

INFLOWS, OUTFLOWS, AND A GIANT DONOR IN THE REMARKABLE RECURRENT NOVA
M31N 2008-12a? — *HUBBLE SPACE TELESCOPE* PHOTOMETRY OF THE 2015 ERUPTION

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ABSTRACT

The recurrent nova M31N 2008-12a experiences annual eruptions, contains a near-Chandrasekhar mass white dwarf, and has the largest mass accretion rate in any nova system. In this paper, we present *Hubble Space Telescope* (*HST*) WFC3/UVIS photometry of the late decline of the 2015 eruption. We couple these new data with archival *HST* observations of the quiescent system and Keck spectroscopy of the 2014 eruption. The late-time photometry reveals a rapid decline to a minimum luminosity state, before a possible recovery / re-brightening in the run-up to the next eruption. Comparison with accretion disk models supports the survival of the accretion disk during the eruptions, and uncovers a quiescent disk mass accretion rate of the order of $10^{-6} M_{\odot} \text{ yr}^{-1}$, which may rise beyond $10^{-5} M_{\odot} \text{ yr}^{-1}$ during the super-soft source phase – both of which could be problematic for a number of well-established nova eruption models. Such large accretion rates, close to the Eddington limit, might be expected to be accompanied by additional mass loss from the disk through a wind and even collimated outflows. The archival *HST* observations, combined with the disk modeling, provide the first constraints on the mass donor; $L_{\text{donor}} = 103_{-11}^{+12} L_{\odot}$, $R_{\text{donor}} = 14.14_{-0.47}^{+0.46} R_{\odot}$, and $T_{\text{eff, donor}} = 4890 \pm 110 \text{ K}$, which may be consistent with an irradiated M31 red-clump star. Such a donor would require a system orbital period $\gtrsim 5$ days. Our updated analysis predicts that the M31N 2008-12a WD could reach the Chandrasekhar mass in < 20 kyr.

Keywords: Galaxies: individual: M31 — novae, cataclysmic variables — stars: individual: M31N 2008-12a — ultraviolet: stars — accretion, accretion disks

1. INTRODUCTION

Novae are a sub-class of the cataclysmic variables (CVs), where a white dwarf (WD) accretes hydrogen-rich matter from a donor star within a, typically close, binary system (see [Bode & Evans 2008](#); [Woudt &](#)

[Ribeiro 2014](#), for review articles). The transferred material usually accumulates in an accretion disk around the WD, but there may also be some element of magnetic accretion at play, depending upon the strength of the WD's magnetic field. Novae distinguish themselves from CVs by virtue of their typically elevated WD mass

accretion rates (\dot{M}_{acc}) and by the nova eruption itself – a thermonuclear runaway within the accreted envelope on the WD surface (see [Starrfield et al. 1976](#)). All novae are inherently recurrent, but their inter-eruption period depends upon the WD mass (M_{WD}) and \dot{M}_{acc} . Systems that combine a large M_{WD} with a high \dot{M}_{acc} exhibit the shortest recurrence periods, and have often been observed in eruption more than once – the recurrent novae (RNe; see [Schaefer 2010](#)). Observed recurrence periods lie in the range $1 \leq P_{\text{rec}} \leq 98$ years (see [Darnley et al. 2014](#); [Pagnotta et al. 2009](#), respectively), where both ends are probably limited by selection effects (see, for e.g., [Kato et al. 2014](#); [Hillman et al. 2016](#); [Shafter 2017](#)).

In most cases, the high values of \dot{M}_{acc} in the RNe are driven by elevated mass loss rates from evolved donors ([Darnley et al. 2012](#)). This is observed to be via Roche lobe overflow of a sub-giant donor (e.g., U Scorpii), or by accretion from the stellar wind of a giant (e.g., RS Ophiuchi, see e.g., [Evans et al. 2008](#)) – both mechanisms lead to an accretion disk around the WD. A handful of RNe, possibly ‘transient’ (rather than long-term) recurrences, such as T Pyxidis, may show evidence of elevated mass transfer driven by the irradiation of their main sequence donors ([Knigge et al. 2000](#); [Godon et al. 2014](#)).

M31N 2008-12a, a RN residing within M31, is the most extreme nova system yet discovered. With an observed $P_{\text{rec}} \simeq 1$ year, it is the prototype of the newly emerging class of ‘rapidly recurring nova’; those with $P_{\text{rec}} \lesssim 10$ years. First detected in 2008, M31N 2008-12a has been discovered in eruption every year since (2008–2016; [Darnley et al. 2014, 2015c, 2016](#); [Itagaki et al. 2016](#)), with three previous eruptions recovered from archival X-ray observations (1992, 1993, 2001; [Henze et al. 2014](#); [Tang et al. 2014](#)). For reference, the observed eruption history is summarised in Table 1 (see [Darnley et al. 2016](#), for a detailed description). By analysis of the eight eruptions between 2008–2015, [Darnley et al. \(2016, hereafter DHB16\)](#) reported $P_{\text{rec}} = 347 \pm 10$ days. However, when including the earlier X-ray detections, [Henze et al. \(2015a, hereafter HDK15\)](#) suggested that P_{rec} could even be as short as 174 ± 10 days.

These rapid-fire eruptions of M31N 2008-12a are powered by the most massive accreting WD yet discovered. Studies of the 2013 eruption yielded $M_{\text{WD}} > 1.3 M_{\odot}$ ([Tang et al. 2014, hereafter TBW14](#)), with a more recent determination of $M_{\text{WD}} = 1.38 M_{\odot}$ ([Kato et al. 2015](#)). We note that the WD mass has not been measured directly, only estimated based on modeling of the system. That same modeling required very large accretion rates, $\dot{M}_{\text{acc}} > 1.7 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ and $\dot{M}_{\text{acc}} = 1.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$, respectively. Under an assumption of spherical ejecta, [Henze et al. \(2015b, hereafter HND15\)](#) concluded that the quantity of ejected hydrogen was $M_{\text{e,H}} = (2.6 \pm 0.4) \times 10^{-8} M_{\odot}$, broadly consistent with the total ejected mass prediction of $M_{\text{e}} = 6 \times 10^{-8} M_{\odot}$ from [Kato et al. \(2015\)](#). This indicates a

Table 1. A summary of the twelve observed eruptions of M31N 2008-12a.

Eruption date ^a (UT)	Inter-eruption time-scale (days) ^b	References
(1992 Jan. 28)	...	1, 2
(1993 Jan. 03)	341	1, 2
(2001 Aug. 27)	...	2, 3
2008 Dec. 25	...	4
2009 Dec. 02	342	5
2010 Nov. 19	352	2
2011 Oct. 22.5	337.5	5, 6–8
2012 Oct. 18.7	362.2	8–11
2013 Nov. 26.95 ± 0.25	403.5	5, 8, 11–14
2014 Oct. 02.69 ± 0.21	309.8 ± 0.7	8, 15
2015 Aug. 28.28 ± 0.12	329.6 ± 0.3	14, 16–18
2016 Dec. 12.32	471.72	19, 20

^aEruption dates in parentheses have been estimated based on an extrapolation of available X-ray data (see [Henze et al. 2015a](#)).

^bThe inter-eruption time-scale is only given when consecutive eruptions have been detected (assuming $P_{\text{rec}} \simeq 1$ year).

NOTE—Compact version of a Table originally published by [Tang et al. \(2014\)](#) and updated by [Darnley et al. \(2016\)](#).

References—(1) [White et al. \(1995\)](#), (2) [Henze et al. \(2015a\)](#), (3) [Williams et al. \(2004\)](#), (4) [Nishiyama & Kabashima \(2008\)](#), (5) [Tang et al. \(2014\)](#), (6) [Korotkiy & Elenin \(2011\)](#), (7) [Barsukova et al. \(2011\)](#), (8) [Darnley et al. \(2015c\)](#), (9) [Nishiyama & Kabashima \(2012\)](#), (10) [Shafter et al. \(2012\)](#), (11) [Henze et al. \(2014\)](#), (12) [Tang et al. \(2013\)](#), (13) [Darnley et al. \(2014\)](#), (14) [Darnley et al. \(2016\)](#), (15) [Henze et al. \(2015b\)](#), (16) [Darnley et al. \(2015a\)](#), (17) [Darnley et al. \(2015b\)](#), (18) [Henze et al. \(2015c\)](#), (19) [Itagaki et al. \(2016\)](#), (20) [Henze et al. \(2017b\)](#).

mass accretion efficiency of $\sim 63\%$, not only is the WD massive, but it is growing.

[Darnley et al. \(2014, hereafter DWB14\)](#) and [TBW14](#) both illustrated the rapid optical development of the 2013 eruption; [Henze et al. \(2014, hereafter HND14\)](#) and [TBW14](#) noted the rapid X-ray development. [DHB16](#) combined all data from the near-identical 2013, 2014, and 2015 eruptions to determine that the optical decay time (t_2 ; the time to decay two magnitudes from the peak luminosity) is only 1.65 ± 0.04 days, and $t_3 = 2.47 \pm 0.06$ days. The accompanying super-soft X-ray source (SSS) ‘turned on’ only 5.6 ± 0.7 days after the 2015 eruption, and turned off after 18.6 ± 0.7 days ([DHB16](#)); only the Galactic RN V745 Scorpii displays more rapid X-ray evolution ([Page et al. 2015](#)).

[DHB16](#) also presented a detailed analysis of the combined spectra of the 2012–2015 eruptions. The earliest post-eruption spectra show fleeting evidence of very high velocity ($v_{\text{ej}} \simeq 13,000 \text{ km s}^{-1}$) outflows. [DHB16](#) proposed that these could be due to a high level of ejecta collimation, in the polar direction, almost along the line

of sight. Darnley et al. (2017, hereafter DHG17) reported similar high velocity material surround the far UV N v (1240 Å) emission line 3 days after the 2015 eruption, again this was linked to possible ejecta collimation or jets from the eruption.

Hints of ejecta deceleration were first reported by TBW14. The DHB16 analysis of the combined 2012–2015 spectroscopy found clear evidence of significant ejecta deceleration, consistent with the adiabatic expansion of a forward shock (cf. Bode & Kahn 1985). DHS15 and DHB16 both proposed that this deceleration could be caused by the ejecta interacting with pre-existing circumbinary material. Given that the circumbinary regime should be cleared by each annual eruption, this environment must be regularly resupplied. Therefore, DHB16 proposed that the M31N 2008-12a donor should be a giant with a significant stellar wind, and not Roche lobe overflow.

Utilising the *Swift* observatory, Kato et al. (2016) undertook the first targeted survey to detect the long-predicted X-ray flash precursor to a nova eruption (see, e.g., Starrfield et al. 1990; Krautter 2002). The campaign was unsuccessful, possibly because of the earlier than predicted 2015 eruption, or because the flash was absorbed by pre-existing material surrounding the system. At the time, there was no strong evidence constraining the mass donor in the system, therefore Kato et al. (2016) favoured the former explanation.

Containing a growing WD that is already close to the Chandrasekhar limit, M31N 2008-12a is therefore the leading pre-explosion supernova type Ia candidate system. *Hubble Space Telescope* (HST) spectroscopy of the 2015 eruption conducted by DHG17 found no evidence of neon within the ejecta. However, as discussed by those authors, that single result cannot yet completely rule out the presence of an ONe WD in the system. Either way, DHG17 argued that the lack of an observational signature of Ne may in itself indicate that the M31N 2008-12a WD is growing in mass. That is, either a CO WD has grown to the Chandrasekhar limit, or a large enough He layer has been accumulated to shield an underlying ONe WD from the nova eruptions.

In this paper, we present the results of a HST program to study the late-decline of the predicted 2015 eruption of M31N 2008-12a, and an updated analysis of archival *HST* and Keck observations of the system. In Section 2 we describe our observations, in Section 3 we present the photometric data. In Section 4 we explore models of the accretion disk in M31N 2008-12a. Finally, in Sections 5 and 6 we discuss our findings and present the subsequent conclusions.

While this manuscript was being prepared, the 2016 eruption of M31N 2008-12a was detected by Itagaki et al. (2016). The observations of the 2016 eruption will be presented in Henze et al. (2017b).

2. OBSERVATIONS

2.1. *Hubble Space Telescope* Observations

Twenty orbits of *HST* Cycle 23 time were awarded to collect early-time UV spectroscopic observations (8 orbits) and late-time imaging of the 2015 eruption of M31N 2008-12a (proposal ID: 14125). The results of the spectroscopy were presented in DHG17. The 2015 eruption was discovered on 2015 Aug. 28.425 UT by an automated monitoring program on the Las Cumbres Observatory 2 m telescope¹ on Hawai'i (Darnley et al. 2015a, see DHB16 for full details). The *HST* photometric observations were conducted between 2015 Sep. 10 and Sep. 30, a log of these observations is provided in Table 2.

We employed 12 *HST* orbits, split into four visits, to collect photometry of M31N 2008-12a using the Wide Field Camera 3 (WFC3) in the UVIS mode. Each visit used identical observing strategies and were approximately one week apart, starting at $\Delta t \simeq 14$ days (post-eruption). Observations were obtained using the WFC3/UVIS F225W, F275W, F336W, F475W, and F814W filters.

For each filter a two-point dither was applied to enable removal of detector defects. To reduce readout overheads, WFC3/UVIS was operated in a $2k \times 2k$ windowed mode utilizing the UVIS2-2K2C-SUB aperture. This part of the chip was selected for its superior performance against charge transfer efficiency (CTE) loss, to further mitigate such effects we included a ‘post flash’ signal of 9–12 electrons.

The WFC3/UVIS data were reduced using the STScI `calwf3` pipeline (v3.1.6; see Dressel 2012), with CTE correction manually applied via the `wfc3uv_ctereverse_parallel` code (v2015.07.22²; see also Anderson et al. 2012). Photometry of the WFC3/UVIS data was then performed on individual exposures using DOLPHOT (v2.0³; Dolphin 2000, following the standard procedure and parameters for WFC3/UVIS given in the manual). For comparative purposes, photometry was also carried out using the combined exposures per epoch for each filter. All data were aligned and final combined images created using the Drizzlepac (v2.0.2) `astrodrizzle` package. Photometry was obtained via the PyRAF `phot` package (v2.2). The results from the DOLPHOT and `phot` methods are consistent. For comparison with previous work, we adopt the DOLPHOT photometry, which is presented in Table 3.

2.2. *Keck* spectroscopy of the 2014 eruption

DHB16 observed that photometrically, the 2013, 2014, and 2015 eruptions were essentially identical; the same is true of the spectra from the 2012–2015 eruptions. Therefore, to support the late-time HST photometry of the 2015 eruption, we also utilise a Keck spectrum of

¹ Formerly known as the Faulkes Telescope North

² http://www.stsci.edu/hst/wfc3/tools/cte_tools

³ <http://americano.dolphinim.com/dolphot>

Table 2. Log of observations of the eruptions of M31N 2008-12a referred to in this Paper.

Eruption	Facility	Instrument	<i>HST</i> Visit	Date (midpoint)	Start $t - t_0$ (days)	End (days)	Orbits	Exposure time (ks)
2014	Keck I	LRIS	...	2014 Oct 21.50	18.80	18.82	...	1.2
2015	<i>HST</i>	WFC3/UVIS	4	2015 Sep 10.64	13.27	13.44	3	6.8
2015	<i>HST</i>	WFC3/UVIS	5	2015 Sep 17.66	20.29	20.47	3	6.8
2015	<i>HST</i>	WFC3/UVIS	6	2015 Sep 23.62	26.26	26.43	3	6.8
2015	<i>HST</i>	WFC3/UVIS	7	2015 Sep 30.58	33.22	33.39	3	6.8

Table 3. *Hubble Space Telescope* WFC3/UVIS NUV and visible photometry of M31N 2008-12a following the late decline of the 2015 eruption.

Date (UT)	Δt (days)	Exposure (secs)	Filter	S/N	Photometry (Vega mag)
2015-09-10.669	13.389 ± 0.027	2×870	F225W	55.0	20.83 ± 0.02
2015-09-17.693	20.413 ± 0.027	2×870	F225W	36.5	21.83 ± 0.03
2015-09-23.656	26.376 ± 0.027	2×870	F225W	25.8	22.31 ± 0.04
2015-09-30.333	33.333 ± 0.027	2×870	F225W	21.2	22.83 ± 0.05
2015-09-10.564	13.284 ± 0.007	2×519	F275W	38.5	20.80 ± 0.03
2015-09-17.587	20.307 ± 0.007	2×519	F275W	25.1	22.18 ± 0.04
2015-09-23.550	26.270 ± 0.007	2×519	F275W	19.9	22.52 ± 0.06
2015-09-30.228	33.228 ± 0.007	2×519	F275W	18.4	22.79 ± 0.06
2015-09-10.580	13.300 ± 0.007	2×519	F336W	76.0	21.08 ± 0.01
2015-09-17.603	20.324 ± 0.007	2×519	F336W	47.2	22.08 ± 0.02
2015-09-23.566	26.286 ± 0.007	2×519	F336W	35.1	22.61 ± 0.03
2015-09-30.245	33.245 ± 0.007	2×519	F336W	33.2	22.75 ± 0.03
2015-09-10.708	13.428 ± 0.010	2×745	F475W	130.7	22.48 ± 0.01
2015-09-17.733	20.453 ± 0.010	2×745	F475W	83.3	23.56 ± 0.01
2015-09-23.696	26.416 ± 0.010	2×745	F475W	65.4	23.97 ± 0.02
2015-09-30.372	33.372 ± 0.010	2×745	F475W	60.1	24.18 ± 0.02
2015-09-10.629	13.349 ± 0.010	2×765	F814W	85.7	22.37 ± 0.01
2015-09-17.654	20.374 ± 0.010	2×765	F814W	47.9	23.39 ± 0.02
2015-09-23.617	26.337 ± 0.010	2×765	F814W	34.7	23.84 ± 0.03
2015-09-30.293	33.293 ± 0.010	2×765	F814W	34.9	23.91 ± 0.03

the 2014 eruption taken 18.81 days after that eruption.

This 2014 Keck spectrum has not been published until now. It was collected using the Low Resolution Imaging Spectrometer (LRIS; Oke et al. 1995; McCarthy et al. 1998; Rockosi et al. 2010), which is mounted at the Cassegrain focus of the Keck I telescope on Mauna Kea, Hawai'i. The spectrum was obtained through the standard low-resolution configuration using the 400/3000 grism (blue camera) and 400/8500 grating (red camera), providing continuous coverage from the atmospheric cutoff to approximately 10300 Å. However, as the nova had faded significantly, crowding and confusion with nearby stars in M31 had started to be problematic, therefore, the object is only clearly detected in the blue camera; only data with $\lambda < 5600$ Å are analyzed here.

2.3. Archival Quiescent Data

The Panchromatic Hubble Andromeda Treasury (PHAT; Dalcanton et al. 2012) was a broadband, multi-colour, NUV–NIR *HST* survey of the bulge and north-eastern disk of M31. As part of the PHAT survey, M31N 2008-12a was observed between eruptions a number of times with *HST*. Initial results from analysis of these data were published in DWB14 and TBW14. Both those works analyzed the optical and NUV *HST* data finding evidence for a very blue source coincident with M31N 2008-12a, indicating the presence of a luminous accretion disk. Although the available *HST* NIR data were also analyzed, DWB14 and TBW14 only presented upper limits on the quiescent photometry of M31N 2008-12a, which was severely blended with nearby sources in the NIR. These upper limits did not place firm constraints on the nature of the donor, only excluding the most luminous red giants (such as that found in the

T Coronae Borealis system). Notably, the initial analysis of the quiescent SED indicated an accretion disk roughly similar in flux distribution, albeit brighter, to that in the RS Oph system; a donor of similar luminosity to the red giant in RS Oph was not ruled out by DWB14.

Williams et al. (2014a) released the NUV to NIR photometric catalog from the PHAT survey, which included the quiescent photometry of M31N 2008-12a. These photometry are provided in Table 4 and are consistent with the independent analysis by DWB14 and TBW14. However, the analysis undertaken by Williams et al. (2014a) was able to successfully de-blend the sources around M31N 2008-12a in the NIR yielding F110W and F160W photometry of the quiescent system. This superior NIR deblending was achieved by simultaneous fitting of the higher spatial resolution F475W data with the NIR data. These F475W data have spatial resolution better by more than a factor of two and allowed for much more robust deblending of crowded sources.

The observations reported in Williams et al. (2014a) are from the first set of PHAT visits and are computed over two separate *HST* visits. Data from another pair of visits are also available and the PHAT collaboration have generously supplied their photometry of M31N 2008-12a from each of the four *HST* visits, and these data are also shown in Table 4. The quiescent photometry reported by DWB14 and TBW14 are combined from observations at different phases in the full eruption cycle of M31N 2008-12a, and from different eruption cycles (as noted by both those papers).

3. LIGHT CURVE ANALYSIS

The five-band *HST* photometry was presented in Table 3, and the subsequent light curves are presented in Figure 1. In the sub-figures, the four epochs of *HST* WFC3/UVIS observations (black data points) are compared with the template M31N 2008-12a eruption light curves from other telescopes. These eruption templates are constructed from Bézier smoothed light curves of the almost identical 2013–2015 eruptions (DWB14, DHS15, DHB16). The uncertainties on the smoothed light curves are computed based on the method employed by Ashall et al. (2016). Here, we compare to the closest filter in wavelength to the *HST* filters. The F275W filter is compared to the *Swift* UVW1 data (central wavelength 2600 Å), F336W to Sloan u' -band (the Sloan ground-based data are converted from the AB system to the Vega system in this plot), F475W to B , and F814W to i' , the *Swift* UVM2 filter (2250 Å) is used for comparison to the F225W data, but these data are not particularly extensive. The solid vertical lines in each plot indicate the epochs of the SSS turn-on and turn-off. The horizontal lines indicate the minimum photometry from the two visits of the PHAT survey (see Table 4).

The *HST* WFC3/UVIS F225W data are well fit by a power law of the form $f \propto t^\alpha$, where $\alpha = -2.04 \pm 0.16$ ($\chi^2_{\text{dof}} = 2.2$). We note that this decline is therefore

consistent with the ‘middle’ relation predicted by the universal decline law of Hachisu & Kato (2006, 2007, $\alpha = -1.75$). This may be connected to the lack of strong emission lines seen in the equivalent region of the NUV spectrum (see DHG17). We also note that DHB16 found that a power-law fit to the u' -band decline of the 2015 eruption, between days 8 and 20, was consistent with the predicted ‘middle’ decline law. The *HST* data from the other filters all show significant deviation from a *single* power law, when taken in isolation, and when compared to the eruption template data – all these filters contain strong flux contributions from lines.

3.1. Quiescent Data

Comparison between the *HST* imaging of the 2015 eruption and the archival data confirm that the object proposed as the quiescent system by DWB14 and TBW14 is associated with the eruptions of M31N 2008-12a. In Table 4 we have also indicated the epoch of the archival PHAT *HST* visits in respect to the M31N 2008-12a eruption cycle. The closest PHAT observations to a known eruption are those from 2011 Jan., which took place 67 days after the 2010 Nov. eruption; significantly later, post-eruption, than the late-time decline data collected for this paper. We also note that the 2011 Aug. observations took place 52 days before the 2011 Oct. eruption.

HDK15 presented evidence that M31N 2008-12a may erupt every ~ 6 months, rather than annually. If this is the case, we must also assess whether the interpretation of the archival *HST* data may be affected by *unobserved* eruptions. The typical eruption date uncertainty is 26 days (HDK15). If we utilize the dates of observed eruptions but assume a ~ 6 month cycle (see HDK15), we can investigate how close to an *unobserved* eruption each PHAT visit potentially occurred (recorded in Table 4). The only observations of note here are those from 2010 Aug. which may lie 68 ± 26 days after an unobserved *early* 2010 eruption. The F475W and F814W data from that time are significantly brighter than those from 2012 Jan. (80 days after the 2011 Oct. eruption), which suggests that these data *may* be coincident with the late decline of an *early* (but missed) 2010 eruption.

By assuming that all M31N 2008-12a eruptions are essentially identical we can roughly fit the 2010 Aug. *HST* observations to the 2015 eruption late decline observations. Therefore, we would predict that a missed eruption of M31N 2008-12a could have occurred on 2010 Jul. 09 $^{+4}_{-3}$ (see the light-blue data points in Figure 1). However, data from PTF rule out an additional eruption between 2010 Jun. 30 and the date of the *observed* 2010 Nov. eruption (Cao et al. 2012, M. M. Kasliwal priv. comm.). As such, we conclude that all PHAT data of M31N 2008-12a were taken at least 67 days after an eruption, and that they represent observations of the inter-eruption, or quiescent, period. We stress that this does not rule out the possibility of an *early* 2010 eruption occurring before this window.

Table 4. PHAT multicolor NUV, optical, and NIR photometry of M31N 2008-12a, in part from [Williams et al. \(2014a\)](#).

Date (UT)	Observed eruptions		Predicted eruptions		<i>HST</i>	Filter	Exposure Time	Photometry
	Δt_{after} (days)	Δt_{before} (days)	Δt_{after} (days)	Δt_{before} (days)	Instrument			
2011 Jan. 25.21	67	270	67	115 ± 26	WFC3/UVIS	F275W	350	$23.13 \pm 0.12^{\dagger}$
2011 Jan. 25.23	67	270	67	115 ± 26	WFC3/UVIS	F275W	660	$22.98 \pm 0.07^{\dagger}$
2011 Aug. 31.51	285	52	103 ± 26	52	WFC3/UVIS	F275W	350	22.73 ± 0.09
2011 Aug. 31.53	285	52	103 ± 26	52	WFC3/UVIS	F275W	575	22.53 ± 0.06
2011 Jan. 25.20	67	270	67	115 ± 26	WFC3/UVIS	F336W	550	$23.07 \pm 0.05^{\dagger}$
2011 Jan. 25.22	67	270	67	115 ± 26	WFC3/UVIS	F336W	800	$23.01 \pm 0.04^{\dagger}$
2011 Aug. 31.51	285	52	103 ± 26	52	WFC3/UVIS	F336W	550	22.59 ± 0.04
2011 Aug. 31.52	285	52	103 ± 26	52	WFC3/UVIS	F336W	700	22.59 ± 0.04
2010 Aug. 07.53	248	104	68 ± 26	104	ACS/WFC	F475W	600	$24.08 \pm 0.02^{\dagger}$
2010 Aug. 07.53	248	104	68 ± 26	104	ACS/WFC	F475W	370	$24.06 \pm 0.03^{\dagger}$
2010 Aug. 07.54	248	104	68 ± 26	104	ACS/WFC	F475W	370	$24.01 \pm 0.03^{\dagger}$
2010 Aug. 07.54	248	104	68 ± 26	104	ACS/WFC	F475W	370	$24.08 \pm 0.03^{\dagger}$
2012 Jan. 10.12	80	282	80	118 ± 26	ACS/WFC	F475W	700	24.46 ± 0.03
2012 Jan. 10.13	80	282	80	118 ± 26	ACS/WFC	F475W	360	24.48 ± 0.04
2012 Jan. 10.13	80	282	80	118 ± 26	ACS/WFC	F475W	360	24.43 ± 0.04
2012 Jan. 10.14	80	282	80	118 ± 26	ACS/WFC	F475W	470	24.51 ± 0.03
2010 Aug. 07.45	248	104	68 ± 26	104	ACS/WFC	F814W	350	$23.87 \pm 0.05^{\dagger}$
2010 Aug. 07.46	248	104	68 ± 26	104	ACS/WFC	F814W	700	$23.80 \pm 0.03^{\dagger}$
2010 Aug. 07.47	248	104	68 ± 26	104	ACS/WFC	F814W	455	$23.83 \pm 0.04^{\dagger}$
2012 Jan. 10.02	80	282	80	118 ± 26	ACS/WFC	F814W	350	23.98 ± 0.05
2012 Jan. 10.05	80	282	80	118 ± 26	ACS/WFC	F814W	800	23.97 ± 0.04
2012 Jan. 10.06	80	282	80	118 ± 26	ACS/WFC	F814W	550	23.99 ± 0.04
2011 Jan. 25.27	67	270	67	115 ± 26	WFC3/IR	F110W	800	$24.19 \pm 0.05^{\dagger}$
2011 Aug. 31.58	285	52	103 ± 26	52	WFC3/IR	F110W	700	23.71 ± 0.03
2011 Jan. 25.26	67	270	67	115 ± 26	WFC3/IR	F160W	400	$24.1 \pm 0.2^{\dagger}$
2011 Jan. 25.28	67	270	67	115 ± 26	WFC3/IR	F160W	400	$24.0 \pm 0.2^{\dagger}$
2011 Jan. 25.29	67	270	67	115 ± 26	WFC3/IR	F160W	400	$23.9 \pm 0.2^{\dagger}$
2011 Jan. 25.29	67	270	67	115 ± 26	WFC3/IR	F160W	500	$24.2 \pm 0.2^{\dagger}$
2011 Aug. 31.57	285	52	103 ± 26	52	WFC3/IR	F160W	400	23.5 ± 0.1
2011 Aug. 31.59	285	52	103 ± 26	52	WFC3/IR	F160W	400	23.5 ± 0.1
2011 Aug. 31.59	285	52	103 ± 26	52	WFC3/IR	F160W	400	23.4 ± 0.1
2011 Aug. 31.60	285	52	103 ± 26	52	WFC3/IR	F160W	400	23.3 ± 0.1

[†]Data derived directly from [Williams et al. \(2014a\)](#); the remainder have been provided directly by the PHAT collaboration.

3.2. A ‘folded’ eruption cycle

The *HST* data covering the quiescent system are admittedly sparse and spread across multiple eruption cycles. However, under the assumption of essentially identical eruptions ([Schaefer 2010](#), [DHB16](#)), M31N 2008-12a appears to take ~ 70 days to return to quiescence, i.e., to reach a minimum flux following an eruption. From this point, the luminosity of the system appears to increase in the lead-up to the next eruption, consistent with the findings of [Henze et al. \(2017a\)](#).

The RN RS Oph is perhaps the best studied Galactic nova both during eruption and at quiescence (see [Evans et al. 2008](#), and references therein). Following the 2006

eruption of RS Oph the system was observed to decline to an optical minimum before the flux began to systematically increase. The increase in flux was more prominent in bluer bands ([Darnley et al. 2008](#)), and coincided with the resumption of optical flickering ([Worters et al. 2007](#)). These observations were proposed to indicate the re-establishment of accretion post-eruption – following the destruction or severe disruption of that disk.

By mapping the quiescent PHAT data onto the template light curves we can combine these multi-color data into two distinct quiescent epochs, based on their approximate phase in the eruption cycle. The first (the red points in Figure 1) ~ 75 days post eruption represents the approximate minimum luminosity state, the second

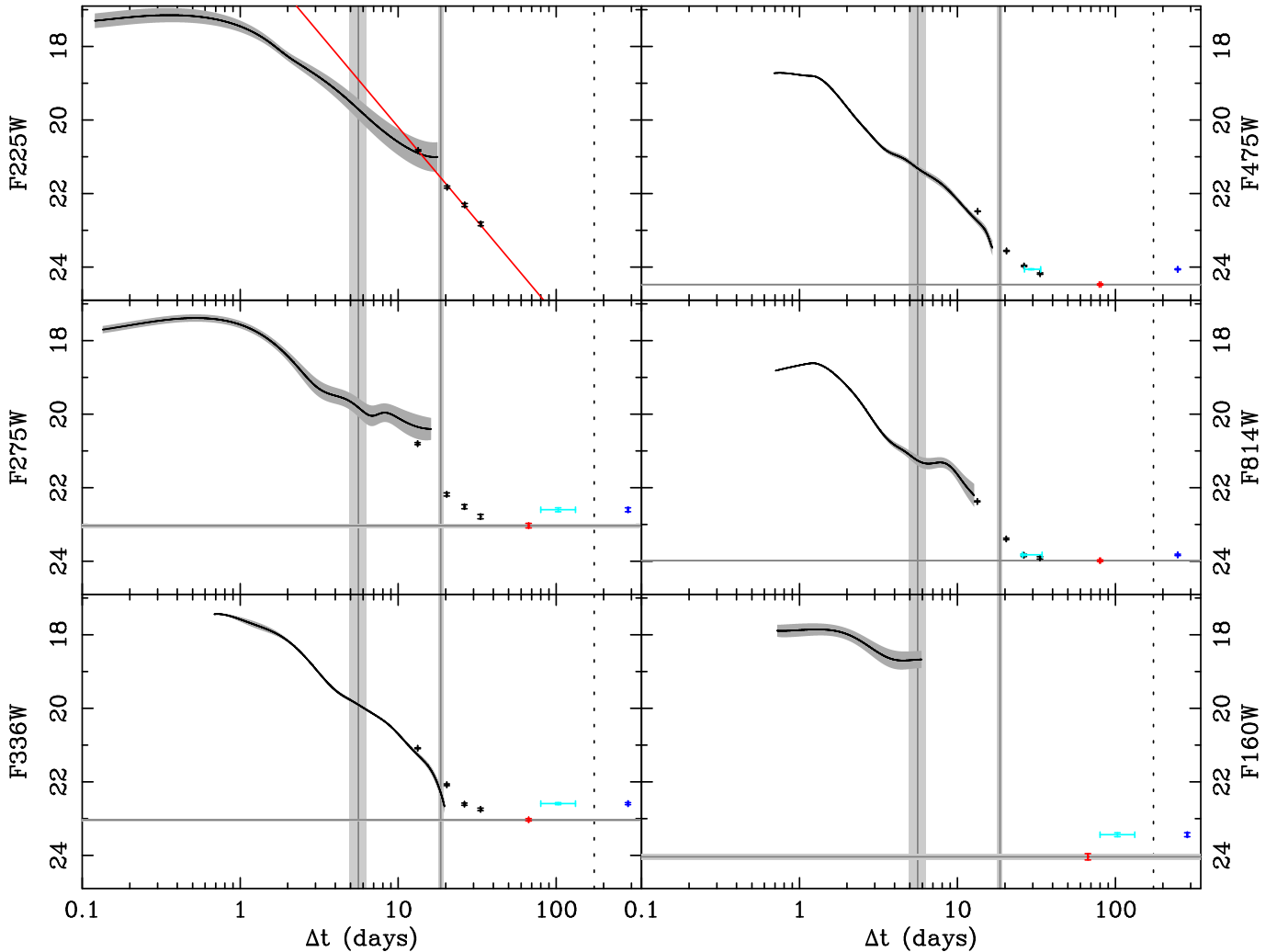


Figure 1. Near-ultraviolet through optical *HST* WFC3/UVIS photometry of the 2015 eruption of M31N 2008-12a (black data points). The time axes runs from 0.1 days post-eruption up to 347 days – the mean observed recurrence period. The vertical dashed line indicates the proposed 147 day recurrence period (HDK15). The horizontal lines, where shown, indicate the faintest detection of the two PHAT epochs – assumed to be the quiescence level. The vertical gray lines indicate the turn on and turn off times of the SSS from the 2015 eruption (the shaded areas their associated uncertainties). The solid black lines show combined and smoothed (Bézier curve) photometry from the 2013, 2014, and 2015 eruptions of M31N 2008-12a, with the surrounding shaded area indicating the 1σ uncertainty; these smoothed lines are provided for illustrative and contextual purposes only, the *HST* and ground-based/*Swift* data are taken through similar, but different, filters. The solid red line in the F225W shows the best fit power-law of index -2.04 ± 0.16 . The red data points indicate the faintest archival PHAT photometry following a detected eruption, the dark-blue points indicate additional PHAT photometry (assuming a year-long cycle), the light-blue points are the same PHAT photometry points extrapolated to a predicted missed eruption (based on a six month recurrence period).

(dark-blue points) ~ 270 days post eruption shows a state of increased flux. We again note that the lighter-blue data points in Figure 1 indicate one possible realization of a six month recurrence period; a realization that is ruled out by PTF data (see Section 3.1).

3.3. Spectral Energy Distribution

Optical and NUV photometric observations of the 2014 eruption of M31N 2008-12a indicated that, due to the low ejected mass, the unusually low maximum radius of the expanding pseudo-photosphere resulted in emission peaking in the UV (DHS15). For all other well observed novae, this peak occurs at visible wavelengths.

DHB16 presented a more comprehensive series of SEDs following the NIR (*H*-band) through NUV (*Swift* uvw1) decline of the 2015 eruption of M31N 2008-12a spanning $t \sim 1 - 10$ days post-eruption. In Figure 2 we reproduce the SED evolution plot from DHB16 and include the *HST* WFC3/UVIS photometry from the 2015 eruption of M31N 2008-12a ($t \sim 13, 20, 26,$ and 33 days post-eruption). We also include the updated quiescent photometry from archival *HST* observations. In Section 4 we will use these new data in conjunction with model accretion disks to constrain the mass accretion rates. The nature of the quiescent system is explored in Section 5.2.

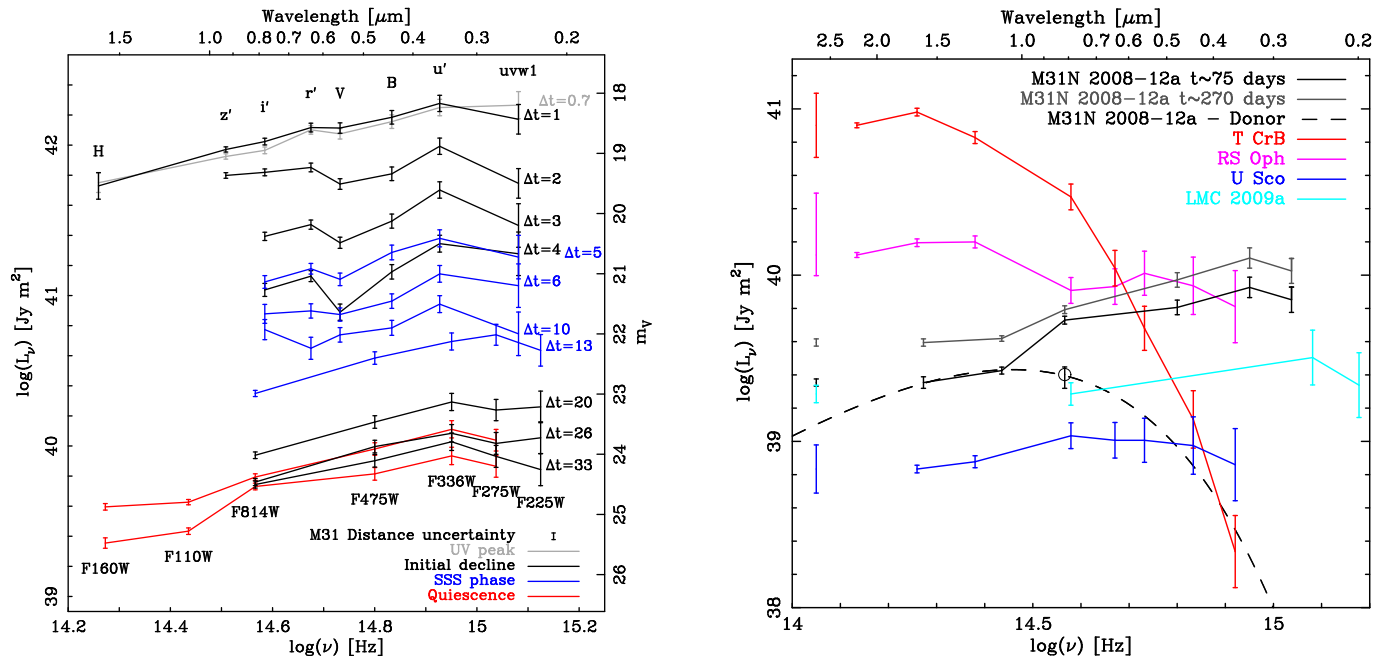


Figure 2. Distance and extinction ($E_{B-V} = 0.1$; [DHG17](#)) corrected SEDs showing **Left:** the evolving SED of the 2015 eruption (blue points indicate epochs when the SSS emission was visible, red points indicate the archival photometry), and **Right:** the quiescent M31N 2008-12a compared to the quiescent RNe RS Oph, T CrB, U Sco, and LMC 2009a. The black data shows M31N 2008-12a at minimum (~ 75 days post-eruption), and the gray data at an elevated state ($t \simeq 270$ d). Throughout, units chosen to allow comparison with similar plots in [Schaefer \(2010\)](#), see their Fig. 71) and [DWB14](#) (see their Fig. 4). The central wavelength locations of the Johnson-Cousins, Sloan, *HST*, and *Swift* filters are shown to assist the reader, see the Keys for line identifications. For each system, the photometric uncertainties are relatively small, indicated error bars are dominated by extinction uncertainties, the isolated error bar to the left of each SED indicates the systematic distance uncertainty. The single circled black data point indicates the excess donor flux once the accretion disk model (see Section 4.5) has been subtracted. The dashed black line indicates a power law-fit to the M31N 2008-12a donor SED (see Section 5.2).

4. MODELING THE ACCRETION DISK

In this Section, and subsequently in Section 5, we discuss in detail our models and interpretation of the accretion disk in M31N 2008-12a. Here, for clarity, we formally define some of the accretion rate terminology that we employ.

The models, discussed below, generate the disk mass accretion rate (\dot{M}), whereas the existing M31N 2008-12a eruption models of [Kato et al. \(2014, 2015, 2016, 2017b,a\)](#) are concerned with the WD mass accretion rate (\dot{M}_{acc}); the amount of material that falls onto the WD surface itself.

In this work we will also consider mass loss from the disk via a disk wind (\dot{M}_{wind}), and mass loss from any outflows from the WD or the disk-WD boundary layer (\dot{M}_{bl}); such that:

$$\dot{M}_{\text{acc}} = \dot{M} - \dot{M}_{\text{wind}} - \dot{M}_{\text{bl}},$$

for most novae \dot{M} is low, therefore it is expected that $\dot{M}_{\text{wind}} - \dot{M}_{\text{bl}}$ are small, and as such $\dot{M}_{\text{acc}} \simeq \dot{M}$.

4.1. Disk Models

The *tlusty*, *synspec*, *rotin*, and *disksyn* suite of codes ([Hubeny 1988](#); [Hubeny et al. 1994](#); [Hubeny & Lanz 1995](#)) are employed to generate synthetic spectra of stellar atmospheres and disks. These include the treatment of hydrogen quasi-molecular satellite lines (low

temperature) and NLTE approximation (high temperature). *synspec* generates continuum spectra with absorption lines. In the present work we do not generate emission lines (see, e.g., [Puebla et al. 2007](#), for a physical description of emission line profiles from disks in CVs). For disk spectra we assume Solar abundances and for stellar spectra we vary the abundances as required.

The *tlusty* code is first run to generate one-dimensional (vertical) stellar atmosphere structures for a given surface gravity, effective temperature and surface composition of the star. H and He are treated explicitly, whereas C, N, and O are treated implicitly ([Hubeny & Lanz 1995](#)).

The *synspec* code takes the *tlusty* stellar atmosphere model as an input, and generates a synthetic stellar spectrum over a given wavelength range from below 900 Å and into the optical. The *synspec* code then derives the detailed radiation and flux distribution of the continuum and lines, to generate the output spectrum ([Hubeny & Lanz 1995](#)). *synspec* has its own chemical abundances input to generate lines for the chosen species. For temperatures > 35000 K the approximate NLTE line treatment is turned on in *synspec*.

Rotational and instrumental broadening, as well as limb darkening (see [Wade & Hubeny 1998](#)), are then reproduced using the *rotin* routine. In this manner, we generated WD synthetic spectra covering a wide range of

temperatures and gravities, all with Solar composition.

The disk spectra are generated by dividing the disk into annuli, with radius r_i and effective surface temperature $T(r_i)$ obtained from the standard disk model for a given WD mass M_{WD} and mass loss rate .

Utilizing input parameters of the disk mass accretion rate (\dot{M}), M_{WD} , the radius of the WD R_{WD} , the inner radius of the disk R_0 , and the outer radius of the disk R_{disk} , `tlusty` generates a one-dimensional vertical structure for each disk annulus (Wade & Hubeny 1998).

In the standard disk model, the radius R_0 is the boundary at which the ‘no shear’ condition is imposed; $d\Omega/dR = 0$ (Pringle 1977). Consequently, the assumed value of R_0 affects the entire solution (not just the boundary) and the temperature profile of the disk.

For moderate disk mass accretion rates, $\dot{M} \sim 10^{-8} M_{\odot} \text{yr}^{-1}$, the boundary layer between the disk and the WD, that region where the angular velocity in the disk decreases from its Keplerian value Ω_K to match the more slowly rotating WD surface Ω_* , is very small ($\sim 0.01 R_{\text{WD}}$) and one can therefore assume $R_0 = R_{\text{WD}}$ (Pringle 1977).

In our present modeling, R_0 is allowed to be larger than the radius of the WD, $R_0 > R_{\text{WD}}$, to accommodate a larger boundary layer (see Godon et al. 2017, for a description of this modified disk model). As \dot{M} increases, the boundary layer becomes larger (Popham & Narayan 1995). As \dot{M} reaches the Eddington accretion limit, the size of the boundary layer rises to the order of the radius of the WD ($R_0 \sim R_{\text{WD}}$; Godon 1997).

Given the large quiescent luminosity and high ejection velocities, DHB16 and DHG17 proposed that the system inclination must be low. Although high inclination systems are not formally ruled out, we note that the large observed disk luminosity would require a significant increase in any derived \dot{M} as the assumed inclination increases.

To model the M31N 2008-12a disk, we assume $M_{\text{WD}} = 1.37 M_{\odot}$ and $R_{\text{WD}} = 2000 \text{ km}$, yielding an Eddington limit $\dot{M}_{\text{Edd}} = 4 \times 10^{-6} M_{\odot} \text{yr}^{-1}$. We generate a grid of disk models for inclinations $i = 10^\circ$, 20° , and 30° . These models are computed for fixed values of \dot{M} in logarithmic intervals of 0.5. The true value of \dot{M} is computed by fitting the observed data by interpolating between the computed values of \dot{M} . For R_0 , we choose:

$$R_0 = \begin{cases} 1.0R_{\text{WD}}, & \dot{M} \leq 10^{-7} M_{\odot} \text{yr}^{-1} \\ 1.1R_{\text{WD}}, & \dot{M} = 10^{-6.5} \text{ and } 10^{-6} M_{\odot} \text{yr}^{-1} \\ 1.5R_{\text{WD}}, & \dot{M} = 10^{-5.5} M_{\odot} \text{yr}^{-1} \\ 2.0R_{\text{WD}}, & \dot{M} = 10^{-5} M_{\odot} \text{yr}^{-1}. \end{cases}$$

For M31N 2008-12a, we use Kurucz stellar spectra of appropriate temperature and surface gravity to extend the outer disk to a radius where $3500 \lesssim T < 10000 \text{ K}$. We also consider disks that are truncated in the outer region as discussed in the results section.

`synspec` uses the `tlusty` results for each disk annulus to generate synthetic spectra. These are integrated into a disk spectrum using `disksyn`, which includes effects of Keplerian broadening, inclination, and limb darkening (Wade & Hubeny 1998)

4.2. Results

Here, we adopt an inclination of 20° , distance of 770 kpc, and reddening $E_{B-V} = 0.1$ (DHG17). In Section 4.4, we take the effects of a different inclination (10° or 30°), an error of $\sim 20 \text{ kpc}$ on the distance, and a reddening error of 0.03 (DHG17) into consideration and assess how these affect the final results. Since the error bars on the data points are themselves rather small (at most 5%), they too are considered at the end of this section. A low inclination system is assumed due to the large UV flux at quiescence, and the large observed ejecta velocities (DHB16, DHG17).

The flux data points from the different epochs were obtained through filters covering given wavelength bands, and as such they represent an average continuum flux level in these regions of the spectrum, possibly also including some prominent lines. One data point (F475W; 4773.7 Å) includes H β (which would be in absorption unless there is a disk wind, which we do not model here). We therefore do not expect the data points, at any epoch, to line up nicely with the continuum of the optically thick standard disk model, *but rather we use our modeling simply to assess the order of magnitude of the mass accretion rate.*

4.2.1. Quiescence

We start by modeling the inter-eruption data at $t \simeq 75 \text{ d}$, as here the flux is at a minimum and we expect the disk to dominate the optical–NUV emission, with negligible contribution from the waning eruption. The modeling at this epoch is then applied to, and adjusted as necessary, the other five epochs.

For the disk models to simply provide sufficient flux to match the observations at the distance of M31, the disk mass accretion rate⁴, \dot{M} , is required to be large, $\gtrsim 10^{-6} M_{\odot} \text{yr}^{-1}$, and therefore not far away from \dot{M}_{Edd} . Such models generate a prominent Balmer discrepancy at $\sim 4000 \text{ Å}$, which is not apparent in the quiescent SEDs (also see the late-time spectrum in Figure 3). However, many CVs *accreting at a high rate* do not exhibit strong Balmer discontinuities (Matthews et al. 2015). We began by fitting the synthetic spectra longward of the Balmer discontinuity to just the F475W photometry (4773.7 Å), which requires $\dot{M} = 1.28 \times 10^{-6} M_{\odot} \text{yr}^{-1}$, see Figure 4e (solid black line). This model has an outer disk radius extending to $1320 R_{\text{WD}} \simeq 3.8 R_{\odot}$, where the temperature falls to 6000 K, but the model is clearly deficient in flux at wavelengths shorter than the Balmer discontinuity.

⁴ We again note that this may be formally different from the WD mass accretion rate, $M_{\text{acc}} \leq \dot{M}$.

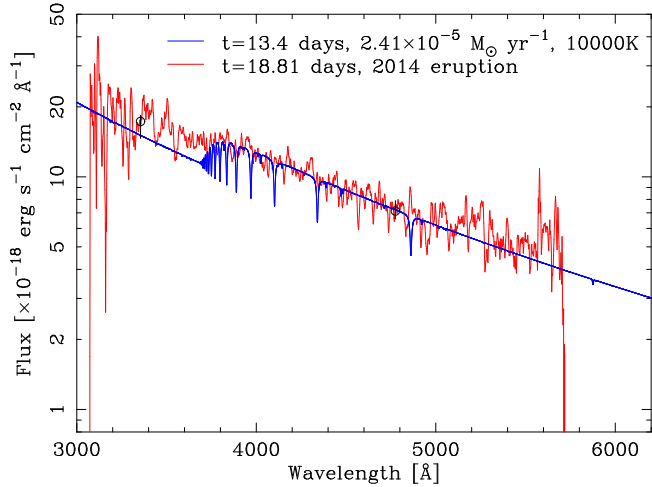


Figure 3. Direct comparison between the $t = 13.4$ day accretion disk model of the 2015 eruption and the dereddened $t = 18.81$ day Keck spectrum of the 2014 eruption. The flux of the Keck spectrum has been increased by $8.85 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$. The spectrum and the model are in good agreement above 4000 \AA , however typical accretion disk absorption lines and the Balmer discontinuity is not present the spectrum, which contains $H\beta$ in emission.

Matthews et al. (2015) proposed that the absence or reduction of the Balmer discontinuity observed in some CVs is due to continuum emission from a disk wind. We note that disk models whose outer radii are truncated also produce spectra with decreased Balmer discontinuities. In Section 5.3 we briefly discuss possible physical explanations for disk truncation. Truncated disk models provide a slightly lower continuum flux level for the same \dot{M} .

If the disk is truncated at $R = 880 R_{\text{wd}} (\approx 2.53 R_{\odot})$, the model fits the first data point (shortest wavelength). This effectively removes regions of the disk cooler than 8000 K . Outer disk truncation results in a reduced flux (for a fixed \dot{M}), therefore the mass accretion rate of this model must rise to $1.51 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, see Figure 4e (red line). To fit the second data point, nearest to the Balmer edge, we further truncate the disk to $750 R_{\text{WD}} (\approx 2.08 R_{\odot})$. Such a disk has $\dot{M} = 3.35 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ and the temperature in the outer disk reaches 12000 K . This model, however, overshoots the first data point. The model is shown in Figure 4e (blue line), and the derived accretion rates are tabulated in Tables 5. We note that all three models underestimate the F814W flux. The excess flux here may be contributed by the donor (see Sections 4.5 and 5.2). We also note that the differing photometric points at both quiescent epochs were taken at different times, therefore any fundamental variability at quiescence could be imprinted on these data.

Next, we turn to the second quiescence epoch, ~ 270 d post-eruption, and (assuming an annual cycle) ~ 70 d pre-eruption. These data are similar to those at $t \simeq 75$ d (see Figure 4e and 4f), but the flux is higher. Consequently, we follow the same modeling procedure. In Fig-

Table 5. Computed accretion rates for M31N 2008-12a during the final decline of the 2015 eruption and at quiescence.

Epoch (days)	\dot{M}	$\dot{M}_{\text{wind}} \simeq \dot{M}_{\text{acc}}$ ($\times 10^{-6} M_{\odot} \text{ yr}^{-1}$)	$\dot{M}_{\text{truncated}}$
13.4	17.2	8.60	24.1
20.4	4.17	2.86	5.59
26.3	3.60	1.80	4.30
33.3	1.55	0.77	1.86
~ 75	1.28	0.64	3.35
~ 270	2.84	1.42	4.90

NOTE— \dot{M} refers to the disk mass accretion rates as directly calculated by the accretion disk models, $\dot{M}_{\text{truncated}}$ refers to the \dot{M} upper limits imposed by heavily truncated disk, \dot{M}_{wind} is the expected maximum mass loss from the disk via a disk wind, with the remaining amount \dot{M}_{acc} being accreted onto the WD (i.e., $\dot{M}_{\text{acc}} = \dot{M} - \dot{M}_{\text{wind}}$, or $\dot{M}_{\text{acc, min}} \simeq \dot{M}/2$; also see the later discussion about possible outflows).

ure 4f we present three models with the outer disk truncated, and $(2.84 \leq \dot{M} \leq 4.9) \times 10^{-6} M_{\odot} \text{ yr}^{-1}$. Again, truncating the cooler outer disk reduces the “jump” of the Balmer edge. From $t \simeq 75$ d to $t \simeq 270$ d \dot{M} has increased by a factor of ~ 1.5 to 2.2.

In all the models presented here, we found that the inclusion of a hot WD did not contribute any significant flux due to the small surface area of the massive WD and to the very large area of the very hot disk. It is also probable that at high \dot{M} the inner disk is swollen and masks the WD.

4.2.2. The decline

We next consider the evolution during the late decline of the 2015 eruption. We model these in reverse, as the complexity of the emission is expected to increase closer to the eruption itself.

At $t = 33.3$ d, the flux from M31N 2008-12a lies approximately midway between that at quiescence ($t \simeq 75$ and $\simeq 270$ d, see left panel of Figure 2), and we find that a standard (non-truncated) disk model with $\dot{M} = 1.55 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ provides a reasonable fit to the data. This model has an outer region extending to where the temperature reaches 6000 K , extending the outer region to 3500 K does not improve the fit to the data points. The fit is presented in Figure 4d. There is a slight flux excess at $\sim 3350 \text{ \AA}$, but as shorter wavelengths are consistent with the model, a truncated disk model does not provide a better fit to the data. Again, there is a flux excess at $\sim 8000 \text{ \AA}$.

Turning to $t = 26.3$ d, the data points are in better agreement with the presence of a weak Balmer edge (see Figure 4c). We fit a disk model while varying the outer truncation radius, and find that the best fit is obtained for $\dot{M} \sim 3.6 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ with the outer disk truncated at 9500 K ($R_{\text{disk}} = 1050 R_{\text{WD}}, \sim 3 R_{\odot}$).

A week earlier, $t = 20.4$ d, we find $\dot{M} = 4.17 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ and $5.59 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ for disks trun-

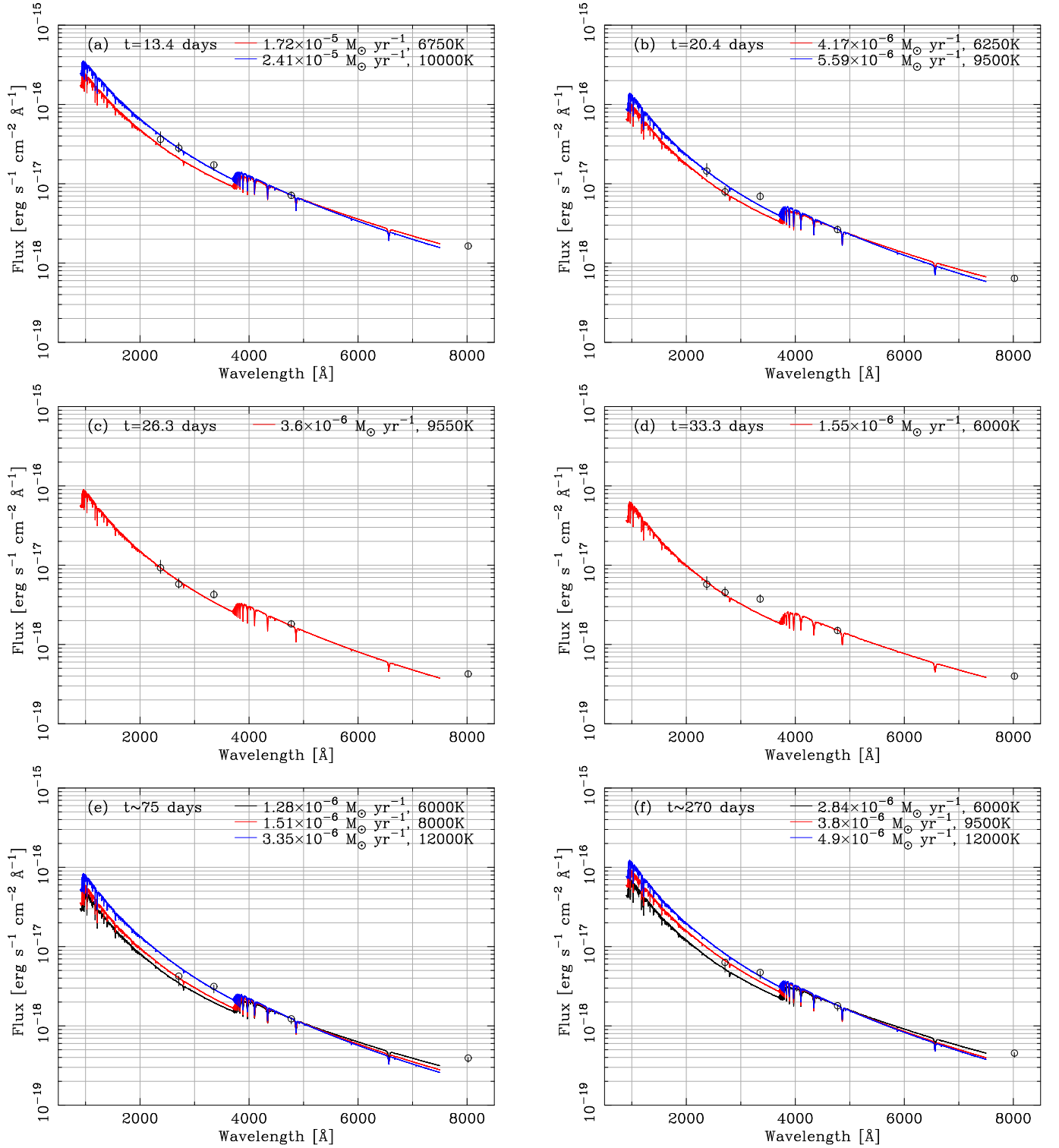


Figure 4. Accretion disk modeling of the declining M31N 2008-12a.. Each sub-plot indicates the epoch post-eruption, and the black data points show the corresponding broad-band *HST* photometry and associated uncertainty. For model fit details, refer to the text and the keys in each sub-plot, which indicate the mass accretion rate \dot{M} and the temperature at the truncated outer edge of the disk. Sub-plots (a)–(d) cover the late decline of the 2015 eruption, (e) and (f) are the two reconstructed inter-eruption, or quiescent, epochs.

cated at 6250 K and 9500 K, respectively. Neither model reproduces the NUV flux well, possibly an indication of a contribution from an additional source or lines. These two disk models are presented in Figure 4b.

4.2.3. The Super-Soft X-ray phase.

Finally we turn to the observations at $t = 13.4$ d, during the SSS phase of the 2015 eruption. As is evident from Figure 4a, not only is this the epoch with the highest flux, but the data exhibit a rather smooth “continuum” – almost a straight line on this logarithmic scale. This is further illustrated by the Keck spectrum taken 18.81 d after the 2014 eruption, which is directly compared to the 2015 $t = 13.4$ d data in Figure 3.

There is no indication of the presence of the Balmer edge from the 2015 data, which is confirmed by the Keck 2014 spectrum. A disk model truncated at 20000 K produces a smooth continuum without a Balmer edge, but the continuum slope is much steeper than inferred from the data. Therefore, we fit the data with disk models that have various degrees of truncation. We find $\dot{M} = 1.72 \times 10^{-5} M_{\odot} \text{ yr}^{-1}$ for a disk truncated at 6750 K, and $2.41 \times 10^{-5} M_{\odot} \text{ yr}^{-1}$ for a disk truncated at 10000 K; see Figure 4a. As with the other epochs, these models cannot fit all the data points simultaneously. We note that such an \dot{M} is above \dot{M}_{Edd} .

4.3. Disk winds

As mentioned earlier, Matthews et al. (2015) proposed that the absence or reduction of the Balmer edge in optical spectra of some CVs is due to the existence of powerful accretion disk winds. Indeed, Matthews et al. (2015) show that a standard disk wind model is successful in reproducing the weak Balmer absorption edge at all inclinations, but particularly for CV systems viewed at high inclination. They further suggest that winds can dominate the continuum emission from CVs. Their modeling shows that the inclusion of the disk wind produces a much weaker Balmer edge, a shallower continuum slope, and the flux level increases due to the contribution of the wind to the disk continuum.

Consequently, compared to the model fit in Matthews et al. (2015), our optically thick non-truncated standard disk models provide an *upper limit* to the mass accretion rate onto the WD (\dot{M}_{acc}), as they produce less flux at the same accretion rate. The discrepancy between the wind disk model and the standard disk model is minimal near the upper edge of the Balmer jump ($\sim 4000 - 5000 \text{ \AA}$) and appears to reach a maximum of about a factor of two in \dot{M} . Therefore, if we assume that disk wind emission has to be taken into account, we have to reduce the mass accretion rates obtained from our non-truncated disk model fits by a maximum of 50% (i.e., $\dot{M}_{\text{acc}} \simeq \dot{M}_{\text{wind}} \simeq 0.5\dot{M}$). Subsequently, the implied mass accretion rates in the presence of a disk wind are shown in Table 5.

4.4. Model uncertainties

Finally, we compute the relative uncertainties introduced from the errors on the reddening, inclination, dis-

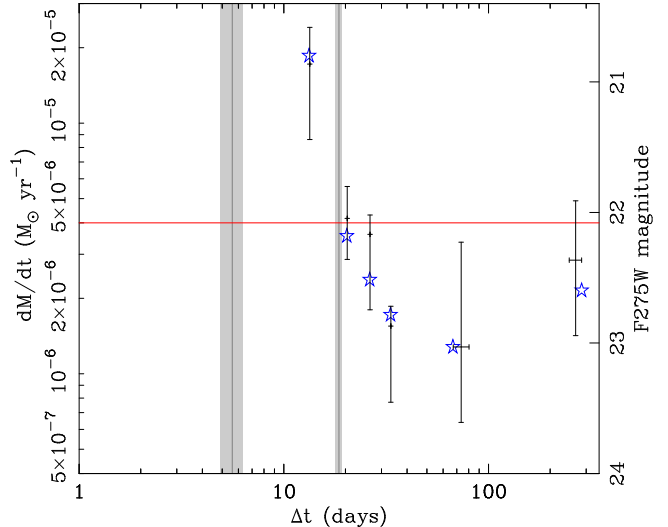


Figure 5. Evolution of the disk mass accretion rate (\dot{M}) of the M31N 2008-12a accretion disk with time within an eruption cycle (black points). The upper limits refer to $\dot{M}_{\text{truncated}}$ and the lower limits $\dot{M}_{\text{wind}} (\simeq \dot{M}_{\text{acc}})$, see Table 5. The *HST* F275W photometry is also plotted for comparison (blue data points), and to indicate that it traces the accretion rate well. The Eddington accretion limit of a $1.37 M_{\odot}$ WD is indicated by the red horizontal line. The vertical gray lines indicate the turn on and turn off times of the SSS from the 2015 eruption (the shaded areas their associated uncertainties).

tance, and fluxes. For this purpose we consider the data for $t = 33.3$ d, with $\dot{M} = 1.55 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$.

DHG17 computed that the reddening toward M31N 2008-12a is $E_{B-V} = 0.10 \pm 0.03$. De-reddening the $t = 33.3$ d data using $E_{B-V} = 0.07$ and $E_{B-V} = 0.13$ gives $\dot{M} = 1.33 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ and $\dot{M} = 1.90 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, respectively. That is, the disk mass accretion rate becomes $\dot{M} = 1.55^{+0.35}_{-0.22} \times 10^{-6} M_{\odot} \text{ yr}^{-1}$.

Similarly we compute the errors for an inclination of $i = 20^{\circ} \pm 10^{\circ}$, distance of 770 ± 20 kpc, and a maximum error of 5% in the fluxes (see Table 4). Assuming \dot{M} varies linearly with small changes in E_{B-V} , i , d , and the fluxes, by quadrature we obtain $\dot{M} = 1.55^{+0.39}_{-0.25} \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, errors of +25% and -16%. These errors are much smaller than the systematics introduced from the use of a truncated disk model (a factor of ~ 2 in \dot{M}) or when comparing our standard disk models to the Matthews et al. (2015) disk wind models (a factor of ~ 0.5). Namely, in fitting the data from $t = 13.4$ d, we obtained $\dot{M} = 1.28 \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, but \dot{M} could be about twice this value if we consider the truncated disk models, or could be about half this value if we consider the possibility of a disk wind continuum.

Taking this into account, we reproduce the computed accretion rates in Table 5, and plot them as a function of time in Figure 5.

4.5. Donor flux excess

As can be seen in Figure 4, there is a flux excess, above the disk models, in the F814W band at all epochs. As the only expected ‘red’ component in the system, this

excess flux is probably from the donor. To examine this effect, we extended the $t \simeq 75$ d and $t \simeq 270$ d accretion disk models to longer wavelengths by fitting a power law to the model spectra, redward of the Balmer limit. This produced a good fit to the model spectra, and enabled us to determine a F814W flux excess at quiescence of $(1.67 \pm 0.18) \times 10^{-19} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$, which corresponds to an apparent magnitude of $m_{\text{F814W}} = 24.8^{+0.2}_{-0.1}$. Further extrapolation of the accretion disk model confirms that any disk contribution in the NIR F110W and F160W filters is negligible.

5. DISCUSSION

5.1. *The accretion disk*

In Section 4 we described the comparison between the *HST* photometry of the final decline of the 2015 eruption, and quiescent observations, of M31N 2008-12a to models of accretion disks around $1.37 M_{\odot}$ WDs. We again state that the ideal datasets for such work would be spectroscopy extending into (and even beyond) the FUV. However, for CVs at the distance of M31 such observations are not yet feasible. Therefore, the *HST* visible and NUV photometry described in this paper currently provide the best, *and only*, data with which to explore the accretion disk in M31N 2008-12a.

The one thing that is immediately clear is the very large luminosity of the M31N 2008-12a accretion disk at quiescence. By necessity, modeling of such a high luminosity disk requires a large disk mass accretion rate (\dot{M}). The results of the modeling show that the broad-band photometric SED of M31N 2008-12a from the epoch of the first post-eruption *HST* imaging, and during quiescence, is consistent with the expected form of an accretion disk. As the first *HST* epoch occurs only 13 days after the 2015 eruption, indeed before the SSS is extinguished, this is evidence that the accretion disk may survive eruptions of M31N 2008-12a.

The basic form of the SED, from the optical to NUV, remains consistent from $t = 13$ d to quiescence, adding further weight to the survival of the disk. Observationally, we first see this disk beginning to dominate the optical/NUV emission about two weeks post-eruption, and possibly as early as $t = 4$ days⁵. The disk luminosity decreases to a minimum just ~ 75 days post eruption, before building again toward the next eruption – presumably as the accretion disk increases in mass. Our working model is that the disk, once struck by the nova ejecta, is initially shocked and heated, but survives largely intact. Further, irradiation from the SSS may begin to affect the disk from as early as $t = 4$ days. These effects cause the disk to begin losing mass at a large rate through a disk wind (with the disk accretion rate at $\dot{M} \sim 2 \times 10^{-5} M_{\odot} \text{ yr}^{-1}$), as is discussed below, some

of this mass may be accreted directly onto the WD. As the surviving disk then cools and relaxes its luminosity decreases until reaching a minimum after ~ 75 days ($\dot{M} \sim 10^{-6} M_{\odot} \text{ yr}^{-1}$). During the next ~ 200 days of quiescence, mass loss from the donor allows the disk to rebuild any matter lost (through the eruption and disk wind), in the run up to the next eruption.

But there is a potential problem, not necessarily with the picture outlined above, but with the mass accretion rates derived from the models, which do not include disk winds. Namely, we computed values of \dot{M} representing the disk mass accretion rate, *not* the accretion rate onto the WD (\dot{M}_{acc}). Up to half of \dot{M} might be lost through a disk wind (Matthews et al. 2015), reducing the effective accretion onto the WD by up to 50%. Our disk models imply that \dot{M} is close to, or even exceeds, \dot{M}_{Edd} throughout the entire eruption cycle, a state where a significant radiation pressure driven disk wind may be expected to be present.

A number of authors have investigated the WD mass – WD accretion rate (\dot{M}_{acc}) phase space, and they arrive at two broad but differing conclusions. The first is that, other than \dot{M}_{Edd} itself, there is (for a given WD mass) no upper limit on the mass accretion rate, and that nova eruptions will occur at any \dot{M}_{acc} (see, e.g., Starrfield 2016). Or alternatively, that there is a clear upper limit to \dot{M}_{acc} (see, e.g., Fujimoto 1982; Nomoto 1982), beyond which nova eruptions cease, with the WD entering a phase of steady state nuclear burning (the persistent SSS). At even higher accretion rates, these models predict that optically thick winds (from the WD) are generated. In recent years, it has been proposed that one important difference between these two scenarios is how mass accretion is treated during a nova eruption (Hachisu et al. 2016), with the former assuming it ceases, the latter assuming it continues. With M31N 2008-12a showing both signs of a surviving disk, therefore continuing accretion, and an elevated \dot{M} , it may be an important system in addressing this long-standing issue.

Discussion of the merits of these two differing pictures is clearly beyond the scope of this paper. But we note, of course, that the former (with no upper limit) poses no clear obstacle to our derived accretion rates. Turning to the latter, we note that the work of Kato et al. (2014, 2015, 2016, 2017b,a), employing such a formulation, has already successfully modeled many observational aspects of the M31N 2008-12a eruptions, while assuming a *constant* $\dot{M}_{\text{acc}} = 1.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ – a factor of four lower than the \dot{M}_{acc} lower limit derived in this work (under the assumption of $\dot{M}_{\text{acc}} \simeq \dot{M}_{\text{wind}}$). As is shown graphically in Figure 6, accretion disks with $\dot{M} = 1.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ significantly under-predict the NUV flux of M31N 2008-12a at quiescence. The discrepancy is a factor of ~ 10 , even at the quiescence minimum of ~ 75 days post-eruption.

In Figure 7 we have recreated Figure 6 of Kato et al. (2014, M. Kato, priv. comm.), which shows the loci of equi-recurrence periods of novae in the WD mass – \dot{M}_{acc}

⁵ Could the optical/NUV light curve plateau presented in DHB16 be caused by the surviving disk being unveiled by the receding photosphere? A similar prediction was made for a number of Galactic RNe by Hachisu et al. (2008).

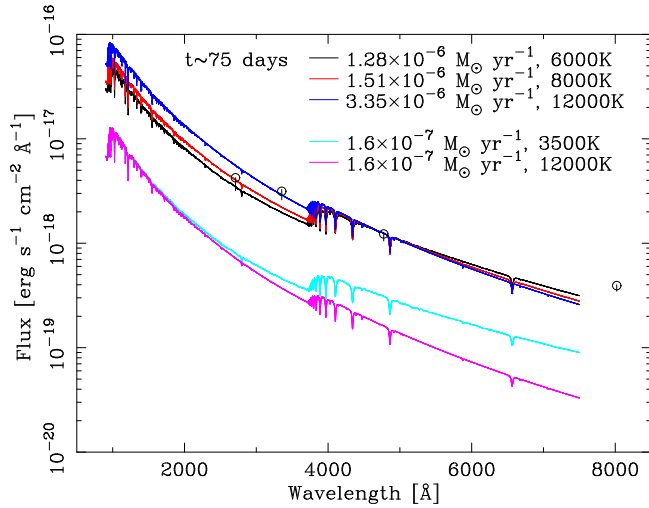


Figure 6. As Figure 4, comparing the best fit accretion disk models to the *HST* photometry ~ 75 days post-eruption – the minimum state. Also shown, in magenta and cyan, are our predicted accretion disk spectra based on an accretion rate of $\dot{M} = 1.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$, as required by Kato et al. (2014, 2015, 2016, 2017b). Disks with the Kato et al. accretion rate under-predict the quiescent flux of M31N 2008-12a by a factor of ~ 10 .

plane. In this plot, as discussed above, the regions of proposed steady burning and optically thick winds are shown. The position of M31N 2008-12a as computed by Kato et al. (2014) is indicated by the red star, and this lies clearly at the extremes of the phase-space permitted by these models. The disk mass accretion rates computed in this work are clearly at odds with the Kato et al. (2014) formulation, unless only a small proportion of the matter from the disk is accumulated on the WD surface. Given our computed mass loss rates, we would require *at least* 80% of \dot{M} to constitute a disk wind (\dot{M}_{wind}), to stop the system undergoing the proposed steady-state nuclear burning. However, such an elevated \dot{M}_{wind} seems unlikely for a sub-critical accretion disk (see, e.g., Poutanen et al. 2007).

Since our disk model indicates such a large \dot{M} , we must explore the validity of the standard α /Shakura & Sunyaev (1973) disk model as \dot{M} approaches \dot{M}_{Edd} , since the basic “geometrically thin” assumption breaks down, i.e. height/radius ≈ 1 . In this regime of $\dot{M} \approx \dot{M}_{\text{Edd}}$, the disk can be represented using the slim-disk equations (Abramowicz et al. 1988), where radial advection and radiation of energy is taken into account, and the flow can be partially supported by gas and radiation pressure. The departure from the standard disk model, however, is noticeable at mass accretion rates reaching $\dot{M} \sim 20\dot{M}_{\text{Edd}}$ (Abramowicz et al. 1988) as the heat trapped within the matter becomes important, and the luminosity increases more slowly than the accretion rate as the matter with its energy content is advected and radiated inward to the inner disk and onto the WD surface. Since the maximum mass accretion we compute in this work is $\dot{M} \approx 4\dot{M}_{\text{Edd}} < 20\dot{M}_{\text{Edd}}$, and the minimum is as low as $0.2\dot{M}_{\text{Edd}}$ during quiescence, our disk models are

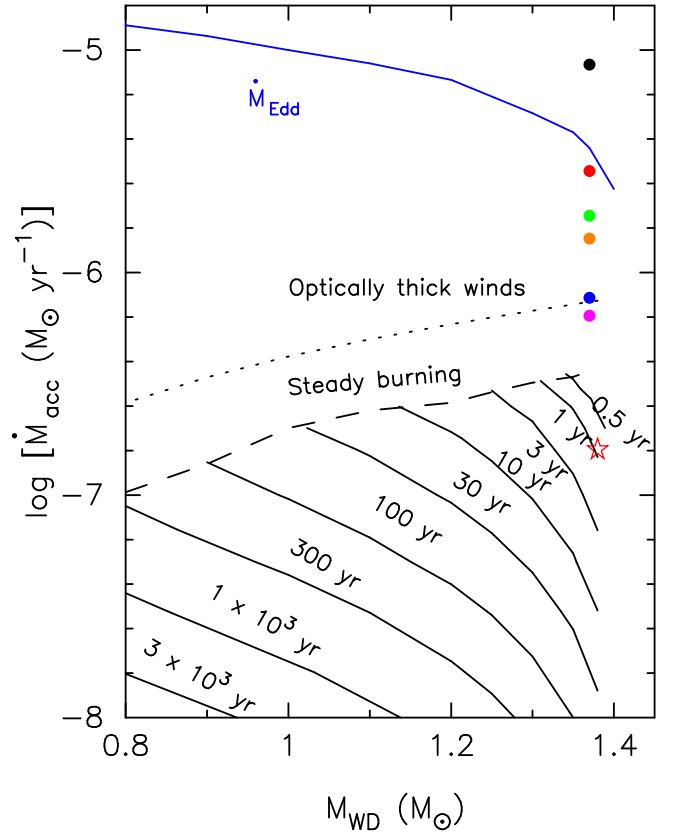


Figure 7. Recurrence period t_{rec} of novae on the $M_{\text{WD}} - \dot{M}_{\text{acc}}$ plane, based on Figure 6 of Kato et al. (2014). We plot the loci of the equi-recurrence periods of novae (black solid lines). Hydrogen burning is proposed to be stable in the region above the dashed line (\dot{M}_{stable}). In the region below \dot{M}_{stable} , H-shell burning is thermally unstable, and the WD experiences shell flashes (i.e., novae). Optically thick winds are accelerated in the region above the dotted line (\dot{M}_{cr}). We note that the Starrfield et al. (2016) interpretation of this plot does not require either the steady burning nor optically thick wind regions (see text for more details). The solid blue line indicates the Eddington limit (based on Nomoto 1982). The red star indicates the quiescent position of M31N 2008-12a, based on the modelling in Kato et al. (2015). The coloured points, at $M_{\text{WD}} = 1.37 M_{\odot}$ indicate the *lower limits* of the computed values of \dot{M} (i.e. assuming that $\dot{M}_{\text{acc}} = 0.5\dot{M}$) from (top–bottom) $t=13.4$ (black), 20.4 (red), 26.3 (green), ~ 270 (orange), 33.3 (blue), ~ 75 d (magenta) post-eruption (see $\dot{M}_{\text{wind}} \approx \dot{M}_{\text{acc}}$ values in Table 5).

probably not strongly affected by neglecting advection of energy.

Advection of energy is, however, more pronounced in the inner disk and can be expected to peak in the boundary layer between the WD and disk, since an additional $L_{\text{bl}} \approx L_{\text{acc}}/2$ is released in that region. This does not affect our disk models either, as the inner annuli in our models do not contribute significant flux at wavelengths $> 2000 \text{ \AA}$, because of their small surface area ($r < 6.5R_{\text{WD}}$) and elevated temperature ($> 3 \times 10^5 \text{ K}$) peaking in the EUV/soft X-ray regime ($\sim 10^5 \text{ K}$).

Having established that our disk models are valid, we furthermore consider the fate of the advected energy in

the inner disk/boundary layer. It has been shown that even at moderately large accretion rates ($\dot{M} \approx \dot{M}_{\text{Edd}}$), advection of energy becomes important in the boundary layer (Godon 1997; Popham 1997). As in advection dominated accretion flows (ADAFs; Narayan & Yi 1994, 1995) the inner disk and boundary layer will radiate significantly less than expected, and the advected energy will heat up the WD and drive a bipolar outflow (\dot{M}_{bl}) in addition to the disk-wind component. This will reduce the amount of material actually accreting onto the WD surface and could bring the WD accretion rate back toward the regime favoured by Kato et al. (2015, and others). The outflow of matter is possibly low in the outer disk and increases inward, where a strong wind forms a bipolar outflow.

For a number of decades some CVs have been suspected to have strong outflows with some systems even exhibiting ejecta, such as the ‘nova-like’ BZ Camelopardalis that is surrounded by a bow-shock nebula (Ellis et al. 1984). However, so far, one finds no collimated outflows (“jets”) in CVs (Hillwig et al. 2004) in spite of the fact that all other disk systems (from X-ray binaries to AGN) exhibit collimated outflows (Livio 1997). Is it possible that M31N 2008-12a is the exception to the *rule*, not just during eruption (see DHB16 and DHG17) but at quiescence?

If we assume a strong disk wind, then about half of the disk material is accreted on to the WD (at a rate of $\dot{M}_{\text{acc}} \sim 6.4 \times 10^{-7} M_{\odot} \text{yr}^{-1}$ at the apparent quiescent minimum), and the other half is deposited into the system (at the same rate, see Table 5 and Section 5.2) – the circumbinary environment. We note that the estimated red giant wind mass (total) in RS Oph at the time of eruption is $\sim 10^{-6} M_{\odot}$ (Vaytet et al. 2011), broadly consistent with the circumbinary contamination predicted by the M31N 2008-12a disk wind. Therefore, such a disk wind mass loss rate alone could be sufficient to account for the observed ejecta deceleration (DHB16) *without a requirement for a wind from the donor*.

The discussion of the accretion disk wouldn’t be complete without considering irradiation of the outer disk by the hot inner disk/boundary layer region. Disk irradiation is known to be important in low-mass X-ray binaries, where accretion occurs onto a neutron star or a black hole, with a much deeper gravitational potential well, while it is usually negligible in CVs (van Paradijs & McClintock 1994; Shahbaz & Kuulkers 1998; King 1998). However, the WD in M31N 2008-12a is very compact with a mass of $1.37 M_{\odot}$, a radius $R_{\text{WD}} = 2000$ km and it is accreting at, or close to, the Eddington limit. We therefore checked the importance of disk irradiation using the approach given by Vrtilik et al. (1990) for different values of \dot{M} . At low disk mass accretion rates ($\dot{M} \lesssim 10^{-7} M_{\odot} \text{yr}^{-1}$; as typical for all other quiescent novae) irradiation increases the outer disk temperature by up to ~ 1000 K, which does not produce any significant change in the disk spectrum. At more moderate accretion rates ($\dot{M} \sim 10^{-6} M_{\odot} \text{yr}^{-1}$; i.e., M31N 2008-

12a at quiescence), irradiation increases the outer disk temperature by up to ~ 3000 K, thereby slightly affecting the spectrum by effectively decreasing the mass accretion rate, since irradiation increases the disk emission. At mass accretion rates above the Eddington limit ($\dot{M} \sim 10^{-5} M_{\odot} \text{yr}^{-1}$; M31N 2008-12a during the eruption), we find that irradiation increases the outer disk temperature by as much as 7000 K, and we would expect that this increase could reduce the mass accretion rate by a factor of ~ 2 . The effect of irradiation within M31N 2008-12a *at quiescence* is therefore only expected to slightly decrease the discrepancy in the disk mass accretion rate and the WD mass accretion rate.

Returning finally to the survival of the accretion disk. The presence of a disk, potentially with a high mass accretion rate, during the SSS phase of a nova eruption opens an intriguing possibility. Could a surviving accretion disk continue to feed significant fuel to the nuclear burning region on the WD, thereby ‘artificially’ extending the SSS phase, compared to a more typical nova (where the disk is assumed to be obliterated by the eruption)? Such a ‘refuelling’ would, for a short time, be akin to the persistent SSSs. Any mass accreted onto the WD during this period would be burnt to He and simply be added to the mass of the WD. Irrespective of the net gain or loss of accumulated mass during the nova eruption, ‘refuelling’ would enable net WD mass growth over the refuelling period. Here we only offer an outline of the concept, this is explored in more detail, observationally and theoretically, in Henze et al. (2017b).

5.2. The donor

There is an expectation that NIR observations of a quiescent nova system will largely isolate the donor star (see Darnley et al. 2012, and Figure 2); particularly for evolved (i.e. luminous) donors. The accretion disk models employed in this work only extend to 7500 \AA , but a simple extrapolation to longer wavelengths confirms that we can expect little, or no, accretion contribution to the quiescent flux in the NIR regime. Therefore we conclude that the PHAT NIR quiescent photometry *should* simply be photometry of the M31N 2008-12a mass donor. In Section 4.5, we used the accretion disk modeling at quiescence to estimate the *I*-band (F814W) contribution from the donor.

As can be seen in the left plot of Figure 8, the position of the quiescent M31N 2008-12a on a NIR color-magnitude diagram indicates that the donor is significantly less luminous and bluer than the red giants contained in the Galactic RG-novae (red points). However, the M31N 2008-12a donor may be consistent with the M31 red clump. A simple black body fit to the F110W ($\sim J$) and F160W ($\sim H$) photometry at quiescence ($t \sim 75$ d) yields $L_{\text{donor}} \sim 100 L_{\odot}$ and $T_{\text{eff, donor}} \sim 4800$ K.

If we include the extrapolated *I*-band luminosity of the donor (see Section 4.5), then the three-point donor SED is very well represented by the same black body fit. Hence we find that the M31N 2008-12a donor may be

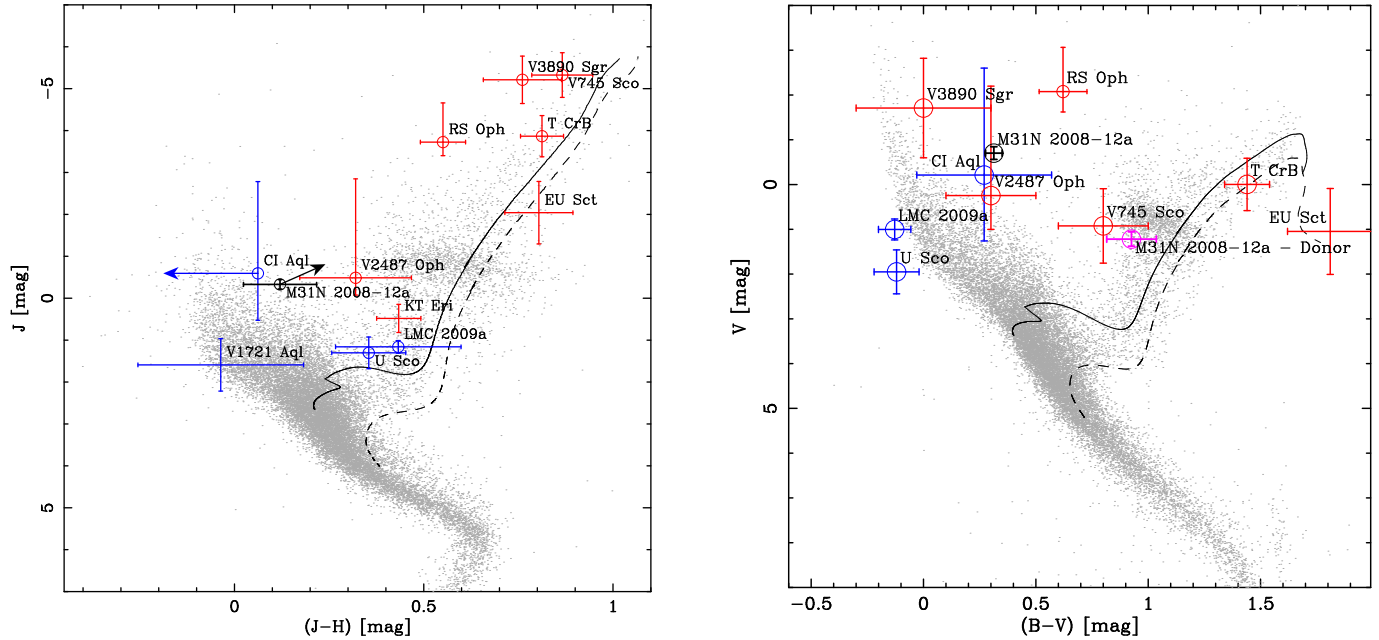


Figure 8. Color-magnitude diagrams showing stars from the Hipparcos data set (gray points; Perryman et al. 1997) with parallax errors $< 10\%$. These stars have been transformed to the distance and extinction of M31N 2008-12a. Photometry is taken directly from Hipparcos or the the 2MASS catalog (Skrutskie et al. 2006). The hashed-black and solid-black lines are evolutionary tracks of $1 M_{\odot}$ and $1.4 M_{\odot}$ stars, respectively (Pietrinferni et al. 2004). The red points represent Galactic RG-novae and the blue points Galactic or LMC SG-novae (see Schaefer 2010; Darnley et al. 2012; Bode et al. 2016, and references therein). The known RNe in this sample have been identified by an additional circle. **Left:** NIR color-magnitude diagram, the black data point shows the assumed location of the M31N 2008-12a quiescent system at minimum ($t \sim 75$ d), considering the uncertainty in the photometry and extinction, the black arrow indicates the increase in emission during quiescence ($t \sim 270$ d). **Right:** standard color-magnitude diagram showing the both position of the quiescent M31N 2008-12a system (black point; accretion disk+donor; $t \sim 75$ d) and the inferred position of the mass donor alone (magenta), which is consistent with the M31 red clump (see Section 5.2).

consistent with a black body of $T_{\text{eff,donor}} = 4890 \pm 110$ K, $L_{\text{donor}} = 103_{-11}^{+12} L_{\odot}$, and $R_{\text{donor}} = 14.14_{-0.47}^{+0.46} R_{\odot}$, at the distance of M31. This black body fit is illustrated by the dashed black line in the right-hand plot of Figure 2. We note that the quoted uncertainties are the formal 1σ errors resulting from the fitting process, the effect of possible systematic uncertainties related to the accretion and black-body disk modeling have not been estimated. Using this black body, we compute the extrapolated B and V photometry (again indicating the above caveats) of the donor and plot it’s position on a standard color-magnitude diagram in the right plot of Figure 8. Here, it is clear that the donor may indeed be consistent with the M31 red clump.

As is illustrated in the left plot of Figure 8, as the quiescent system evolves from its minimum state ($t \sim 75$ days) toward the next eruption ($t \sim 270$ days) the donor brightens by $J \sim 1$ mag, but becomes redder – consistent with an increase in the donor radius. Such behavior may be related to irradiation of the donor, causing heating and expansion of the atmosphere. It could also be related to the orbital phase of the system, with a tidally locked donor, for example, being non-spherically symmetric and unevenly heated – but such phase effects would imply a high system inclination. The very high luminosity of the accretion disk almost certainly now rules out high inclinations.

Given the available evidence, we must conclude that the mass donor in the M31N 2008-12a system is either a (low luminosity) red giant or post red giant branch star (e.g. horizontal branch), and/or that it is affected by significant irradiation from the primary, the disk, and the eruptions. Of course we must point out that there is a possibility that the star identified as the donor may simply be another star at a very similar position on the sky within M31. If we were relying on WFC/IR photometry of the donor alone, this probability would be quite large, but given the F814W spatial resolution of *HST*, the likelihood will be relatively small (around 2%, see DWB14 and Williams et al. 2014b).

One piece of evidence may be key to constraining the donor however. DHS15 proposed that the M31N 2008-12a ejecta interacts strongly, and immediately, with material in the circumbinary medium with a $1/r^2$ density dependence. This picture was strengthened by the re-analysis presented in DHB16. With such behavior being seen consistently across four consecutive eruptions (2012–2015), the circumbinary material must be continuously replenished. One feasible source of such material seems to be a stellar wind from a red giant donor as observed in RS Oph (see e.g., Bode et al. 2006). As ejecta-circumbinary shocks are not observed in CNe, we must infer that Roche-lobe overflow is too efficient a mass transfer process to build up significant material in the

circumbinary environment. Therefore the donor would be strongly constrained to any star capable of generating such a wind.

However, in this paper we have explored the possibility that the extremely luminous accretion disk generates a significant disk wind. Therefore, is it possible that such a disk wind is the source of the circumbinary pollution, not the stellar wind of a red giant donor? As such a giant donor could be transferring matter at a high rate to the disk via Roche lobe overflow. Indeed, could such a scenario be the only feasible manner in which such a high sustained WD mass accretion rate could be achieved?

5.3. *Orbital Period*

To date, the orbital period of the M31N 2008-12a system has eluded observation. However, given we now know the mass of the WD and have constrained the radius of the donor, we can place some restrictions on P_{orb} . We will assume that the donor has evolved at least enough to reside on the red giant branch and that it was originally the lower mass component of the binary. Therefore the donor mass must be somewhere in the range $0.8 - 8 M_{\odot}$ ⁶. If we assume that the donor is Roche lobe filling, then the orbital separation must be in the range $25 - 44 R_{\odot}$ ($0.12 - 0.20$ AU), hence $5 \lesssim P_{\text{orb}} \lesssim 23$ days. If the accretion is stellar wind driven then $P_{\text{orb}} \gg 5$ days. We note that such minimum orbital separations, and hence Roche lobe sizes, are far too large to account for natural accretion disk truncation by the presence of the donor.

Based on a suspected red giant donor, [DHB16](#) suggested that M31N 2008-12a might be the only known nova with $P_{\text{rec}} < P_{\text{orb}}$. Therefore, we again point out that as \dot{M} is a function of the donor–WD separation, that if $P_{\text{rec}} < P_{\text{orb}}$ any orbital eccentricity may affect the accretion rate and inter eruption timescale on an eruption by eruption basis.

5.4. *The X-ray flash non-detection*

The production of an X-ray flash at the onset of a nova eruption is a long standing prediction ([Starrfield et al. 1990](#); [Krautter 2002](#)). [Kato et al. \(2016\)](#) reported the results of an intensive *Swift* observing campaign to detect any X-ray flash associated with the 2015 eruption of M31N 2008-12a. This campaign did not detect such a flash and [Kato et al. \(2016\)](#) presented two explanations for the non-detection. Firstly, that the X-ray flash simply occurred before the *Swift* monitoring began – which requires the X-ray flash to precede the optical/NUV nova by $\gtrsim 8$ days. The second proposed that significant circumbinary material masked the flash signature. At the time, with little firm evidence for the nature of the donor, [Kato et al. \(2016\)](#) preferred the

first explanation. However, given the work reported in this paper, we emphasize that the X-ray flash could have been missed due to significant absorption from circumbinary material. This material could consist of some combination of a donor wind and a disk wind. For a low inclination system with a significant disk wind, the bulk of the circumbinary material could even reside along the line of sight. But the material in a disk wind dominated scenario may be expected to already be highly ionised, and hence unlikely to be able to mask any flash. Therefore, the X-ray flash could have been absorbed if there was significant pollution of the circumbinary environment by the wind of the donor, but likely only if the donor isn't Roche lobe filling.

5.5. *Time to reach the Chandrasekhar mass*

[TBW14](#) presented a prediction of the time required for the WD within the M31N 2008-12a system to grow to the Chandrasekhar mass (or at least to $1.37 M_{\odot}$). For example, they presented a $1.36 M_{\odot}$ WD, with $\dot{M}_{\text{acc}} = 1.7 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$, and a mass retention rate (the amount of accreted material remaining on the WD surface post-eruption) of 35%. This resulted in a time scale to grow to the Chandrasekhar mass of ~ 200 kyr.

We can update this simple calculation using the results from this paper and from [DHS15](#), but using the same approach as [TBW14](#). We will assume that the WD mass is actually $1.37 M_{\odot}$, as used for the disk modeling and that a further $0.01 M_{\odot}$ of accretion is required to reach the Chandrasekhar mass (the same required mass growth as [TBW14](#)). [HND15](#) determined that $(2.6 \pm 0.4) \times 10^{-8} M_{\odot}$ of H is ejected in each eruption – as a conservative estimate we will therefore assume that the total ejected mass is $6 \times 10^{-8} M_{\odot}$ ⁷ (consistent with [Kato et al. 2015](#))⁸. Retaining this conservative stance, we will assume that the *average* WD mass accretion rate over the entire 1 yr cycle is in fact the absolute *minimum* rate predicted by this paper, $\dot{M}_{\text{acc}} = 6.4 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$; assuming a similar amount of mass is lost in the form of a disk wind. Even then at such a lower accretion limit, the retained mass, or accretion efficiency is a staggeringly high 90%. Much higher than the $\sim 30\%$ predicted by [TBW14](#) and the 63% calculated by [Kato et al. \(2015\)](#). Combining these new data, we arrive at an updated prediction of the time to reach the Chandrasekhar mass of < 20 kyr — *possibly* much shorter.

6. SUMMARY & CONCLUSIONS

In this paper we have presented our analysis of an unrivalled series of *Hubble Space Telescope* photometric observations of the final stages of the 2015 eruption of M31N 2008-12a. In this analysis, we have also exploited

⁷ Assuming roughly equal mass of H and He in the ejecta. [Kato et al. \(2014\)](#) assume $X=0.55$, $Y=0.43$, and $Z=0.02$ in the ejecta.

⁸ We further note that this ejected mass assumes a spherical geometry. With highly asymmetrical ejecta and the proposed low inclination, it is possible that the ejected mass is much greater.

⁶ The lower limit to allow evolution onto the red giant branch by the present day, the upper is the approximate zero-age upper mass limit to form a WD.

archival *HST* imaging during quiescence, and Keck spectroscopy 2014 eruption of M31N 2008-12a. Our main findings include:

1. The *HST* WFC3/UVIS photometry of the late decline of the 2015 eruption shows a steady decline toward quiescence from the $\sim I$ -band to the NUV, broadly consistent with the general trends established from the early-decline ground-based and *Swift* photometry.
2. When combined with the archival *HST* photometry – shown to have been taken between eruptions – the system appears to reach a flux minimum (in all bands) ~ 75 d post-eruption, before again increasing in luminosity by ~ 270 d post-eruption; ~ 100 d before the next eruption.
3. The broadband SEDs of the late decline and during quiescence were explored using accretion disk models. The results indicate that these SEDs, even as early as 13.4 d post-eruption, are consistent with the emission being dominated by an accretion disk – one that has survived the eruption.
4. The inferred accretion rates are initially above the Eddington accretion limit ($t = 13.4$ d), indicating a disk that has survived albeit in a severely shocked and heated state. The disk luminosity and inferred \dot{M} then decline toward minimum ($t \sim 75$ d) before increasing again ($t \sim 270$ d) presumably as the disk fully reestablishes.
5. The computed accretion rates, even at quiescence are large, with the disk luminosities still close to the Eddington limit. We speculate that mass loss from the disk will lead to a disk wind.
6. Could such a disk wind contribute a significant quantity of material to the circumbinary environment, and could provide the matter source with which the ejecta are observed to interact, and possibly even the ejecta collimation mechanism?
7. Our disk modelling computed a range of $\dot{M} = (1.2\text{--}2.8) \times 10^{-6} M_{\odot} \text{ yr}^{-1}$ during quiescence. Even when accounting for disk winds, which might account for half of \dot{M} , the derived accretion rates onto the WD at quiescence are still in the range $\dot{M}_{\text{acc}} = (0.6\text{--}1.4) \times 10^{-6} M_{\odot} \text{ yr}^{-1}$, significantly larger than any other nova. If confirmed, WD accretion rates this high cause significant problems for a number of well-established nova eruption models.
8. [DHB16](#) and [DHG17](#) both proposed the presence of highly collimated outflows or even jets during the eruption. Could such a high \dot{M} drive a bipolar outflow from the inner disk and boundary layer, even at quiescence?
9. Archival *HST* WFC3/IR photometry on the system isolates the donor. Coupled with a strong $\sim I$ -band excess from the accretion disk modeling, this photometry indicates a donor with $T_{\text{eff,donor}} = 4890 \pm 110$ K, $L_{\text{donor}} = 103_{-11}^{+12} L_{\odot}$, and $R_{\text{donor}} = 14.14_{-0.47}^{+0.46} R_{\odot}$ – consistent with the M31 red clump.
10. The NIR colors of the donor are slightly redward of the red clump and there is significant variation in the donor luminosity at quiescence. These may be signs that the donor is significantly irradiated by the WD, disk, and ejecta, or may also be due to orbital phase affects.
11. Based on the work presented in this paper, the updated time-scale for the system to reach the Chandrasekhar mass has fallen to < 20 kyr.

These *HST* observations of the late-decline of the 2015 eruption, combined with serendipitous archival detections during quiescence have started to shed some light on the inter-eruption behaviour of M31N 2008-12a. It is clear that UV observations of this remarkable system are key to fully entangling the extreme physics at play throughout the entire eruption cycle. Vital questions that should be addressed over the coming eruptions include the balance between accreted matter and ejected matter, in the light of the apparent large variation in quiescent \dot{M} – to fully assess the ultimate fate of M31N 2008-12a.

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