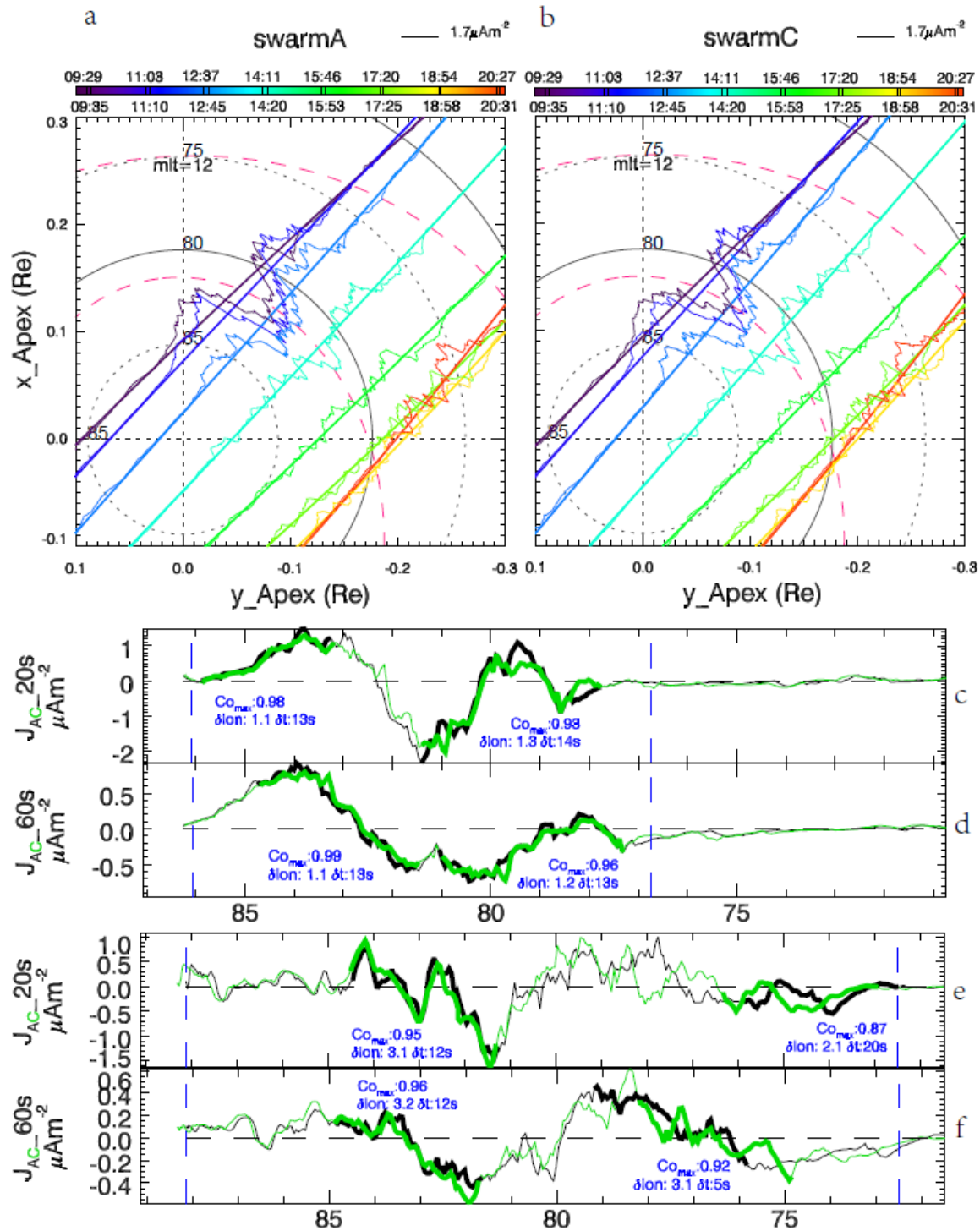
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680

681 **Figure 1:** Swarm A/C FACs for an example interval on 6th, June, 2014, corresponding to a sequence of
 682 consecutive orbits (colour coded with time from black to red shown above the colour bar in **a**) and **b**),
 683 where: **a**) shows 20s filtered FACs observed by Swarm A plotted on the orbits within time period 09:29-
 684 20:31 UT where the FAC magnitude scale is denoted on the top-right, and **b**) shows the same as **a**) but for

Swarm C. The model [Xiong and Lühr, 2014] poleward and equatorward auroral boundaries of current intensity on 14:30 UT are shown on Figure 1 (**a**, **b**) by the magenta dashed curves. The lower panels (**c** to **f**) show two, dual spacecraft (Swarm A (black) and Swarm C (green)) intervals within this sequence, plotted as a function of APEX latitude, and the current value not on the same scale: **c**) and **d**) show the first north descending orbit track (black orbit in panels **a** and **b**), where the R1 and R2 boundaries are estimated to be at 9:30 and 9:33 UT, and **e**) and **f**) are for three orbits later (light blue orbit in panels **a** and **b**) with boundaries at 14:13 and 14:18 UT. Each pair shows the 20s (upper panels, **c** and **e**) and 60s (lower panels, **d** and **f**) moving average data and also indicates the sliding maximum correlation achieved for the two intervals between R1 and R2 boundaries (blue dashed lines) with the longitude and time shift between A and C for each interval (blue text).

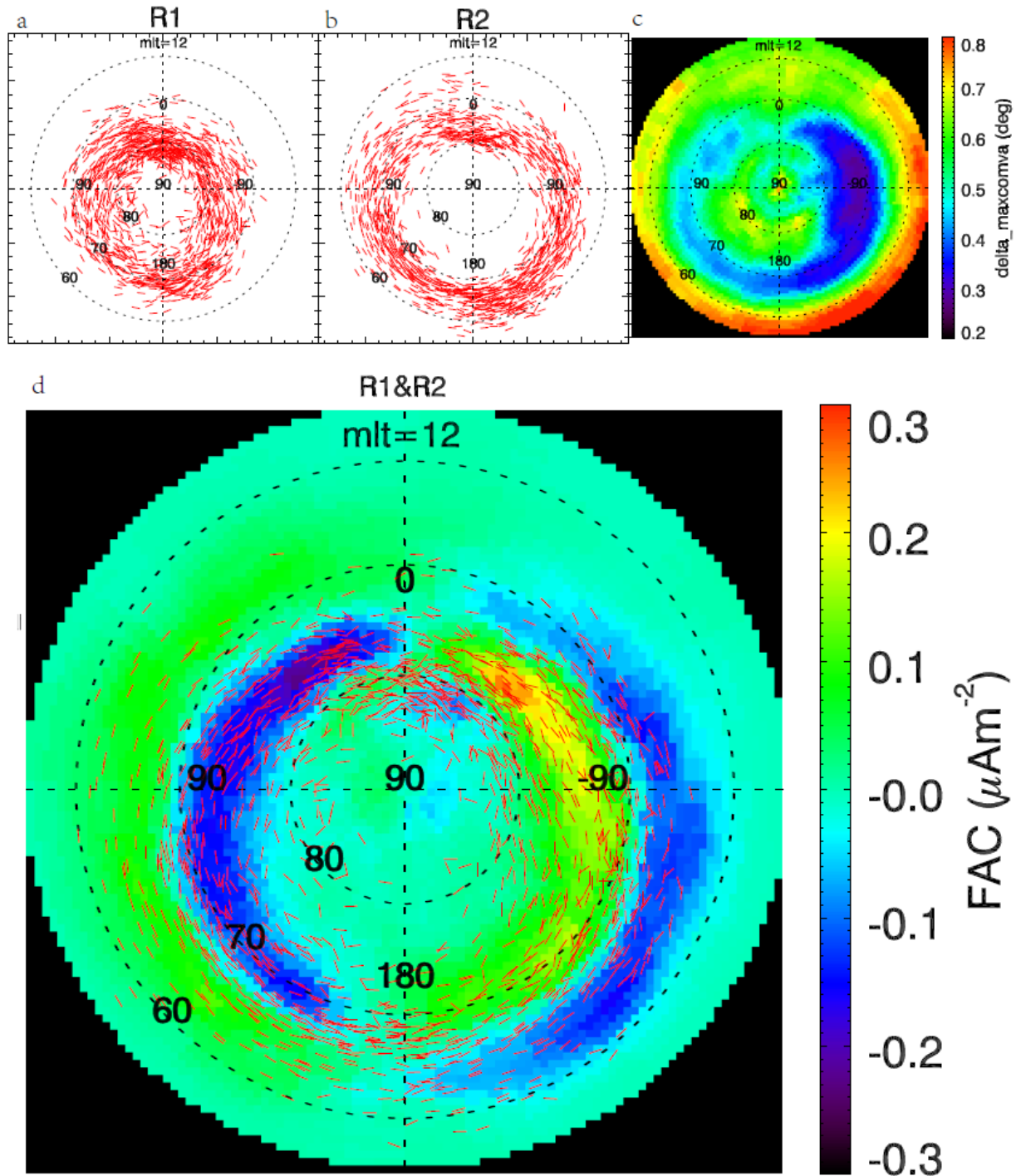


Figure 2: **a, b).** Northern hemisphere polar map, showing inferred current sheets, for Swarm A and C data from 17th Apr. 2014 to 20th Aug. 2014, during which time Swarm A and C's orbit has covered 24hrs MLT. These are plotted using lines of normalised length which connect the average Swarm A and C positions of those orbit segments producing the maximum correlations (drawn for 20 seconds filtered data), where: **a)** shows the current sheet orientations found for the higher latitude regions, **b)** shows those for the lower latitude regions, **c)** shows the intersection angle of current sheets calculated by maximum correlation and MVA method, for Swarm A and C data from 17th Apr. 2014 to 30th Apr. 2016 and **d)** shows a Northern hemisphere polar map, showing the average FACs for Swarm A and C data from 17th

Apr. 2014 to 30th Apr. 2016, overlain with similar current sheet orientations to **a)** and **b)** but for both higher latitude regions and lower latitude regions.

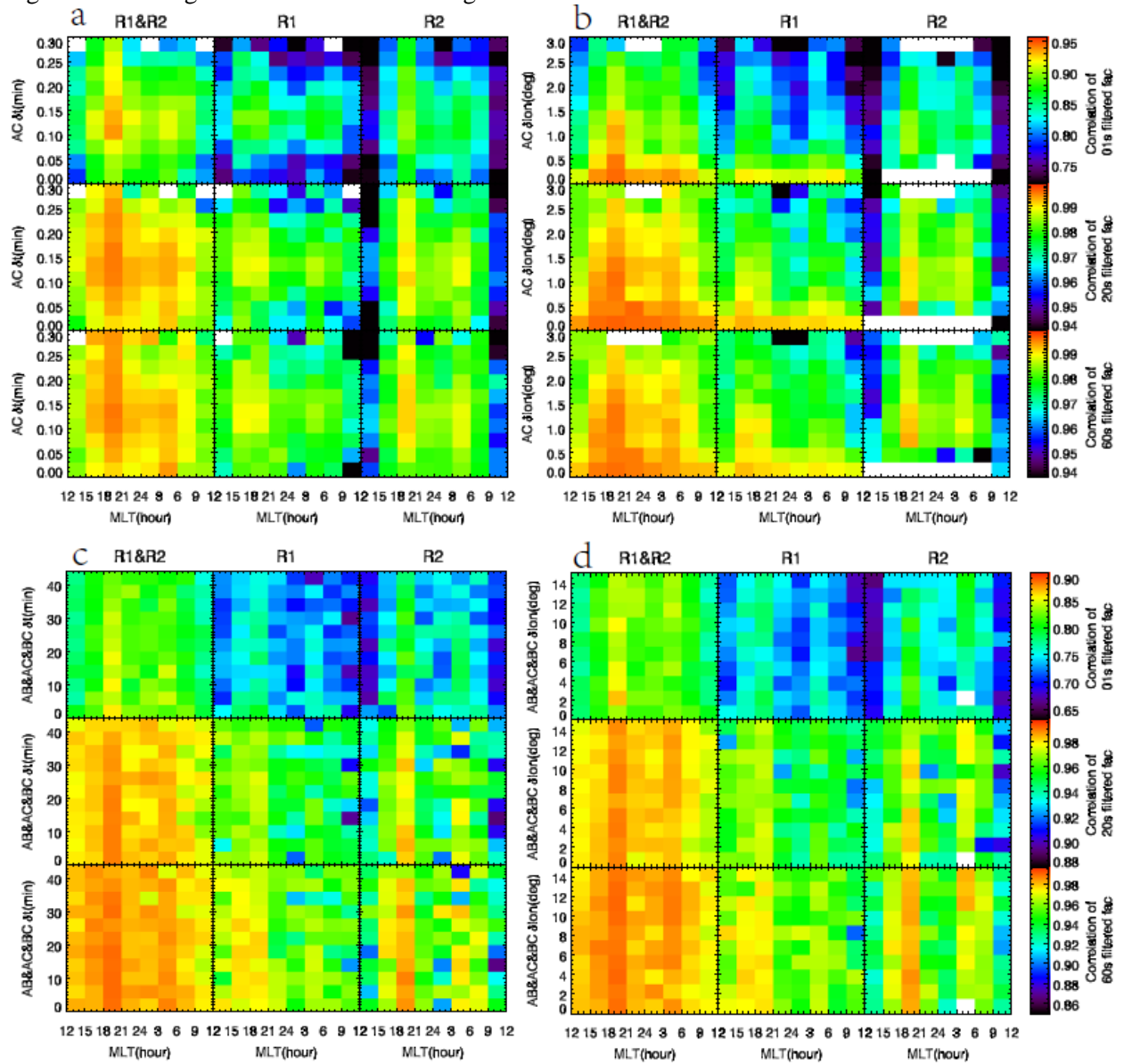


Figure 3: **a)** nine panels showing vertically from the top the A-C correlation, binned with the δt difference, as a function of MLT, for unfiltered data, 20 seconds low pass filtered data and 60 seconds low pass filtered data. The three columns are for correlations over the whole interval between the R1 and R2 boundaries and for the intervals covering R1 and R2 respectively (denoted by R1&R2, R1 and R2 each at the top); **b)** similar array of correlations but binned with δlon ; **c, d)** similar array of correlations but of A-B and B-C correlation.

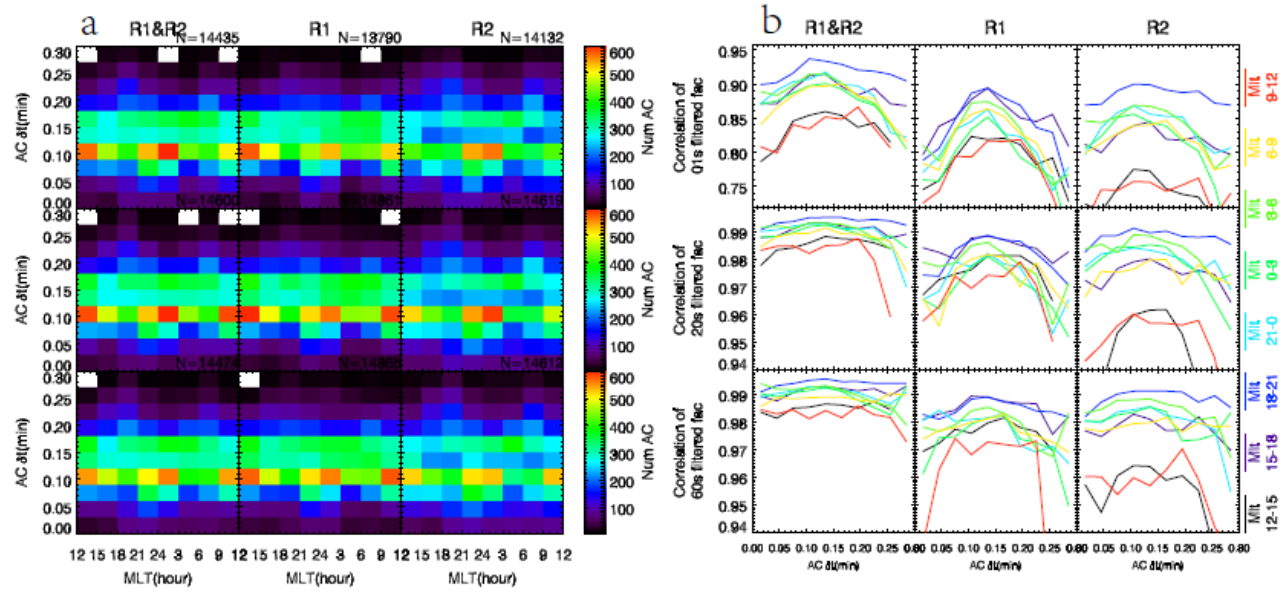


Figure 4: a) number of cases in each bin of Figure 3a, and b) similar information to that in Figure 3a but correlations are indicated by the y axis on the left of each panel, here different colours represent different MLTs and the δt difference is now shown by the x axis.

Figure 1.

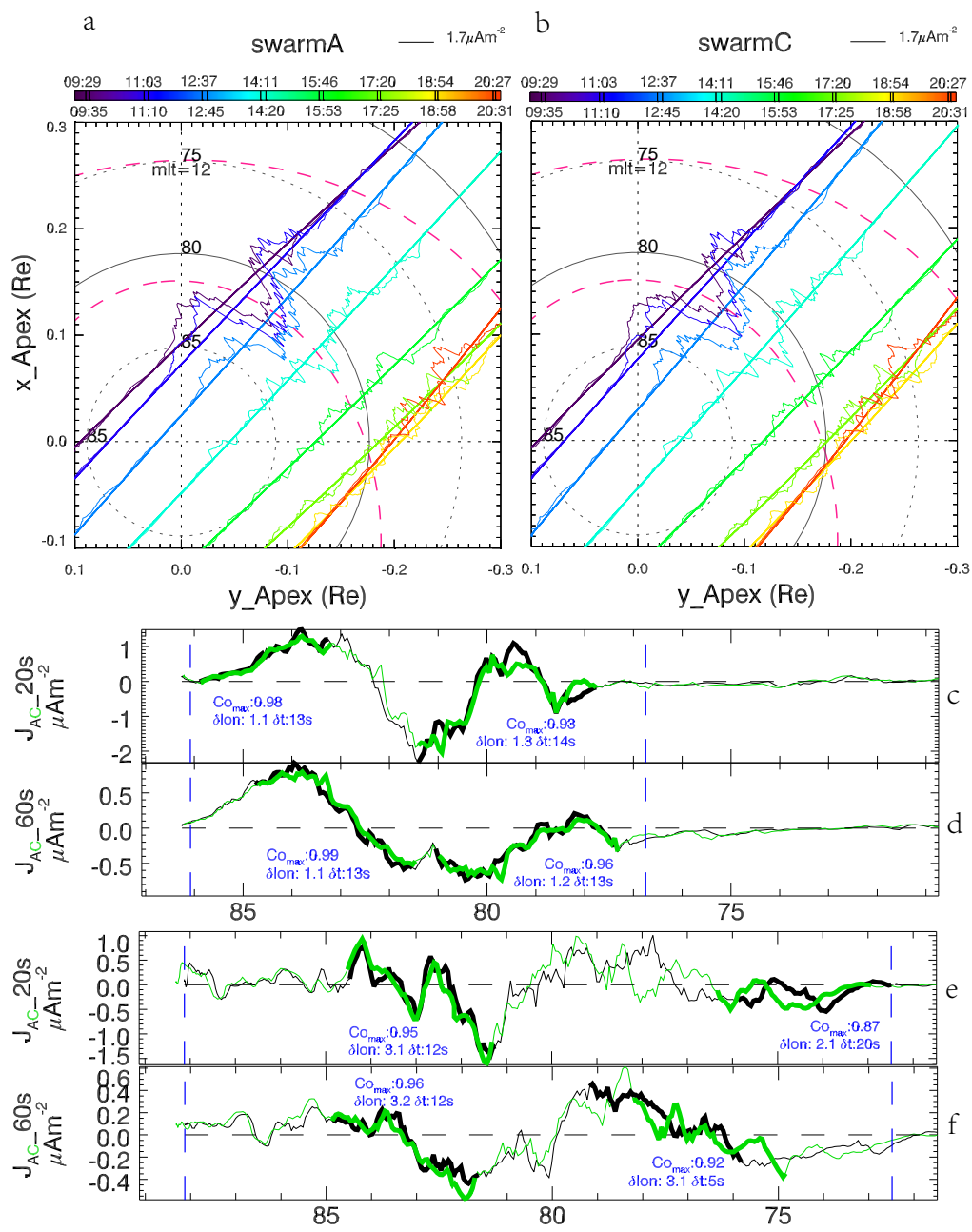


Figure 2.

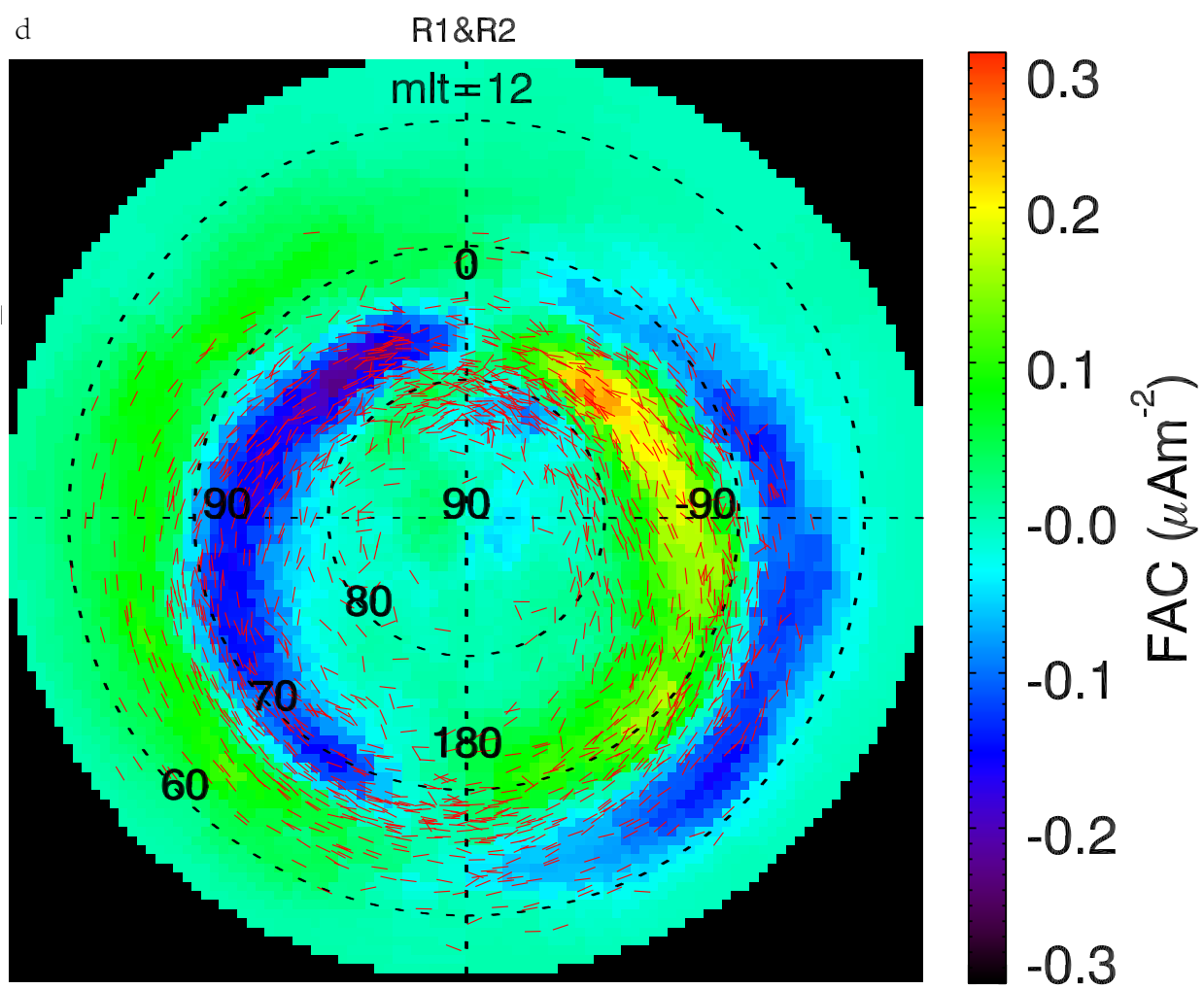
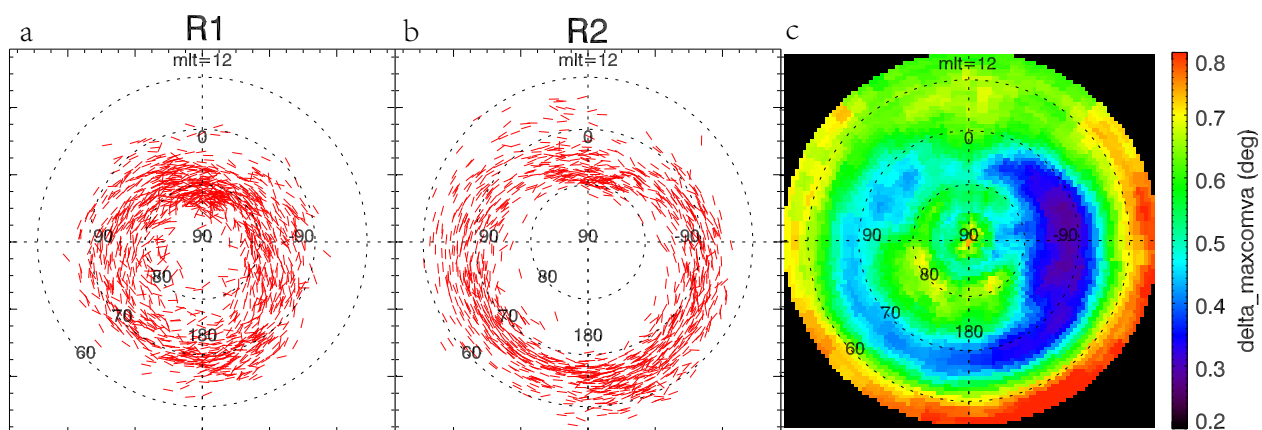


Figure 3.

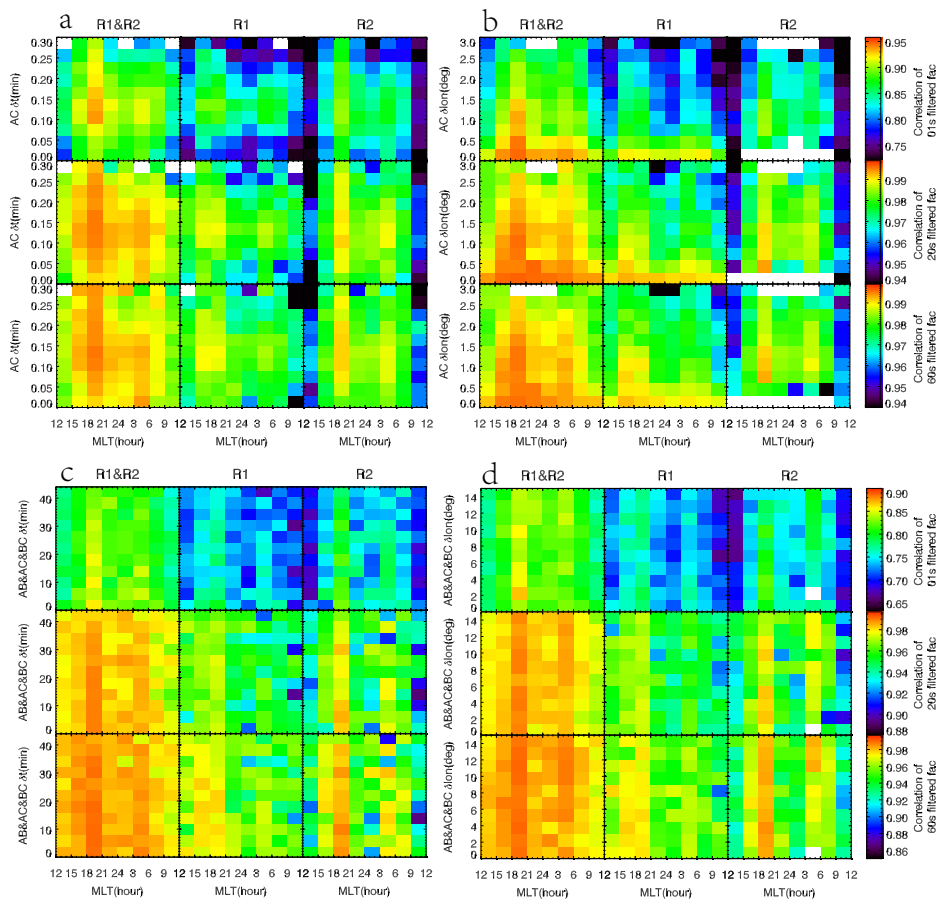


Figure 4.

